New insights on galaxy evolution since z ~ 1.2 from the CFHT Legacy Survey

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Jean Coupon (東北大学)

Collaborators

- Paris (IAP): H. J. McCracken, Y. Mellier
- Marseille (LAM): O. Ilbert
- CFHT: S. Arnouts
- Garching (Excellence cluster Universe): M. Kilbinger
- VVDS & VIPERS teams: U. Abbas, S. de la Torre, Y. Goranova, P. Hudelot, J.-P. Kneib, and O. Lefèvre

Overview

- cosmologists have built a “standard model” relating the dynamics of the Universe to its constituents, but leaving us with unresolved questions regarding their origin
- astronomers face a troubling paradox: the only constituent (also the rarest) we best know about it, baryonic matter, is observed today in a remarkable diversity of galaxies we struggle to explain with simple arguments
- the field has known spectacular progresses from deep and large scale surveys, combined with semi-analytic simulations and new analysis techniques
- the halo occupation distribution (HOD) model is based on the well-known dark matter halo model and the simple assumption that the number of galaxies only depends on halo mass to probe the relationship between galaxies and their host haloes
- we applied this new method for the first time to the CFHTLS-Wide survey which remains unmatched in terms of volume at high redshift and image quality

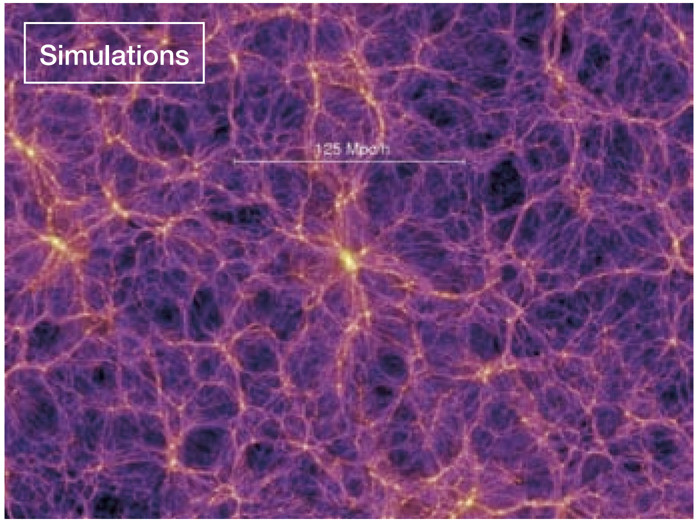
Plan

- I. Introduction: cosmological context and open questions
- II. Linking the galaxy distribution to Dark Matter: the HOD model
- III. Measurements in the CFHTLS Wide
- IV. Results: new insights on galaxy evolution since z ~ 1.2
- V. Conclusions

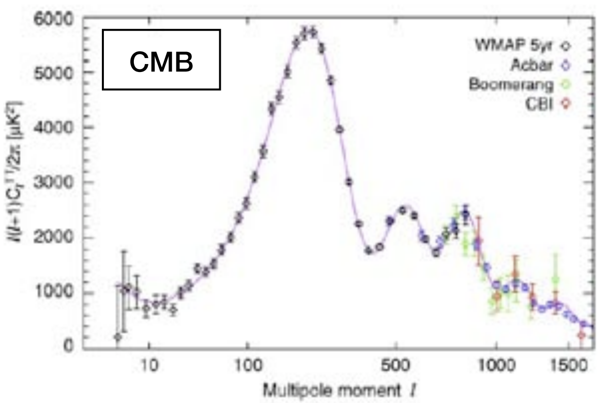
I. Introduction: cosmological context and open questions

The Universe in one slide

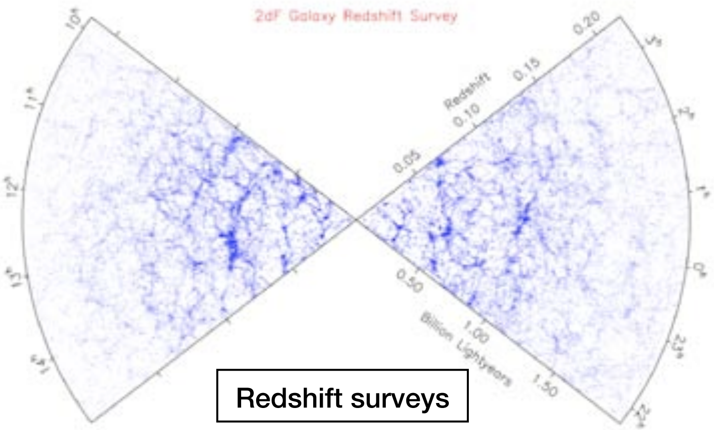
- the Universe is homogeneous and isotropic
- flat
- expanding and accelerating
- is dominated by (cold) dark matter and dark energy
- structures grow by gravitational instabilities of primordial tiny fluctuations that experienced a huge expansion during an early inflation period

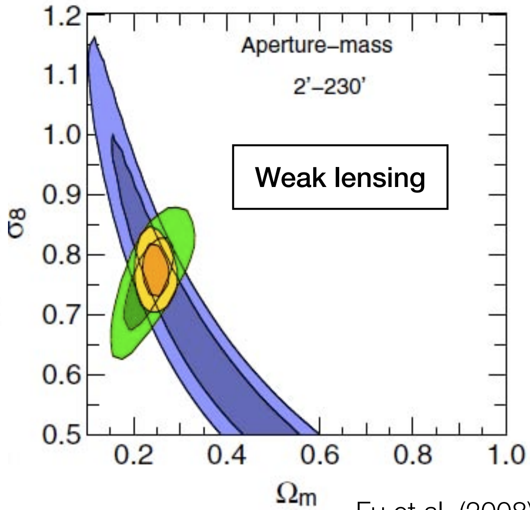


Millenium simulation



Nolta et al. (2009)





Fu et al. (2008)

Measuring the Universe

- General relativity (Einstein, 1915) is the law of gravitation. It relates the space-time curvature to its content:

$$G_{ik} = R_{ik} - \frac{1}{2}g_{ik}\mathcal{R} - \Lambda g_{ik} = \frac{8\pi G}{c^4}T_{ik}$$

- Einstein's equation in the Friedmann-Robertson-Walker metric (consequence of the cosmological principle):

$$\mathrm{d}s^2 = c^2\mathrm{d}t^2 - a^2(t)\left(\frac{\mathrm{d}r^2}{1-kr^2} + r^2\mathrm{d}\theta^2 + r^2\sin^2\theta\mathrm{d}\phi^2\right)$$

- which leads to Friedmann's equations:

$$\frac{\ddot{a}}{a} = -\frac{4\pi G}{3}(\rho + \frac{3P}{c^2}) - \frac{\Lambda c^2}{3}$$

$$\left(\frac{\dot{a}}{a}\right)^2 = \frac{8\pi G}{3}\rho - \frac{kc^2}{a^2} + \frac{\Lambda c^2}{3}$$

matter/energy is a fluid described by its equation of state: $p=w\rho c^2$

- where we define:

Hubble parameter

$$H(t) = \frac{\dot{a}}{a}$$

redshift

$$1 + z = \frac{a(t_o)}{a(t_e)}$$

$$z = \frac{\lambda_o - \lambda_e}{\lambda_e}$$

Expansion rate and the cosmological parameters

Equation of state: $p \qquad w\rho c$

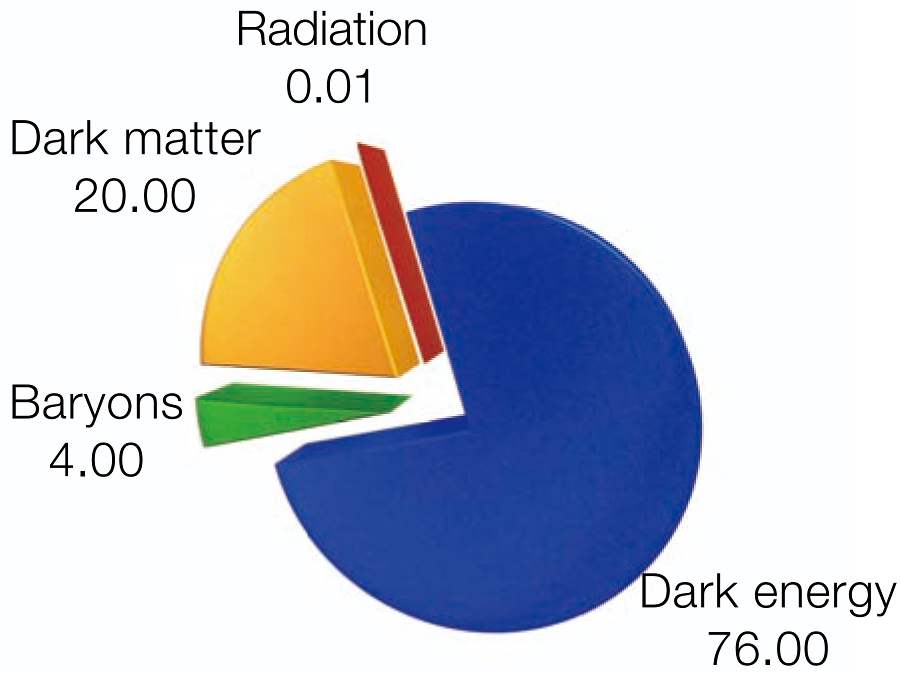
- Matter: $w_m = 0$
- Radiation: $w_r = 1/3$
- Vacuum energy: $w_\Lambda = -1$

From dimensionless Friedmann's equation:

$$(\ )^2 = \ {}^{2}_{0}\left(\Omega \quad {}^{-3} + \Omega \quad {}^{-4} + \Omega_\Lambda + \Omega \quad {}^{-2}\right)$$

with $\Omega \ := \dfrac{\rho}{\rho}$ and the constraint today:

$$\Omega \quad + \Omega \quad + \Omega_\Lambda + \Omega \quad = 1$$



- What's the nature of dark matter?
- What's the nature of dark energy?
- time evolution of w_Λ?

≡ ΛCDM model: we assume it's our representation of the Universe

The growth rate of structures

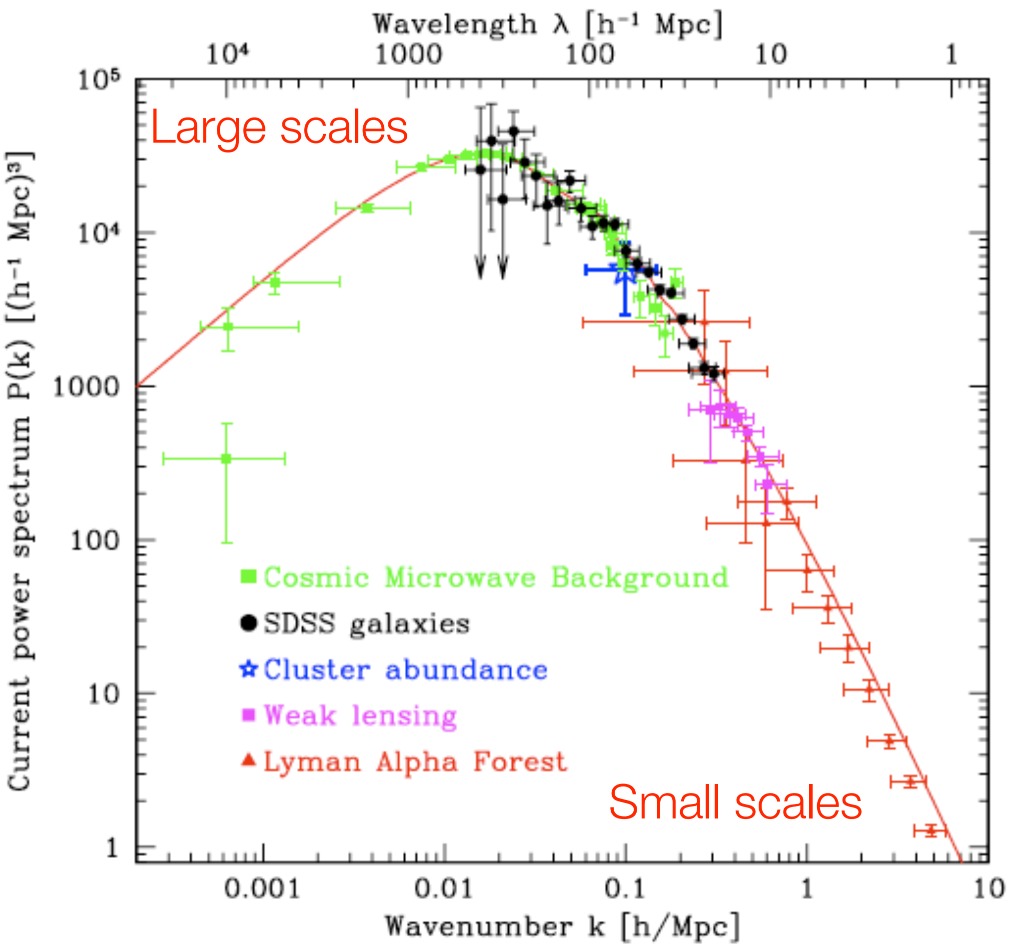
Matter fluctuations are characterised by P(k,t), which depends on cosmological parameters:
H_0, Ω_m, Ω_Λ, w_Λ + $\boldsymbol{\sigma}_8$, n_s

Initial power spectrum

$$P(k,t) = D_+^2(t)P(k,0) = D_+^2(t)Ak^{n_s}$$

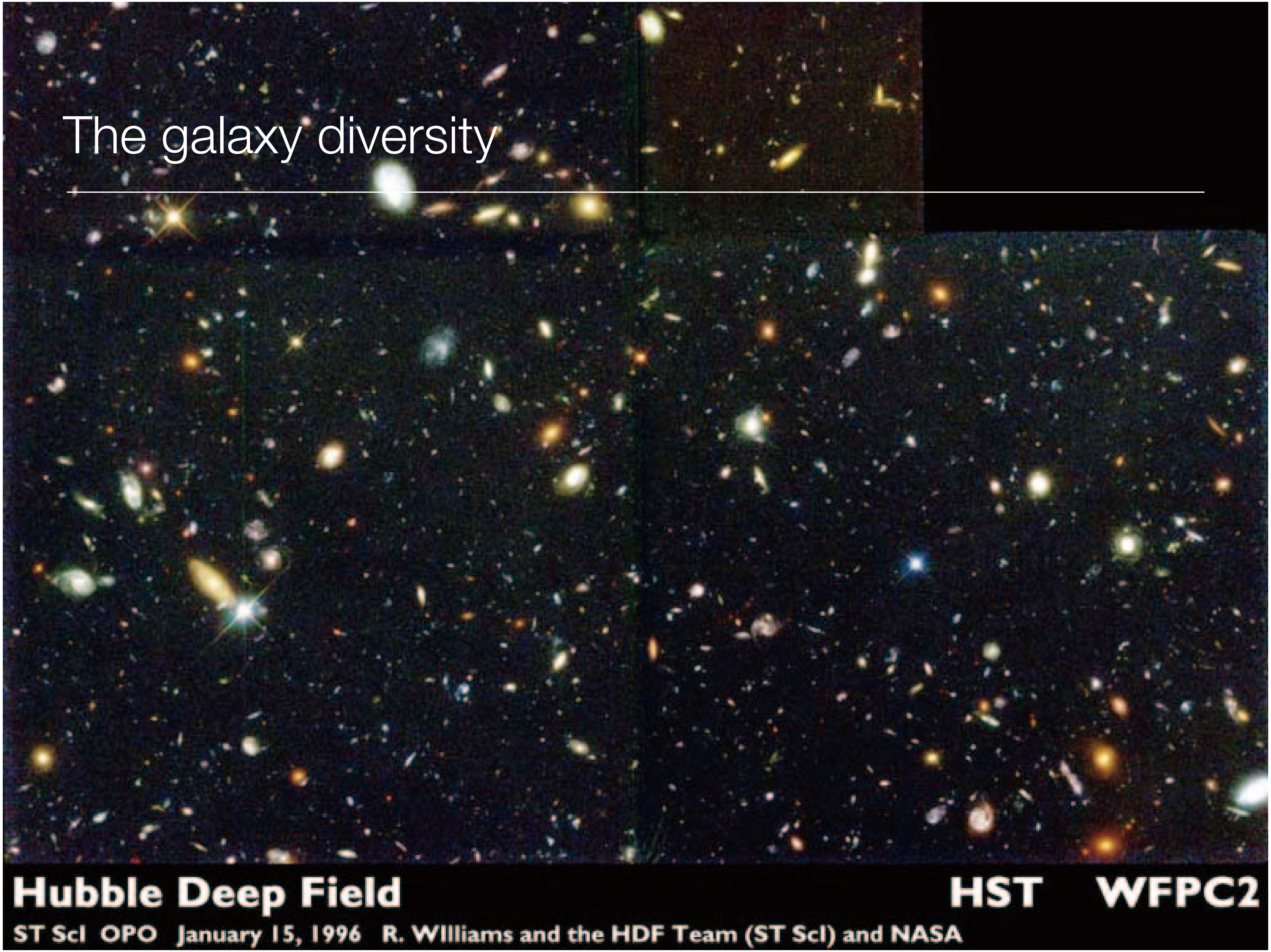
Growth of structures (strongly depends on DE and DM):

$$D_+(a) = \frac{5}{2}\Omega_m \frac{H(a)}{H_0}\int_0^a \left(a\frac{H(a)}{H_0}\right)^{-3} \mathrm{d}a$$

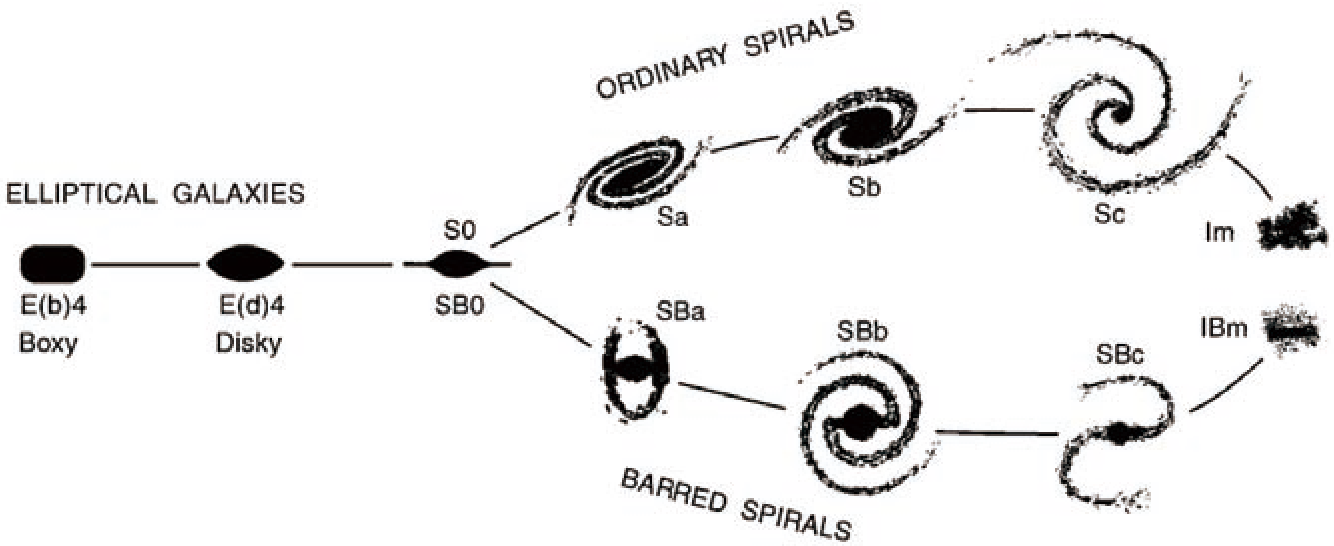


Tegmark et al. (2004)

The galaxy diversity



The Hubble sequence



Kormendy & Bender (1996)

dwarf, irregular,
\+ peculiar galaxies and
active galactic nuclei

Elliptical galaxies
or early-type galaxies
or “red” galaxies

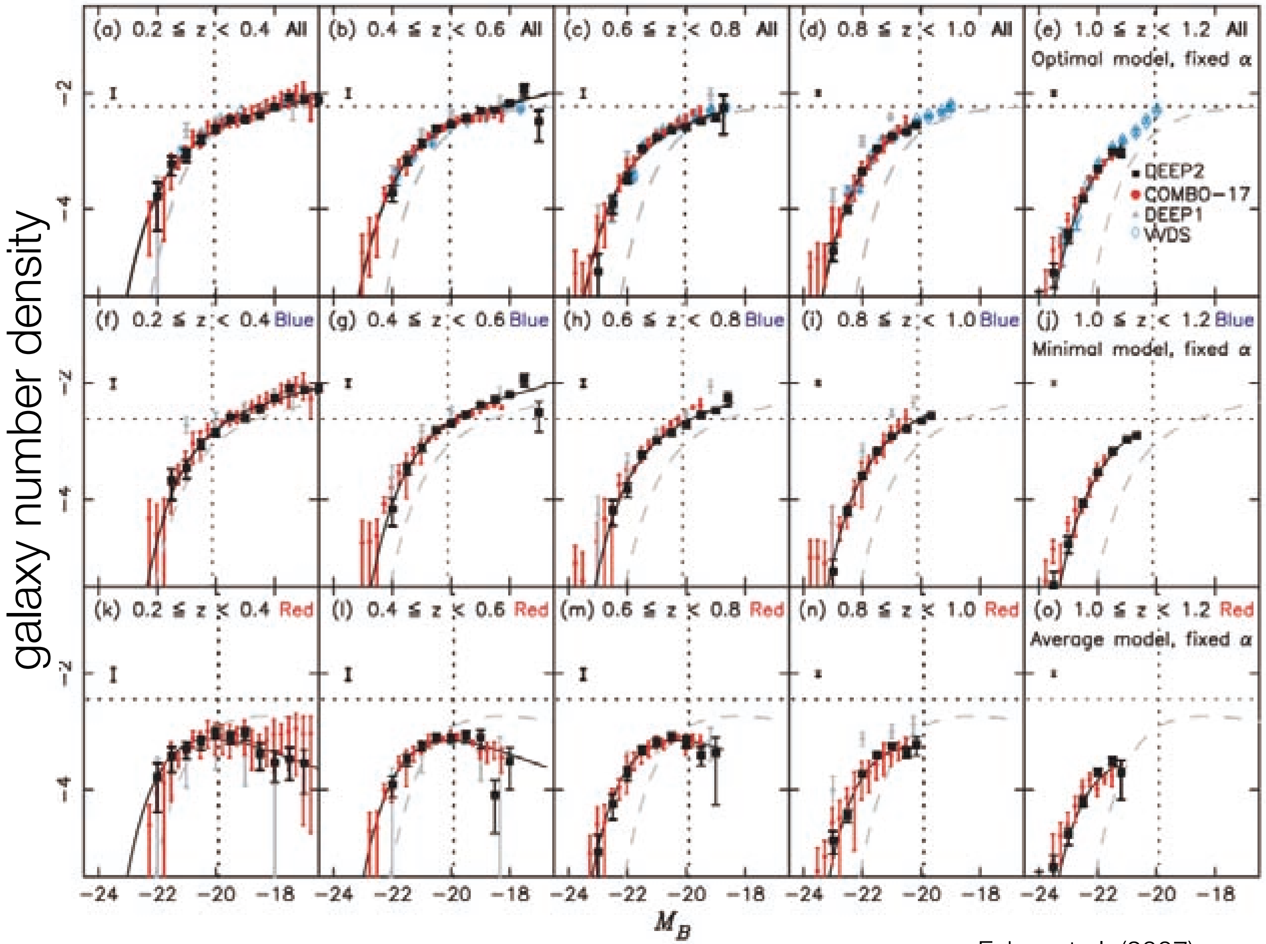
Spiral (disk) galaxies
or late-type galaxies
or “blue” galaxies

rare objects but carry
some precious information
about galaxy evolution

How did galaxies form and evolve from the initial baryon density field to the galaxy diversity as seen today?

Galaxy luminosity function

redshift →



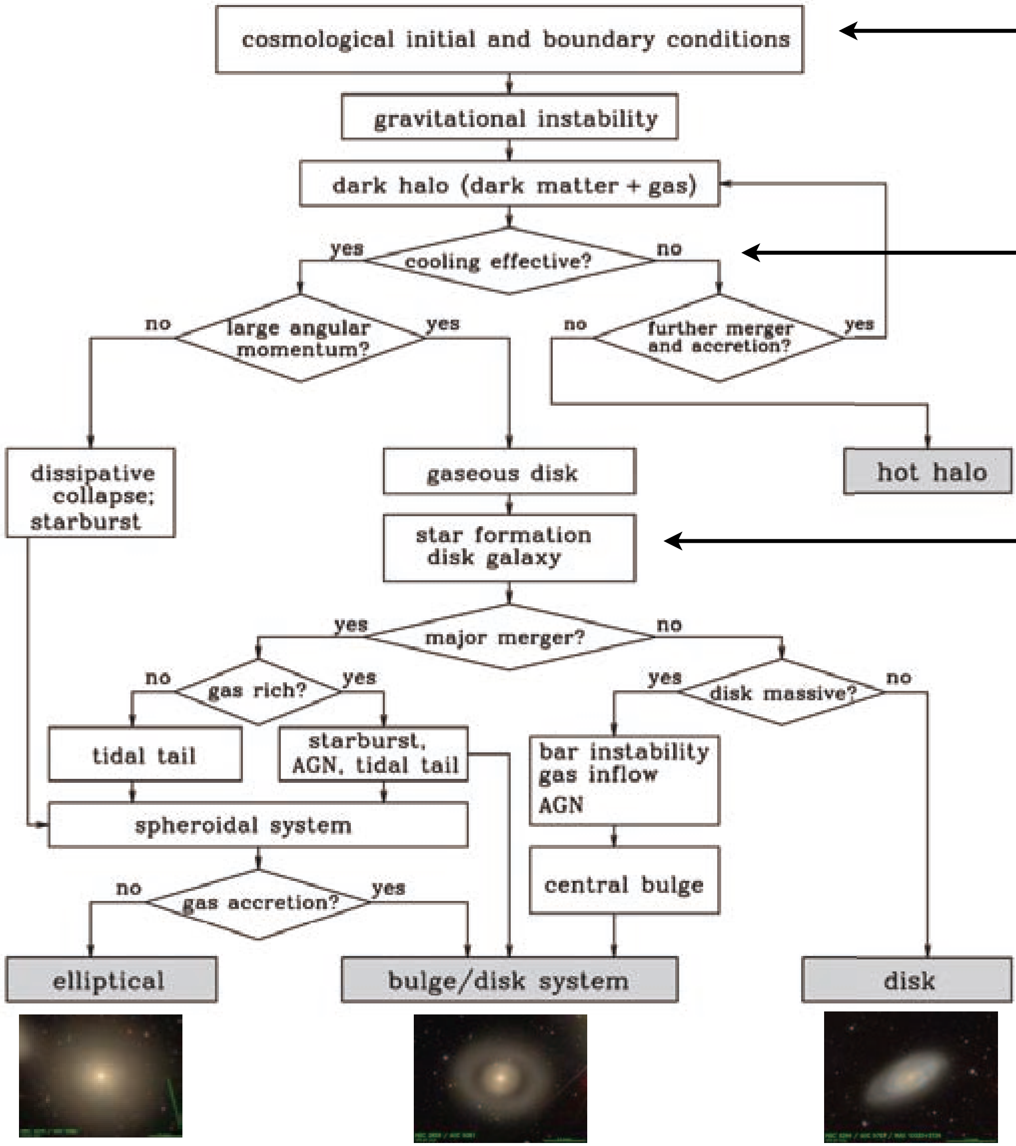
Faber et al. (2007)

- all galaxies. In grey: results from the local Universe (SDSS) -> redshift evolution
- blue galaxies. Dominate at faint luminosity
- red galaxies. Dominate at bright luminosity

WHY?

Basics of galaxy formation and evolution

Mo et al. (2011)



ΛCDM cosmology

No cooling, no star

Star formation stops

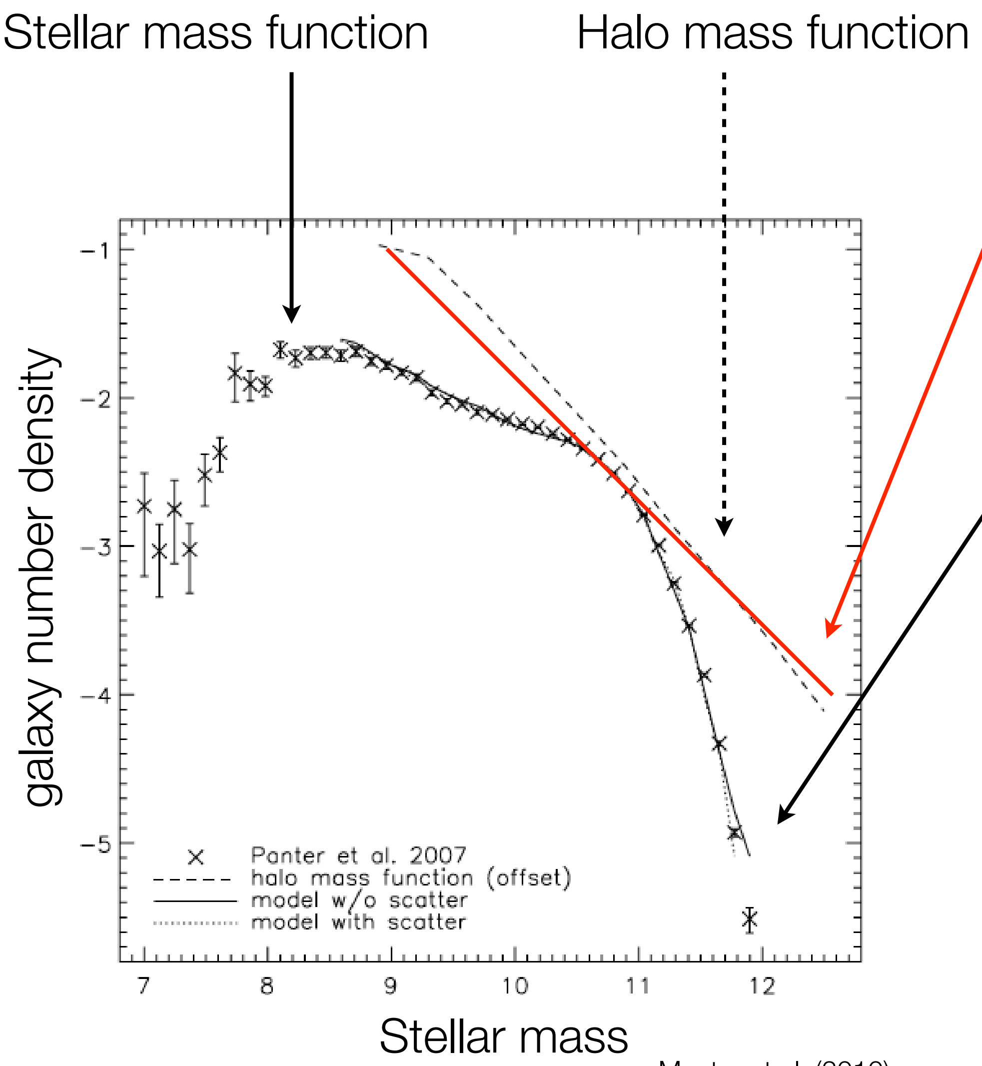
Star formation goes on

Galaxy evolution: depends on halo mass, environment and redshift

Hubble sequence observed today

SDSS, Blanton & Hogg

Star formation efficiency in dark matter haloes



Moster et al. (2010)

- stellar mass function in a naive picture where all baryons are transformed into star
- but star formation efficiency depends on halo mass

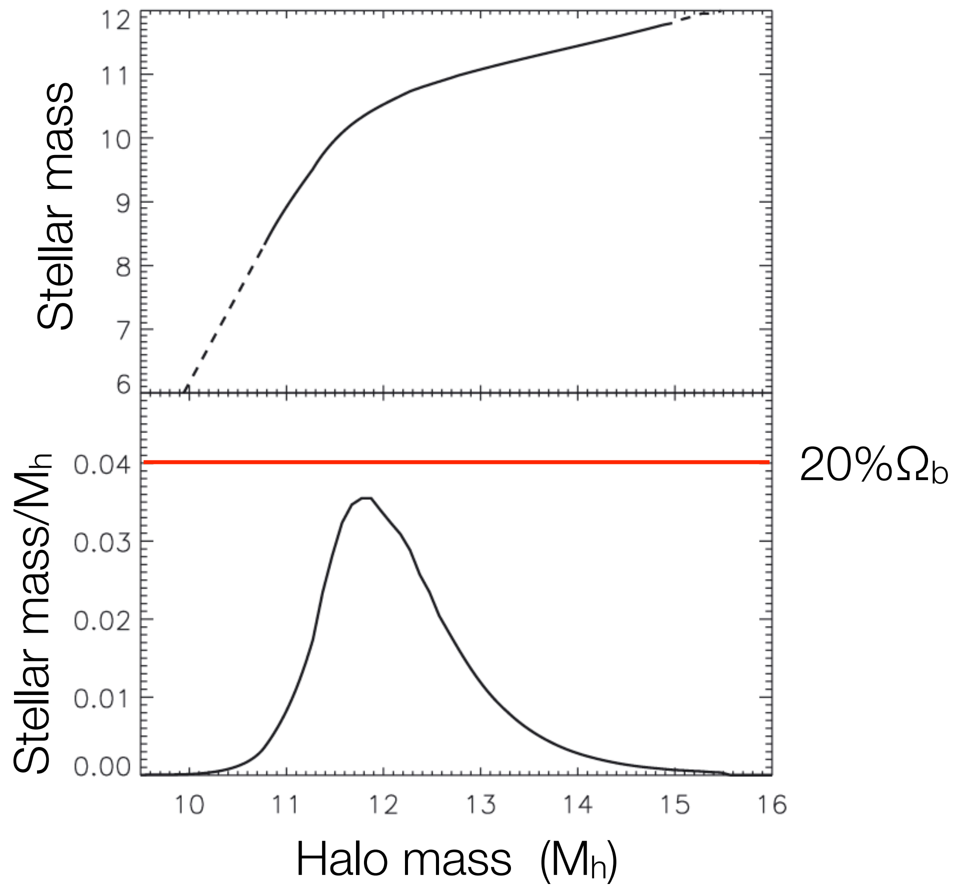
"no cooling, no star": there must be physical processes capable of

1. preventing the gas from cooling
2. or reheating the gas

Star formation efficiency in dark matter haloes

- stellar mass = only few % of the available baryon “fuel” (Ω_b)
- reaches a maximum of ~ 20% (at z = 0) for galaxies ~ Milky Way
- star formation less efficient in low and high mass haloes

“abundance matching” Guo et al. (2010)

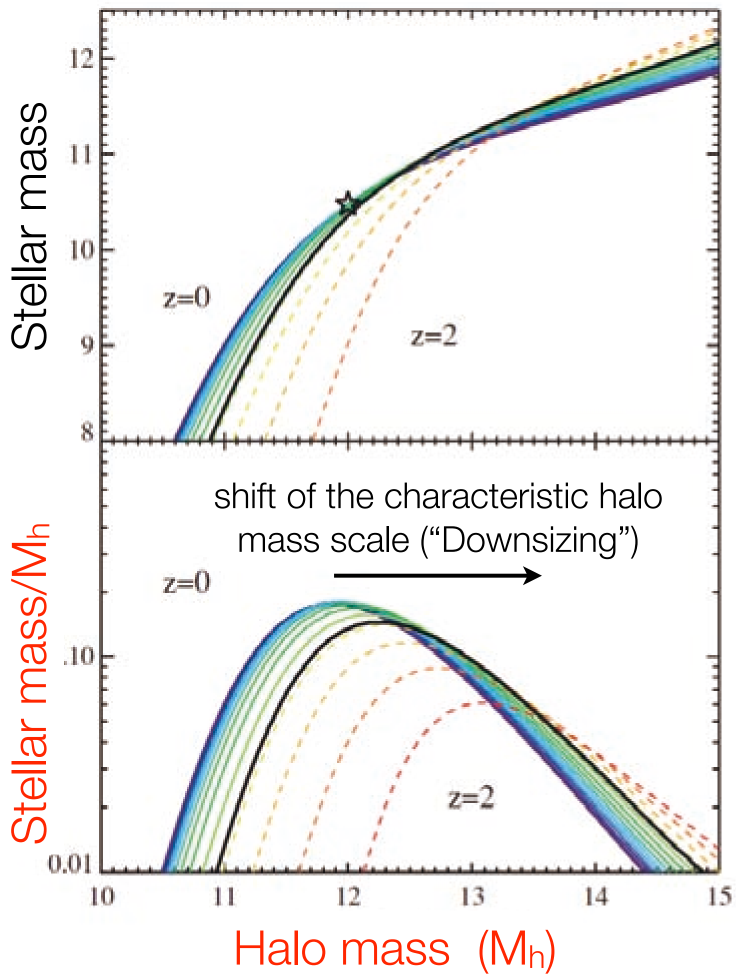


depends on:
- the available gas
- feedback processes “quenching” star formation

low-mass haloes:
supernovae, stellar winds

high-mass haloes:
AGNs

Open questions



Conroy & Wechsler (2009)

"Pure" hydrodynamical simulations face difficulties to reproduce the observed stellar abundance.

Key questions:

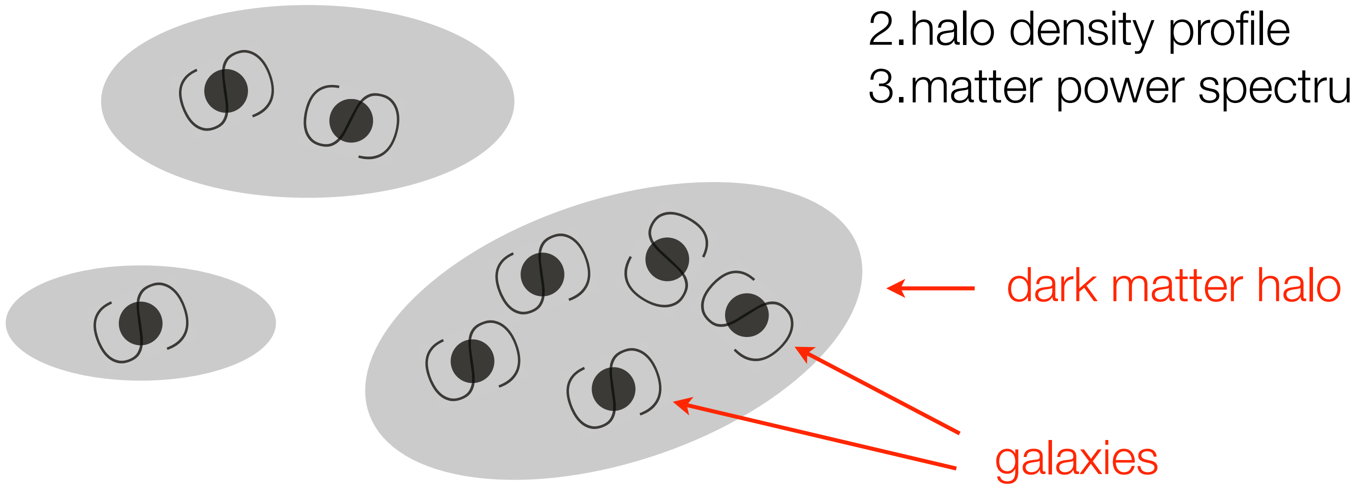
- the role of galaxy merging?
- AGN and supernovae feedback processes?
- the importance of the environment?
- impact of a different cosmology?

☆Strategy☆:

The relationship between M_{star} and M_h changes with time and at a rate which depends on halo mass. Observing the stellar-to-halo-mass ratio as function of redshift helps to understand the physical processes involved in galaxy evolution.

II. Linking the galaxy distribution to Dark Matter: the HOD model

The halo occupation distribution model



Dark matter distribution defined by:
1. halo mass function
2. halo density profile
3. matter power spectrum (halo bias)

Main hypothesis: the number of galaxies only depends on halo mass: $\propto \quad {}^{\alpha}$

+ dark matter (halo) space distribution, one can predict:
- the galaxy number density
- the galaxy distribution (clustering)

The halo model

Cooray & Sheth (2002):

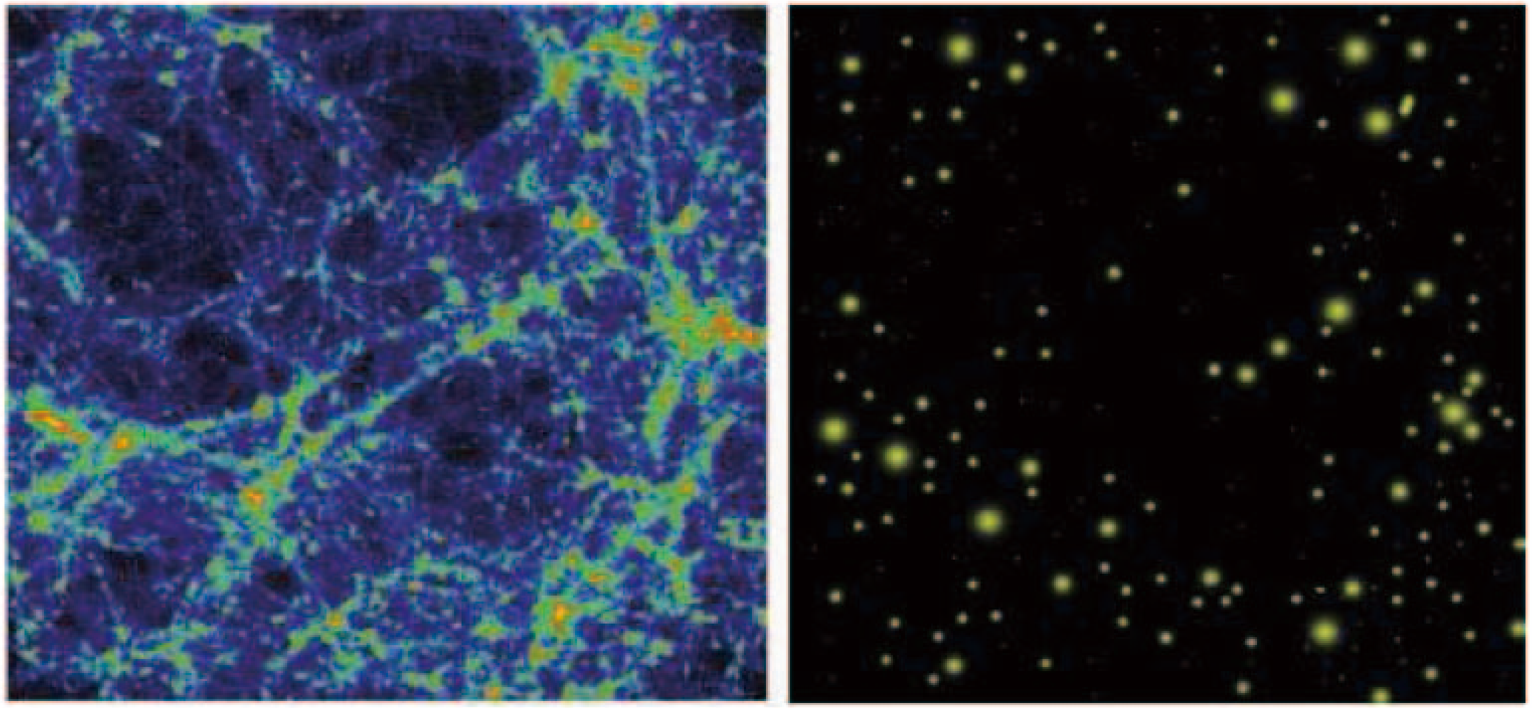
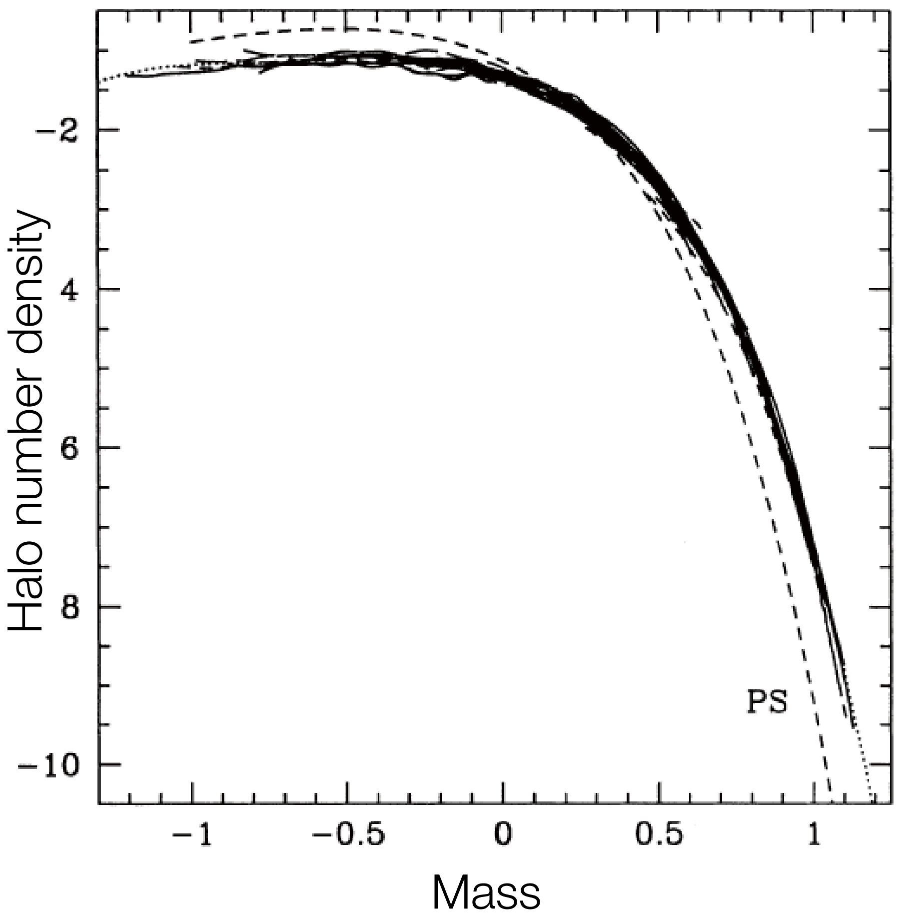


Fig. 1. The complex distribution of dark matter (a) found in numerical simulations can be easily replaced with a distribution of dark matter halos (b) with the mass function following that found in simulations and with a profile for dark matter within halos.

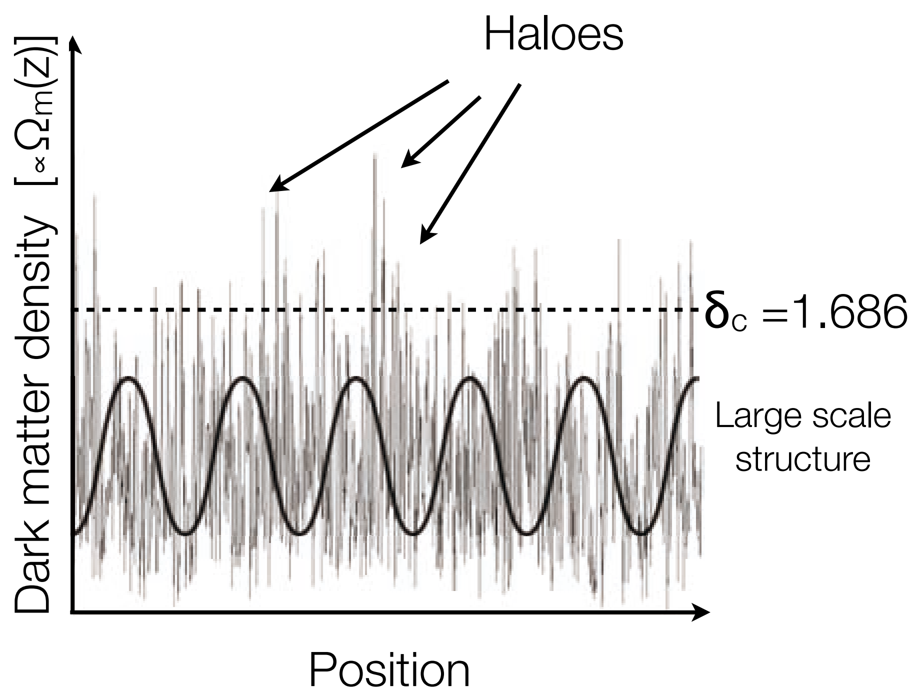
1. The halo mass function

Is the number density of haloes as function of halo mass.

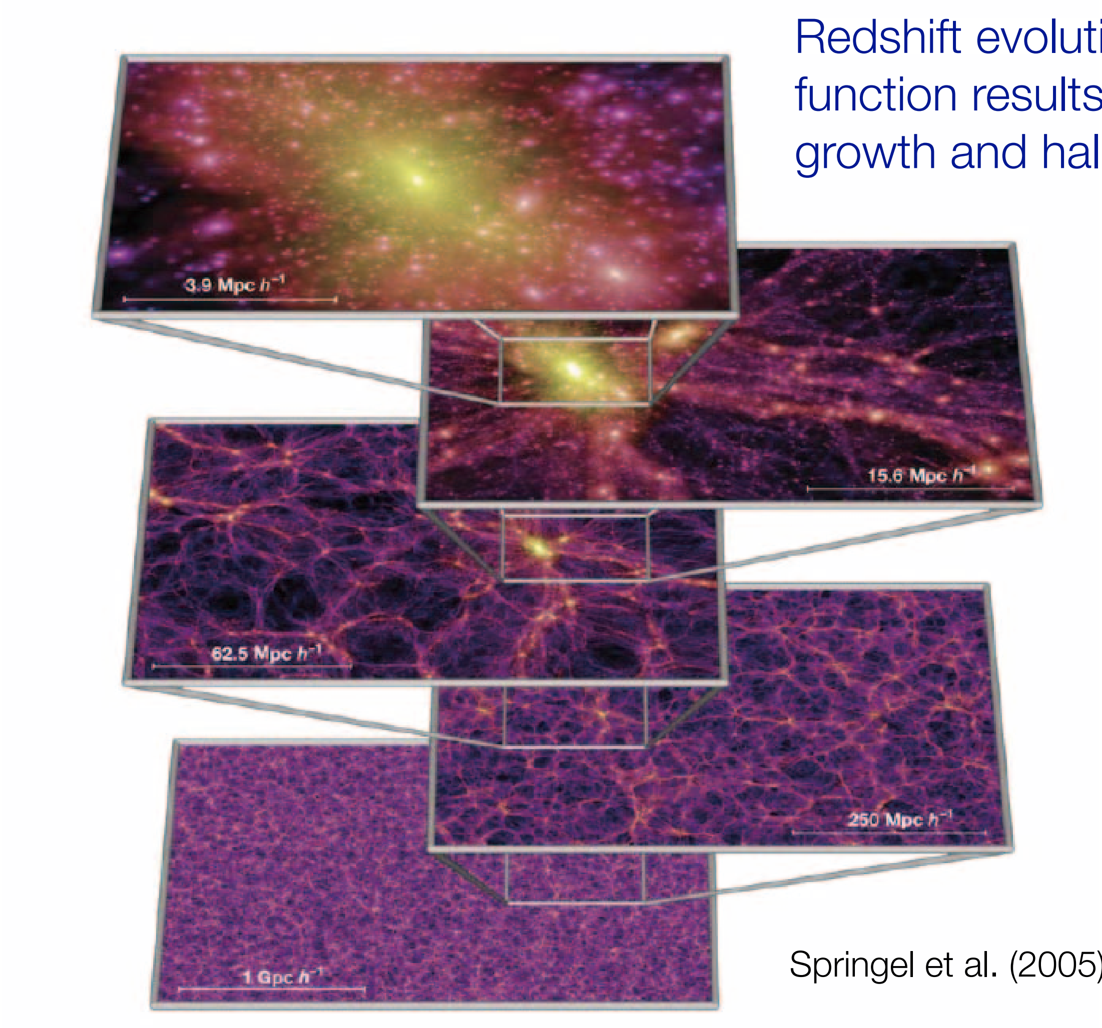


Jenkins et al. (2001):

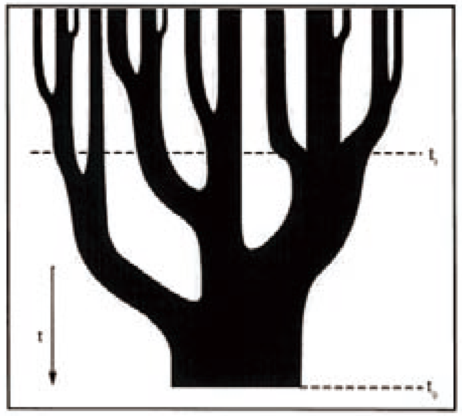
Press & Schechter (1974) formalism:
in a CDM Universe, dark matter collapses in overdense regions above the critical density δ_c to form virialised haloes.



1. The halo mass function (redshift evolution)

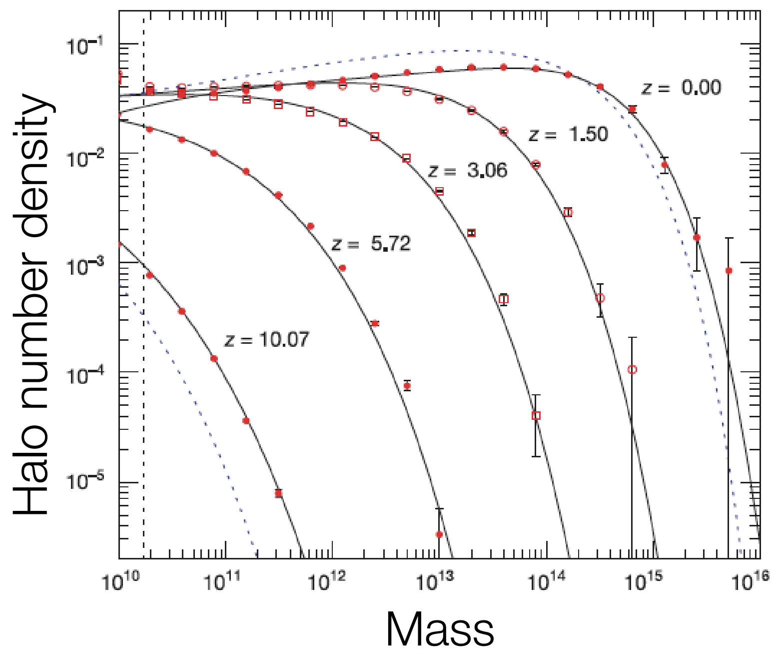


Redshift evolution of the mass function results from structure growth and halo merging.

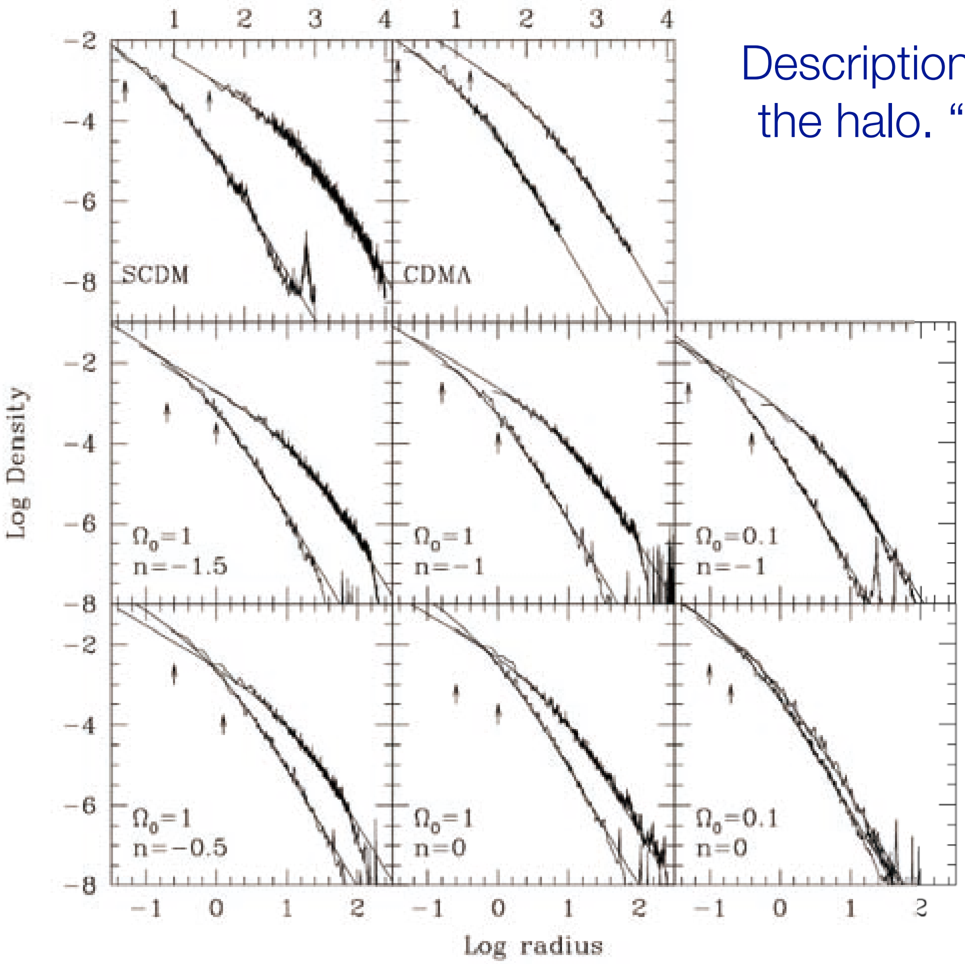


Lacey & Cole (1993)

Springel et al. (2005)



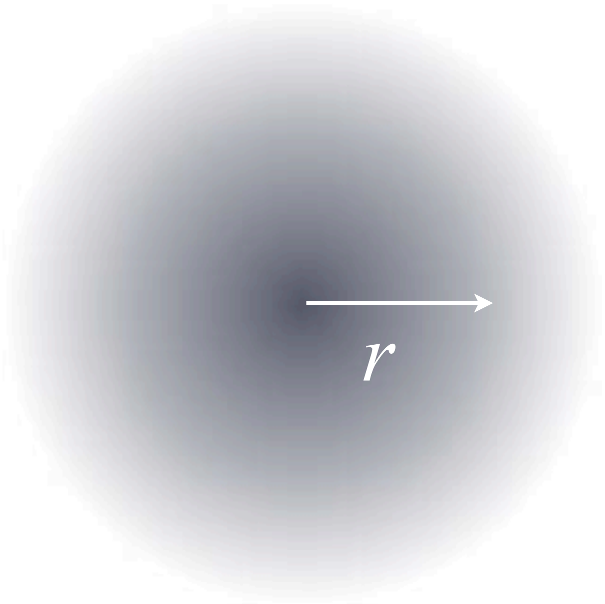
2. Halo density profile



Navarro, Frenk & White (1997)

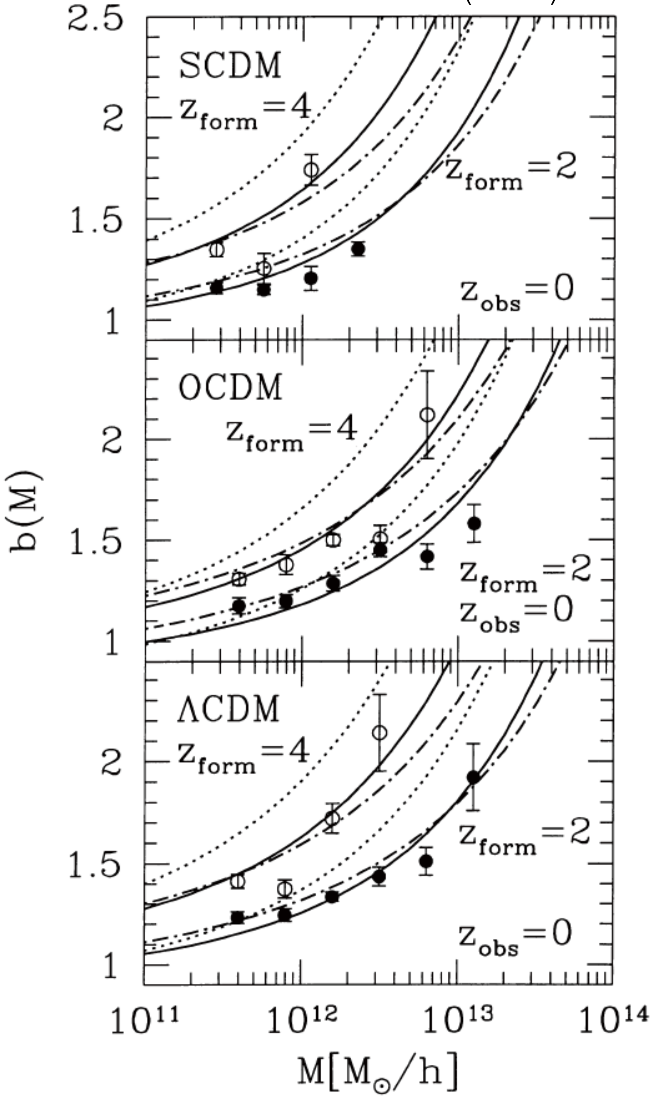
Description of how dark matter concentrates within the halo. “NFW” profile accurately fits simulations:

$$\rho_h(r|M) = \frac{\rho_s}{(r/r_s)(1+r/r_s)^2}$$

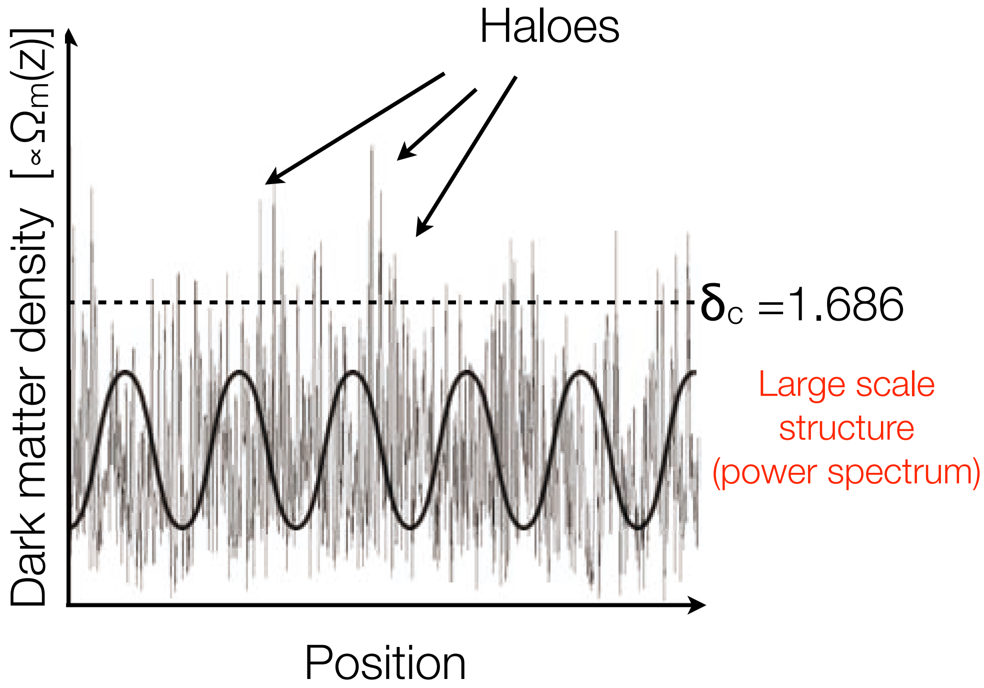


3. Halo bias

Sheth & Tormen (1999)



Describes dark matter halo clustering.

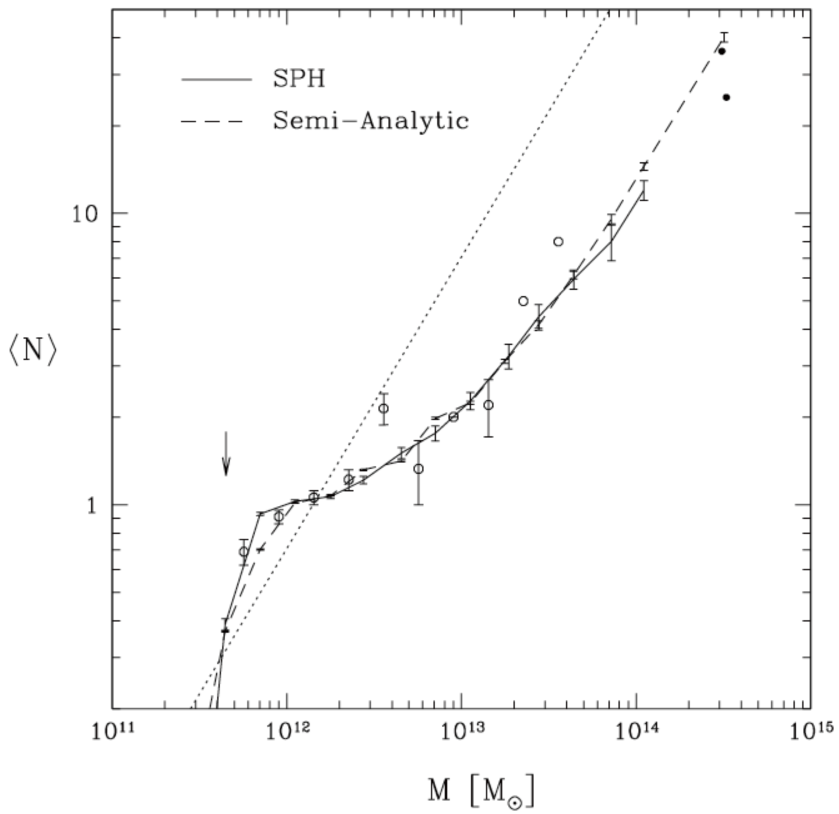


Haloes are "biased" tracers of the matter density fluctuations. b(M) is deduced from "Press & Schechter" formalism and simulations.

HOD parameterisation

Idea: the number of galaxies ONLY depends on halo mass
(Berlind & Weinberg 2002, 2003; Kravtsov et al. 2004; Zheng et al. 2005).

Berlind & Weinberg (2003)



Semi-analytic simulations:

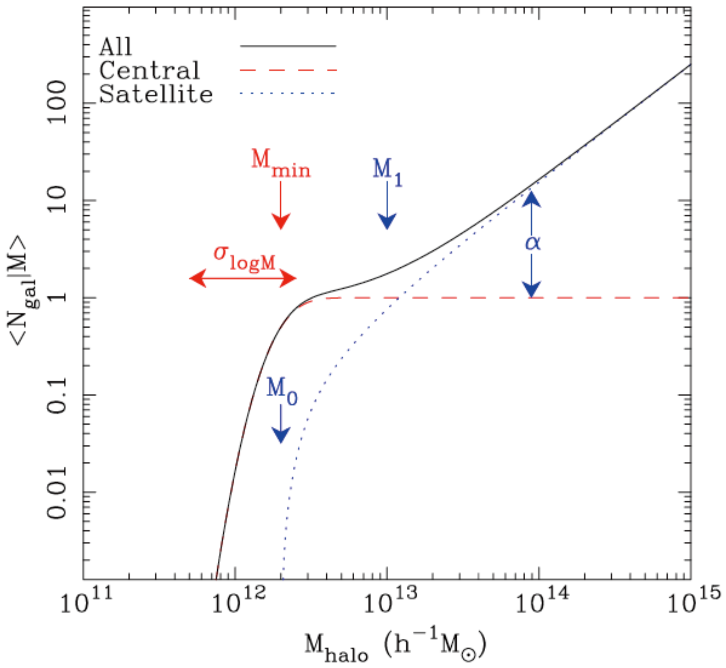
- the number of galaxies only depends on halo mass
- there is a mass below which no galaxy forms
- the number of galaxies then follows a power law as function of halo mass

$$\propto \quad {}^{\alpha}$$

Observable (galaxy properties)

Halo mass

HOD parameterisation



Wake et al. (2011)

Separation of the contribution from central and satellite galaxies (Zheng et al. 2005, 2007):

$$N_c(M) = \frac{1}{2}\left[1 + \mathrm{erf}\left(\frac{\log M - \log M_{min}}{\sigma_{\log M}}\right)\right],$$

$$N_s(M) = N_c(M) \times \left(\frac{M - M_0}{M_1}\right)^{\alpha}.$$

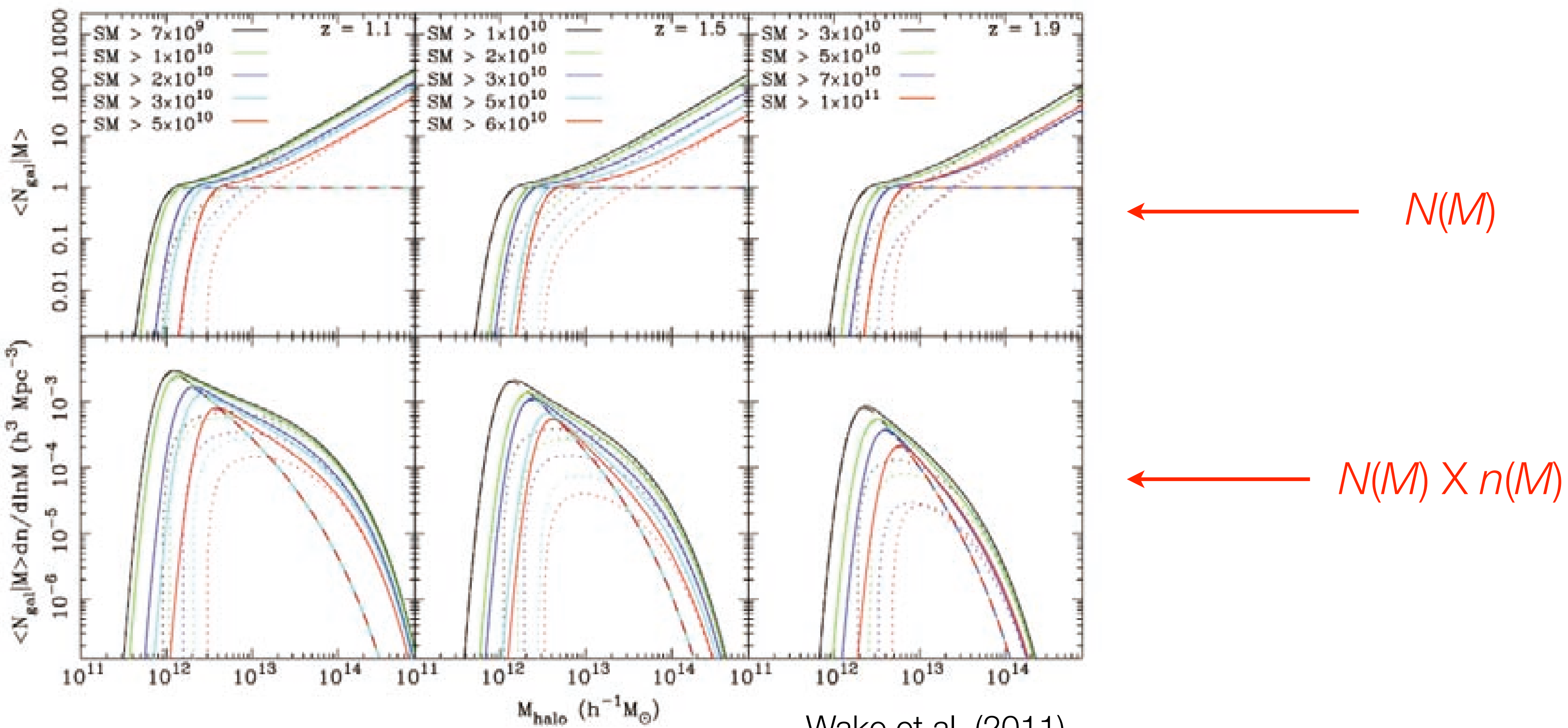
Important assumption: allows to treat central and galaxy contribution separately.

- the smooth transition takes into account the scatter in galaxy formation
- M_{min} corresponds to the halo mass of central galaxies
- M_1, **α** describe the satellite number

From HOD to observables: galaxy number density

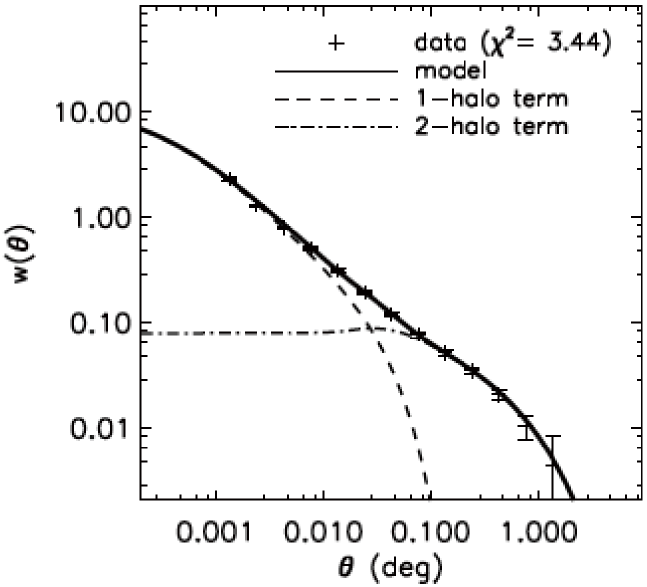
Convolution between halo mass function and HOD:

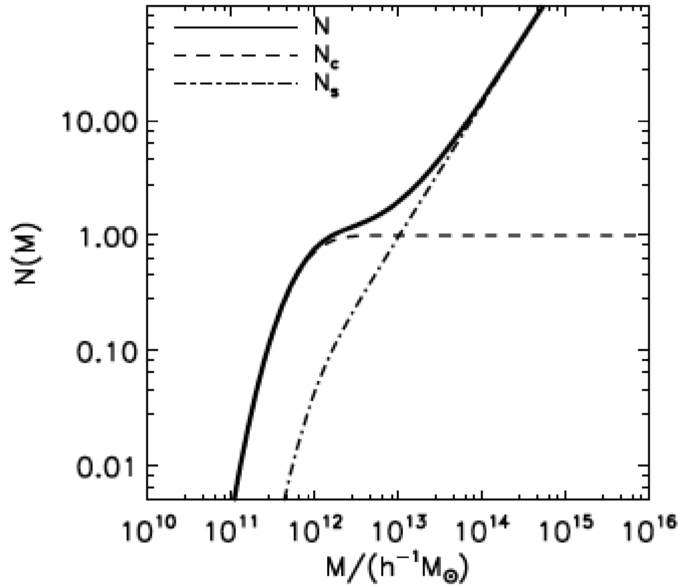
$$n_{\rm gal}(z) = \int N(M)\, n(M,z)\, \mathrm{d}M$$



Wake et al. (2011)

From HOD to observables: galaxy clustering





Halo density profile convolved with HOD

Halo bias convolved with HOD

Last step: projection of xi(r) with redshift distribution

$$\xi(r) = 1 + \xi_1(r) + \xi_2(r)$$

$$w(\theta) = 2\int_0^\infty dx\, f(x)^2 \int_0^\infty du\, \xi(r = \sqrt{u^2 + x^2\theta^2})\,,$$

Physical parameters

$$b_g(z) = \int \mathrm{d}M\, b_h(M,z)\, n(M,z) \frac{N(M)}{n_{gal}(z)}$$

Galaxy bias

$$\langle M_{halo} \rangle(z) = \int \mathrm{d}M\, M\, n(M,z) \frac{N(M)}{n_{gal}(z)}$$

Mean halo mass

$$f_c(z) = \int \mathrm{d}M\, n(M,z) \frac{N_c(M)}{n_{gal}(z)}$$

Satellite fraction (1-f_C)

III. Measuring galaxy clustering in the CFHTLS Wide

Photometric redshifts

LePhare: PHotometric Analysis for Redshift Estimations (S. Arnouts & O. Ilbert). Code for photometric redshift estimation based on template fitting and bayesian approach.

magnitudes
+ errors

$\chi^2(z, T, A)$

SED libraries:
galaxies/stars/QSO
+ filter set
+ extinction, IGM

(many assumptions here)

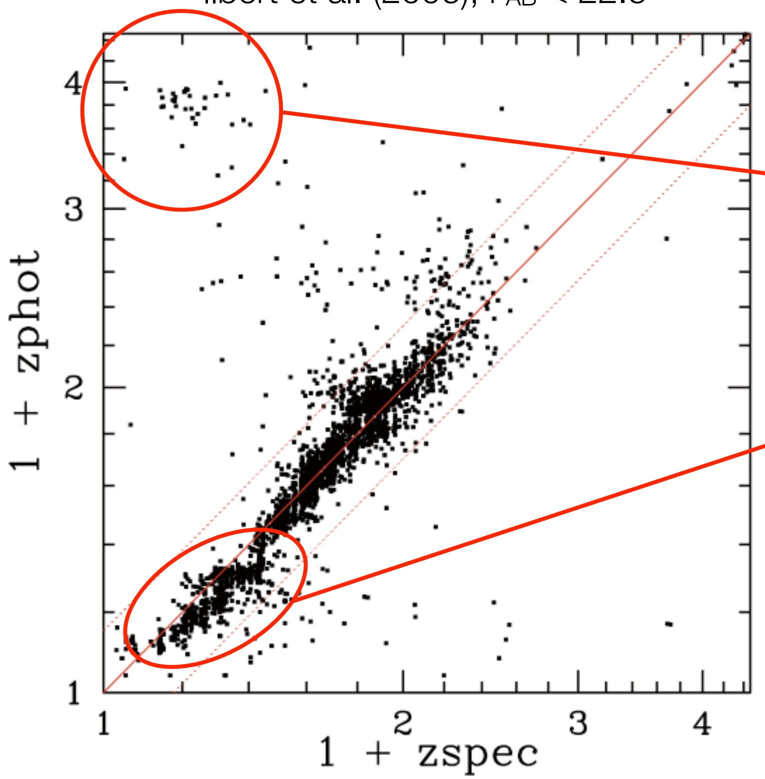
- Best redshift, second solution
- Type, E(B-V), extinction
- Returns a χ^2 for galaxy, star and QSO library
- Uncertainties, PDF(z)
- Absolute magnitudes
- Physical parameters: mass,mean age,SFR,etc.

advanced features

Results with LePhare

Raw results with template fitting
(no spectroscopic calibration)

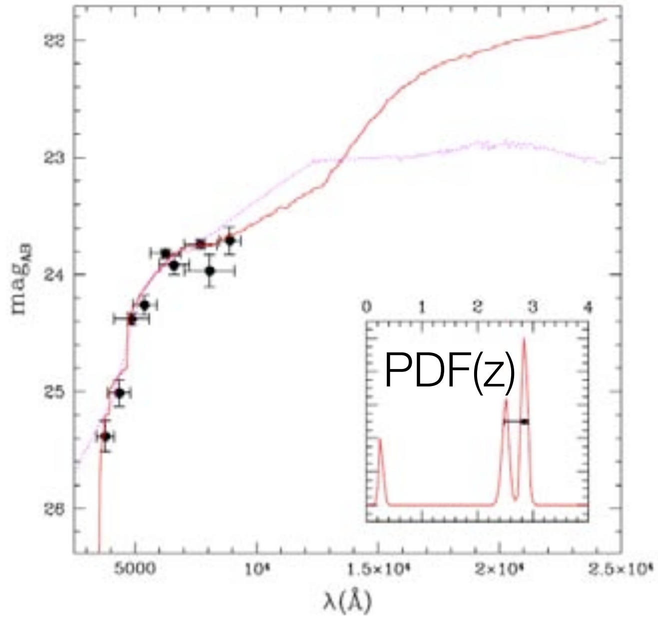
Ilbert et al. (2006), i’$_{AB}$ < 22.5



CFHTLS deep vs VVDS deep (le Fevre et al. 2005)

Catastrophic errors
= “Outliers”

Systematic trends



Idea: calibrate photo-z’s with spectro-z’s.
Ilbert et al. (2006):

1. correction of systematic offsets
2. template optimization
3. galaxy distribution priors

Improved results with LePhare

CFHTLS photo-z's vs VVDS spectro-z's, i'$_{AB}$ < 24

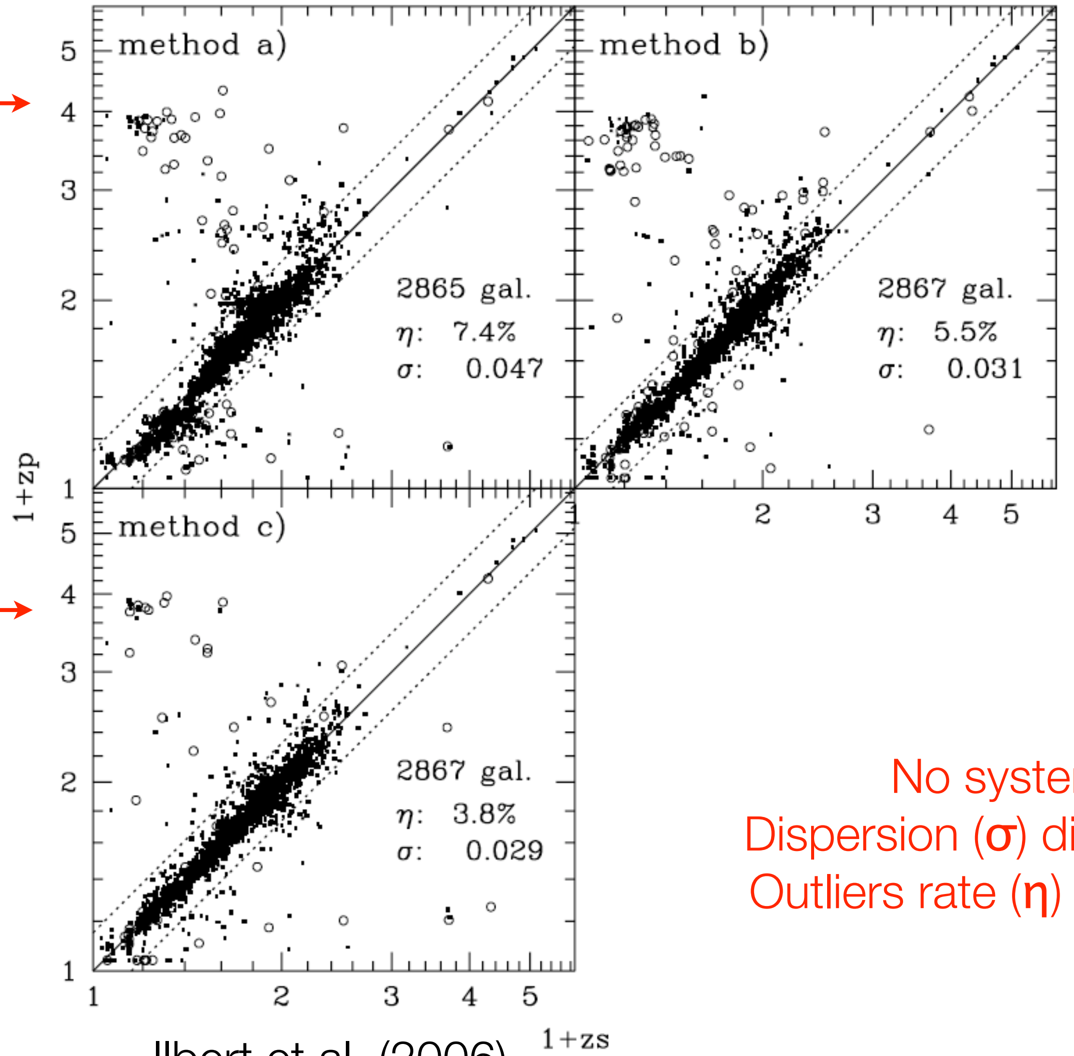
Raw template fitting - no improvement

Offset correction + template optimization

Removes systematics

priors

Reduces catastrophic errors



No systematic
Dispersion (σ) divided by 1.5
Outliers rate (η) divided by 2

Ilbert et al. (2006)

Photo-z errors

Comparison with spectro-z's
= "real errors"

photo-z dispersion, **σ**, robust estimator of the standard deviation:

$$\sigma_{\Delta z/(1+z_s)} = 1.48 \times \mathrm{median}(|\Delta z|/(1+z_s))$$

outlier rate, **η** (number of catastrophic errors), as the number of objects with:

$$|\Delta z| \geq 0.15 \times (1+z_s)$$

Computed by Le Phare
= "estimated errors"

From the PDF(z), 68% confidence limits are defined as:

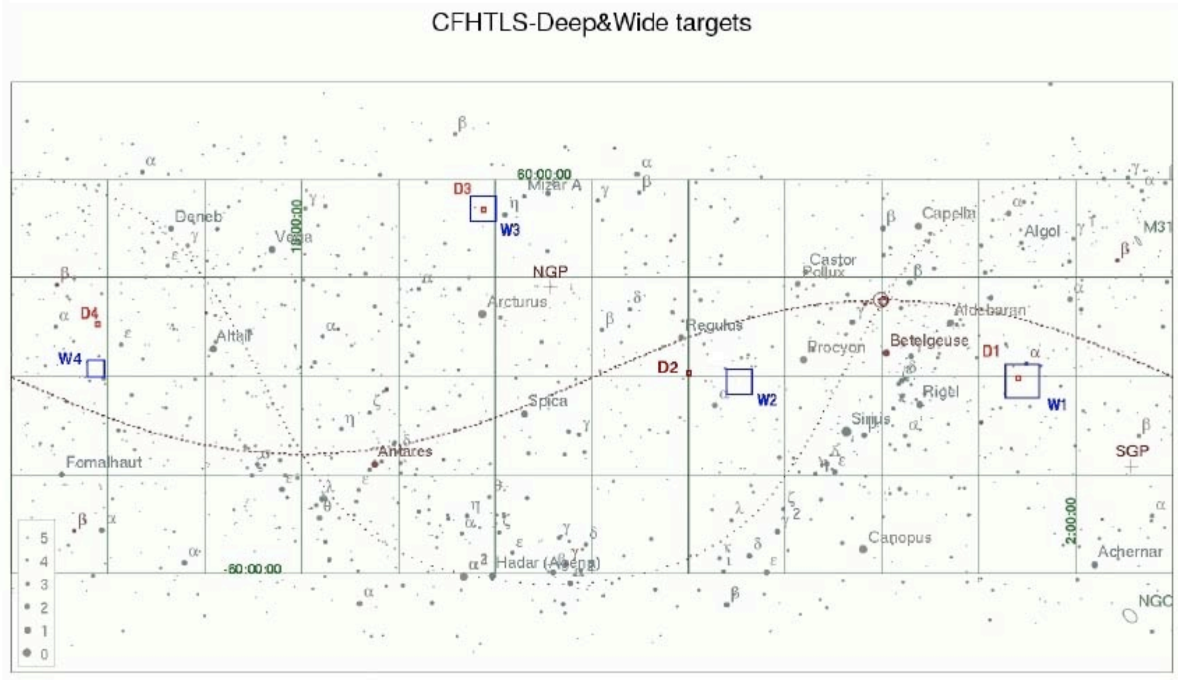
$$\Delta\chi^2(z) = \chi^2(z) - \chi^2_{\mathrm{min}} = 1$$

and our error estimate becomes:

$$\sigma_{z_p} = \frac{|z_{\mathrm{left}}(68\%) - z_{\mathrm{right}}(68\%)|}{2}$$

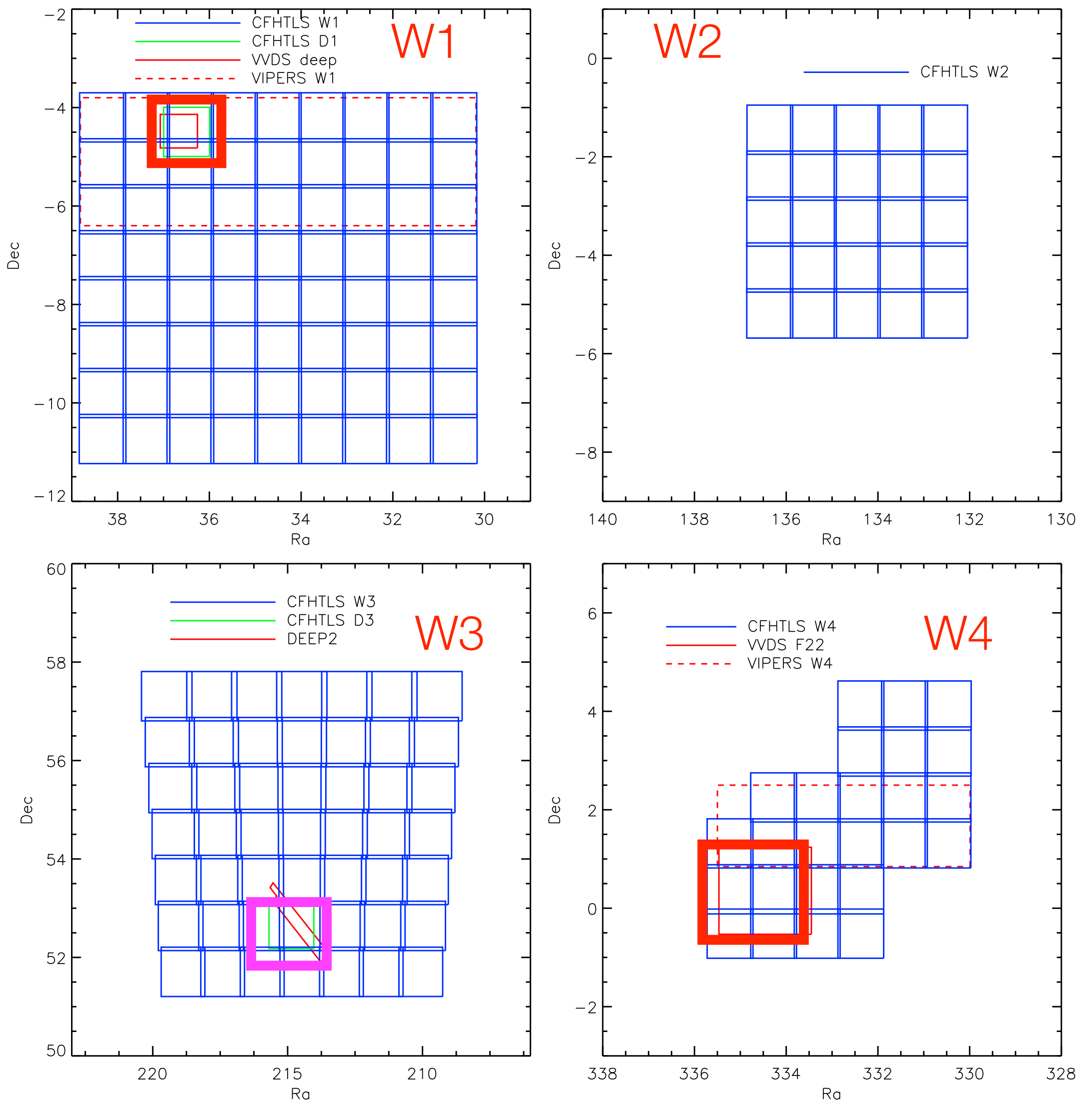
$$\sigma_{\Delta z/(1+z_s)} = \sigma_{z_p}$$

The Canada-France-Hawaii Telescope Legacy Survey (CFHTLS)



- 450 nights observed with Megacam @ CFHT in u,g,r,i,z filters
- Terapix is in charge for the data reduction (latest release T0006)
- the survey is completed since december 2008
- Deep survey: 4 independent fields (total 4 deg2) i < 27.5
- Wide survey: 4 independent fields (total 133 deg2) i < 25.5
- Very Wide survey

CFHTLS-Wide and spectroscopic data



Effective area of the Wide: 133 deg^2
3,000,000 photo-z's in the Wide part
600,000 photo-z's in the deep part
(S/N ~ 40)

calibrated with 20,000 spectra:

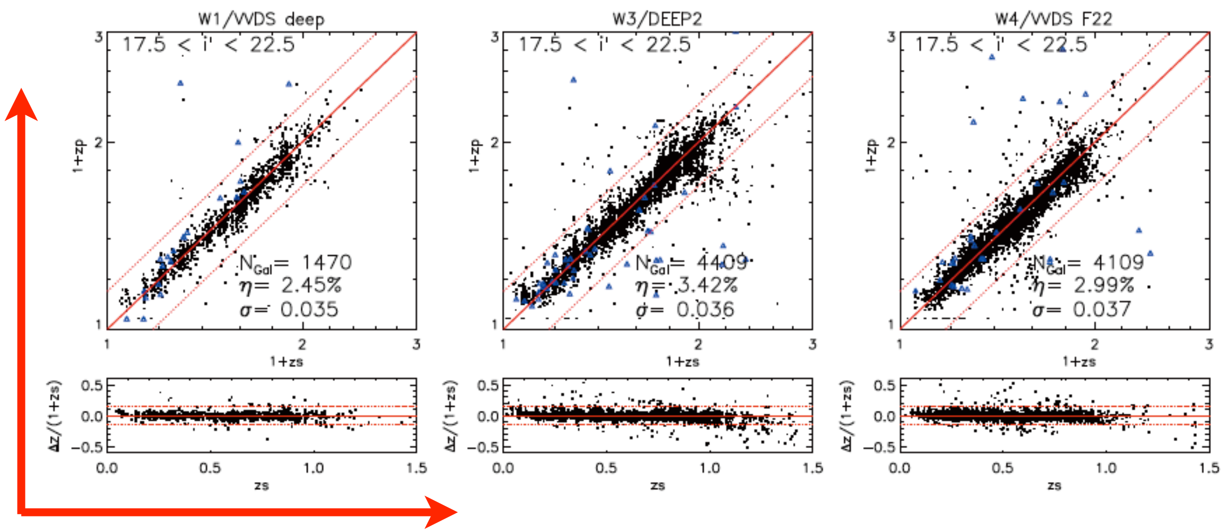
- VVDS Deep (Le Fèvre et al 2005)
- VVDS F22 Wide (Garilli et al 2008)
- DEEP2 (Davis et al 2007)

Photometric redshift accuracy

Photo-z's computed from the Terapix T0006 release using template fitting method and spectro-z calibration (Ilbert et al. 2006, JC et al. 2009):

1. correction of systematic offsets
2. adaptation of templates
3. use of n(z) prior

photo-z



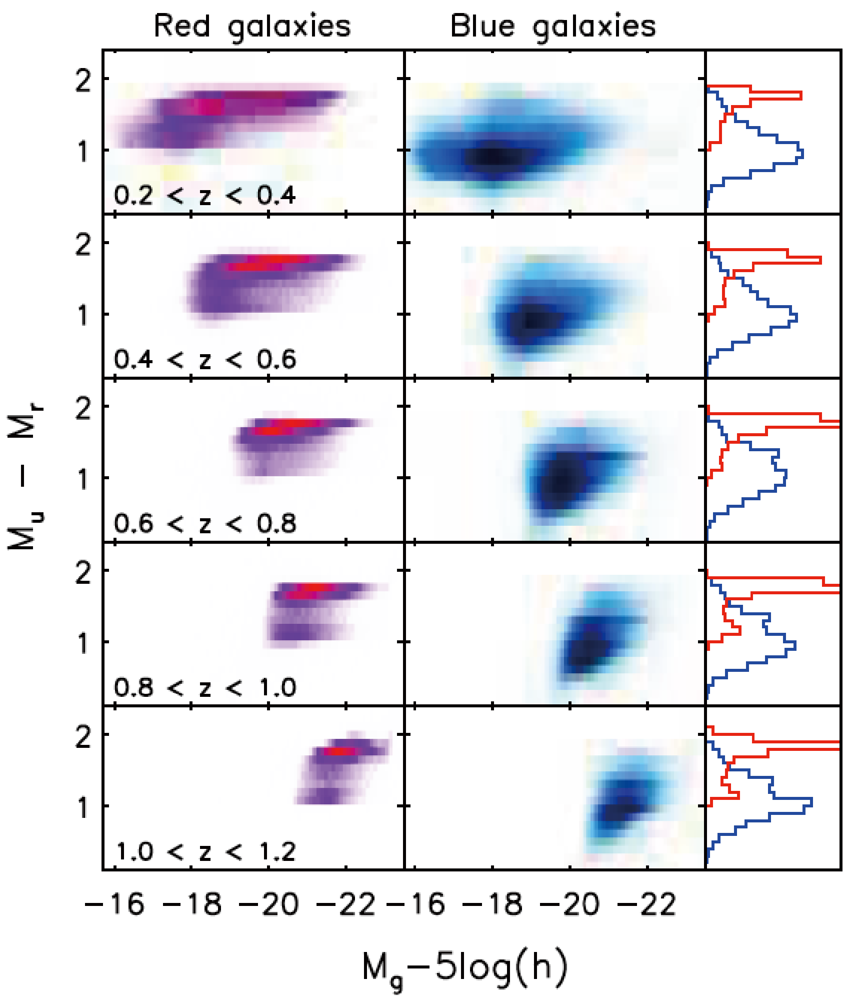
spectro-z

Dispersion: $\sigma = 0.04(1+z)$

Outlier rate: $\eta < 4\%$

The full sample includes 3,000,000 reliable redshifts (i < 22.5) over 4 independent fields covering 155 deg^2 (effective area).

Sample selection - galaxy type



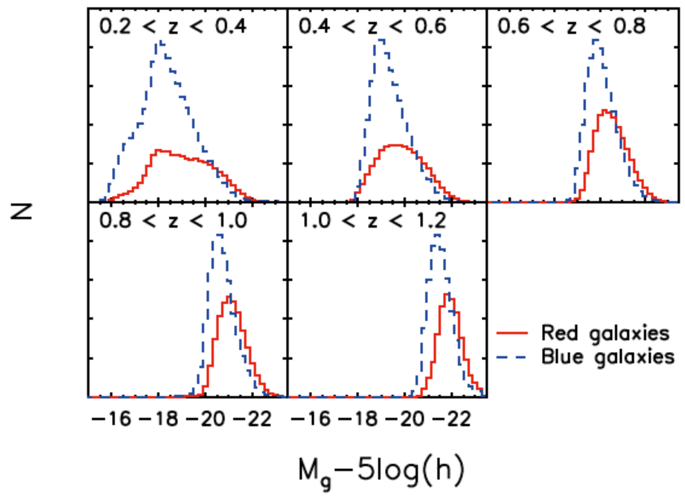
Total sample (i < 22.5): 3 000 000 galaxies divided into:

- red galaxies: El, Sbc
- blue galaxies: Sbc, Scd, Im, SB1, SB2

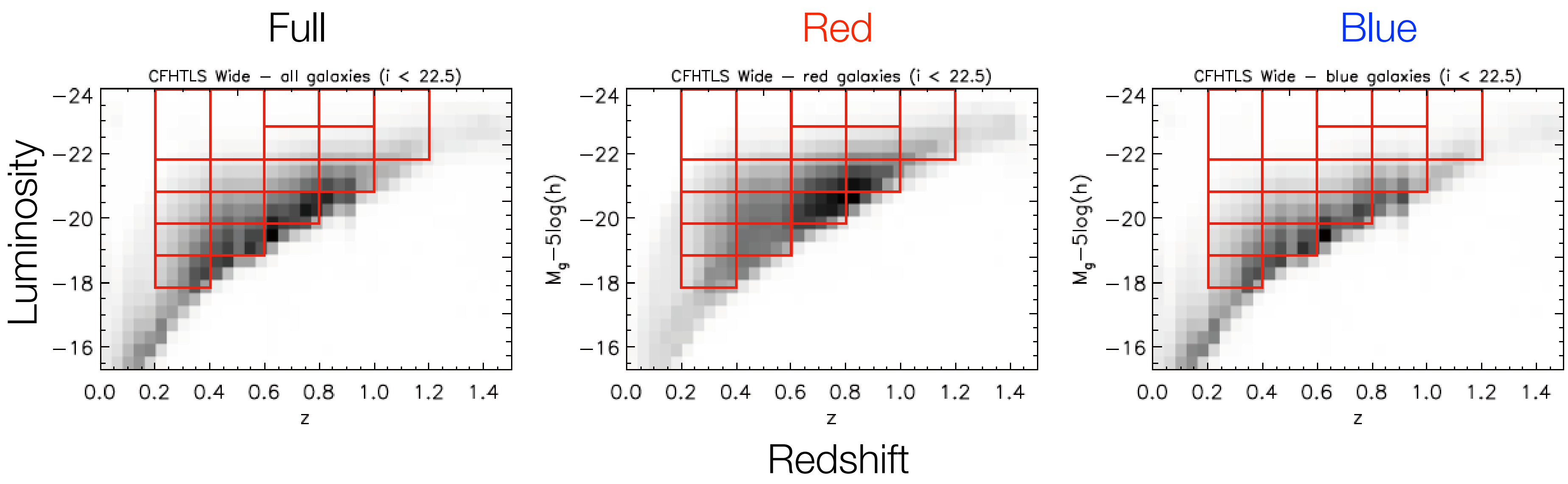
(equivalent to a colour selection)

In general blue galaxies are more numerous than red ones but:

- blue ones dominate faint samples
- and red ones dominate bright samples



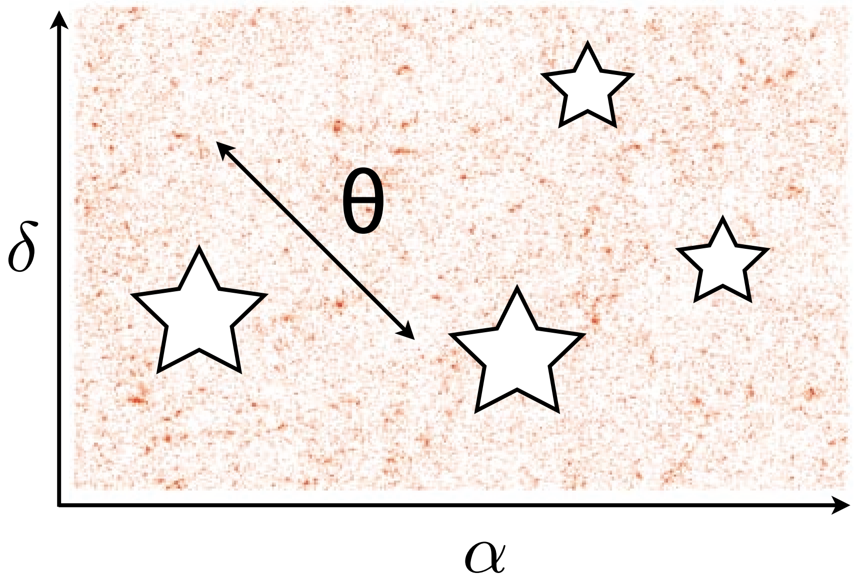
Sample selection - redshift/luminosity



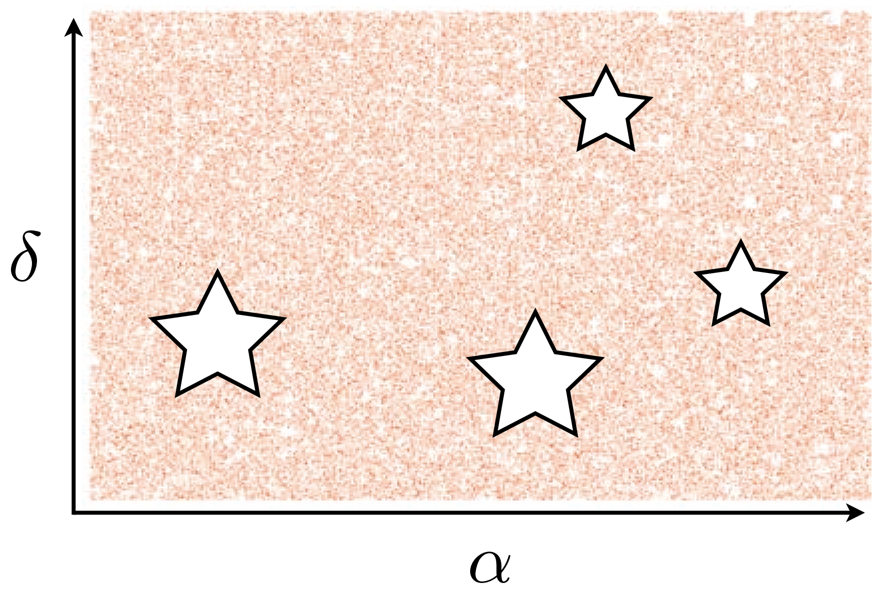
- threshold samples (guarantees the presence of central galaxies)
- five redshift bins, $0.2 < z < 1.2$
- 45 samples
- blue bright samples are discarded (weak clustering signal)
- larger samples have over 1 000 000 galaxies
- galaxy number density estimate: N_{total}/volume: $n_{\rm gal}^{\rm obs} = N_{\rm total} / \left[\Omega \int_{z_{\rm min}}^{z_{\rm max}} \frac{{\rm d}V}{{\rm d}z} {\rm d}z \right],$

Angular two-point correlation function $w \ \theta$

Data



Random (= same geometry)



Landy & Szalay (1993) estimator:
(low variance and bias)

$$w(\theta) = \frac{N_r(N_r - 1)}{N_d(N_d - 1)} \frac{\langle DD \rangle}{\langle RR \rangle} - \frac{N_r - 1}{N_d} \frac{\langle DR \rangle}{\langle RR \rangle} + 1$$

Fast angular correlation function measurements

PB: classic estimator scales as N^2 would take weeks for ~ 1 000 000 object samples

SOLUTION: for large angular separations, correlate boxes instead of individual objects and build optimised boxes with kd-tree

now scales as N log N

adaptive approximation ~ 0.5

size of the box

$$\alpha_w = \theta / d_b$$

angular separation

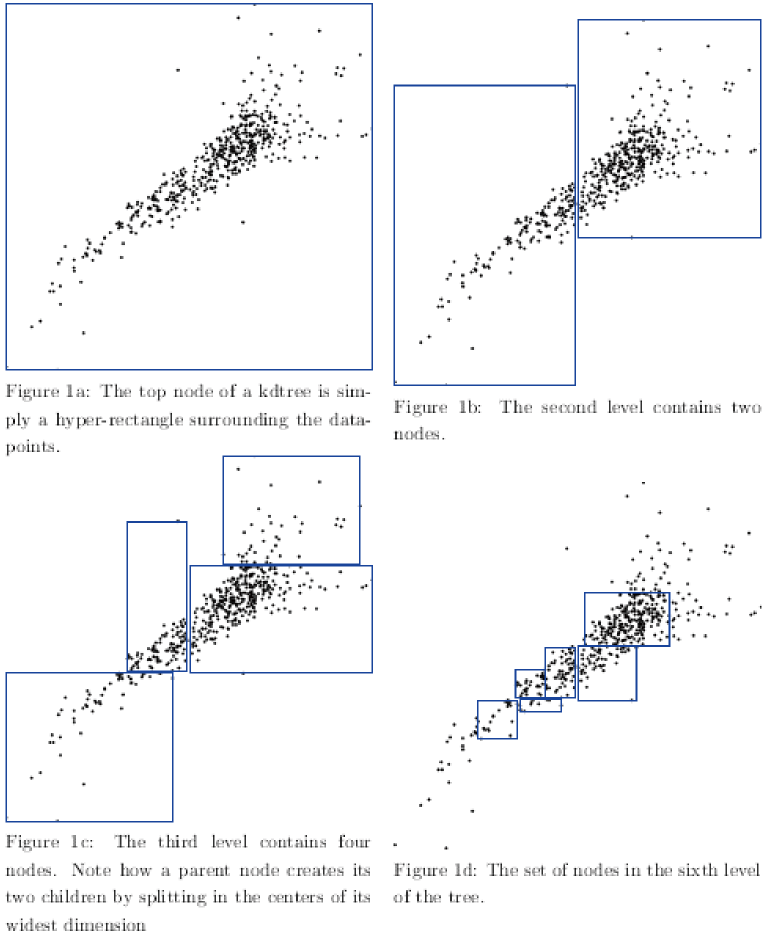


Figure 1a: The top node of a kdtree is simply a hyper-rectangle surrounding the datapoints.

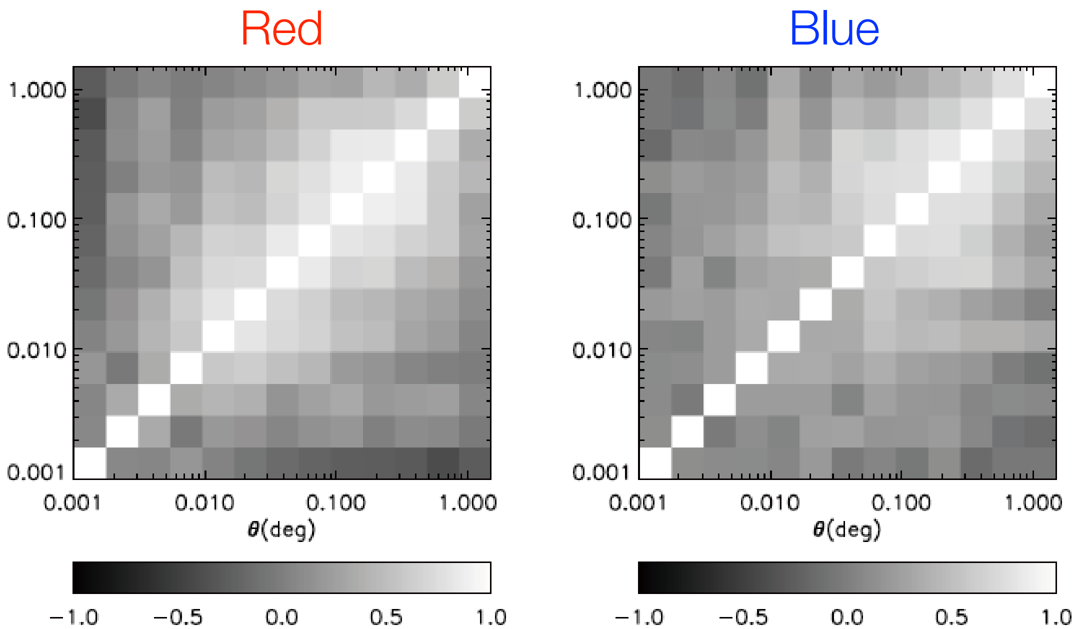
Figure 1b: The second level contains two nodes.

Figure 1c: The third level contains four nodes. Note how a parent node creates its two children by splitting in the centers of its widest dimension

Figure 1d: The set of nodes in the sixth level of the tree.

Moore et al. (2001)

Estimating errors - w(θ)



$$C(w_i, w_j) = \frac{N-1}{N} \sum_{l=1}^{N} (w_i^l - \overline{w}_i)(w_j^l - \overline{w}_j)$$

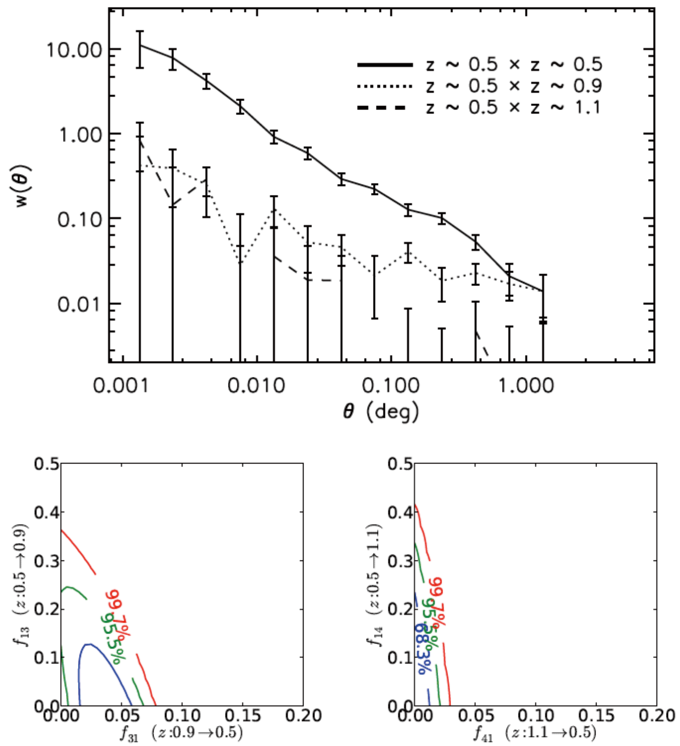
+ Error on n_{gal} from the field-to-field variance

Covariance matrices estimated from Jackknife estimator using 62 realisations. 4 independent fields allow a non-biased (although noisy) cosmic variance estimate

- large scales highly correlated
- small scales correlated for red galaxies

Estimating errors - photometric redshifts

- photo-z error estimated from PDF (cf previous slide)
- ξ(r) projected using redshift distribution of photo-z convolved with photo-z error
- cross-correlation check (a la Benjamin et al 2010): photo-z contamination should create a positive cross-correlation between redshift bins



results:

- small cross-correlation between adjacent bins likely due to the large-scale structure
- no significant contamination found between distant field
- red samples (better photo-z) perform the best

IV. Results: new insights on galaxy evolution since z ~ 1.2

Fitting HOD parameters

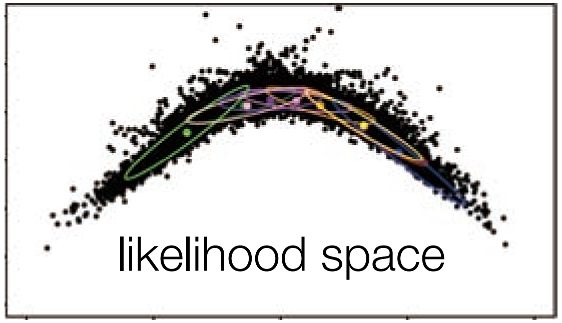
Likelihood: $\chi \quad \sum [w \quad \theta \; - w \qquad \theta \;] \, (C^{-} \;) \; [w \quad \theta \; - w \qquad \theta \;] \quad \frac{[n \quad - n \qquad]}{\sigma}$

- constraints from w(θ) + galaxy number density

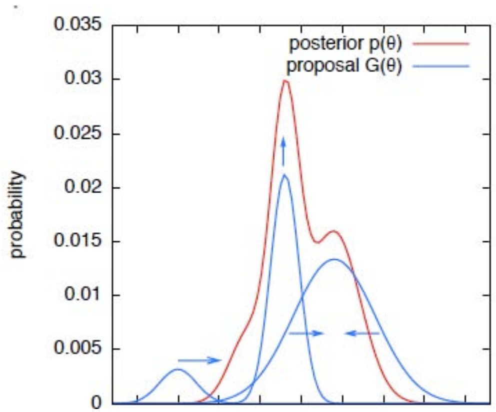
- population monte carlo (PMC): likelihood space is sampled from a proposal (importance sampling method, see Cappé et al. 2004)

- the proposal is iteratively adapted to match the posterior (convergence: “perplexity” → 1)

- same results as MCMC but not point is rejected and the method is easy to parallelize (10 times faster)



Wraith et al. (2009)

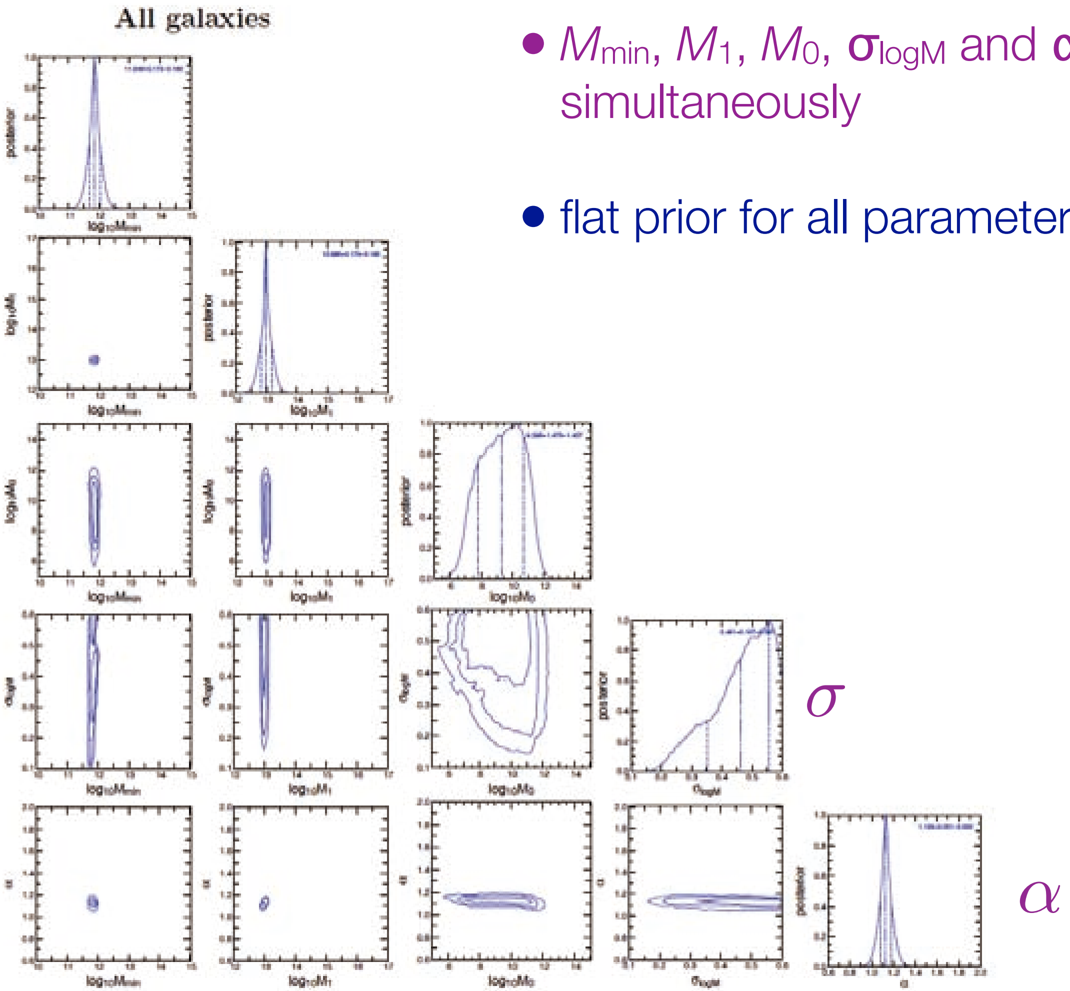


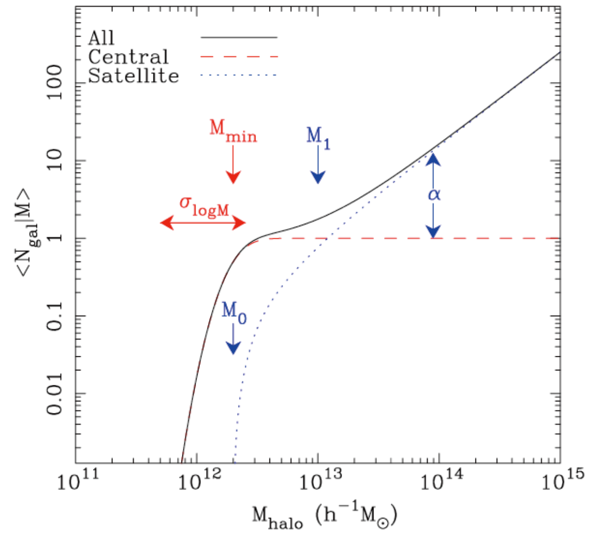
(M. Kilbinger)

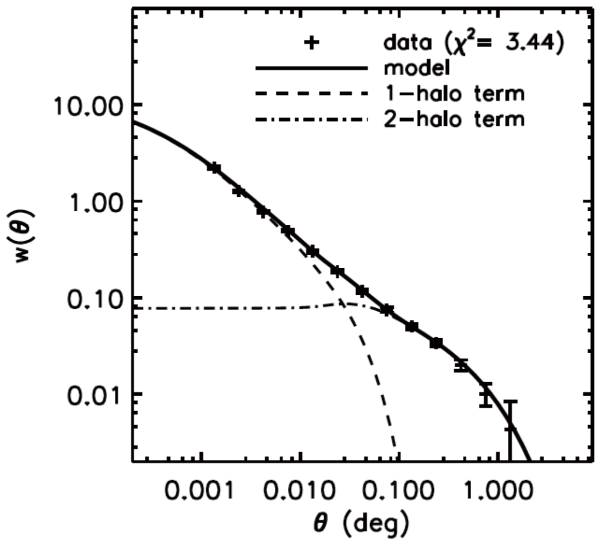
Fitting HOD parameters

- *M*min, *M*1, *M*0, **σ**logM and **α** fit simultaneously

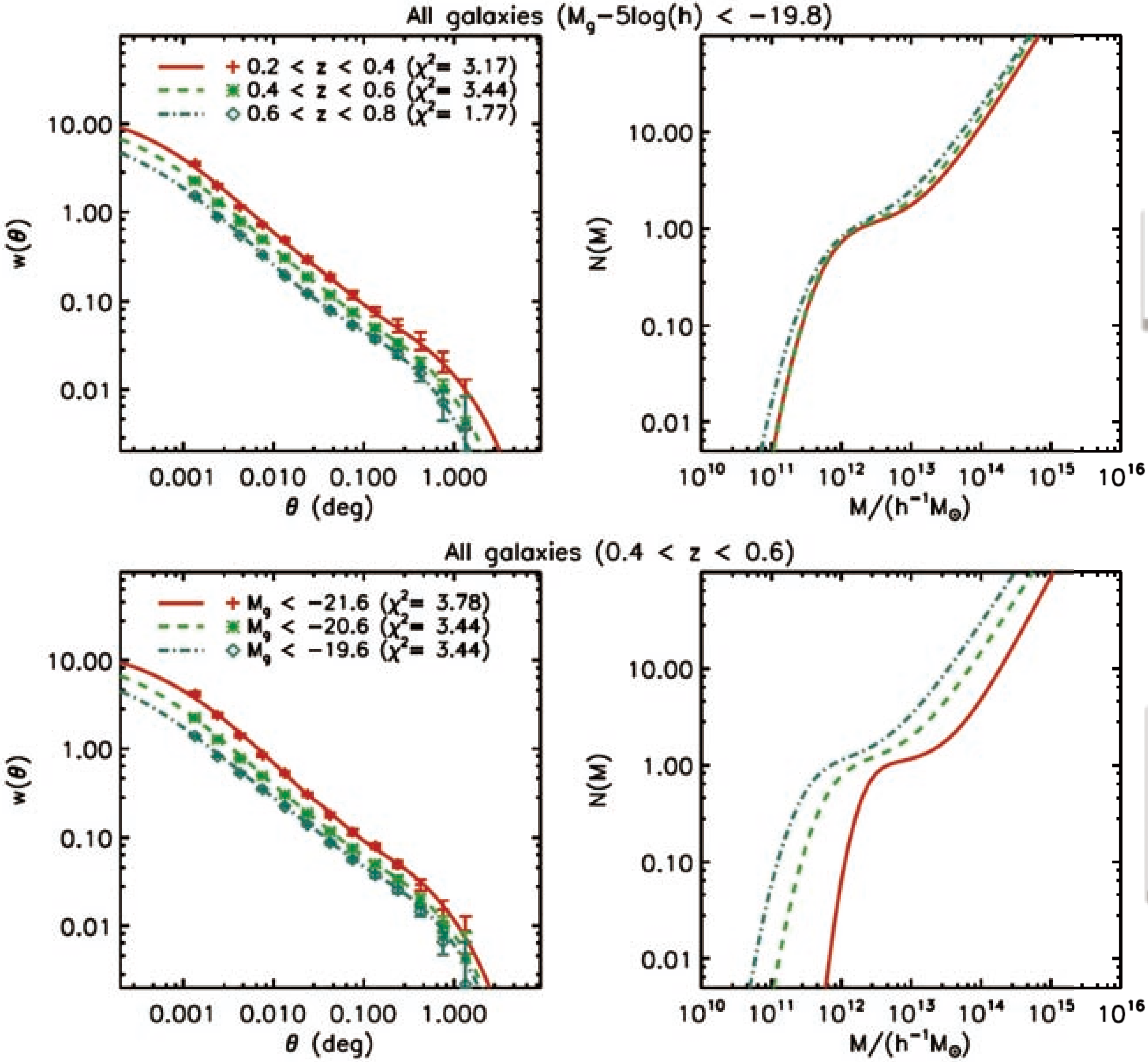
- flat prior for all parameters







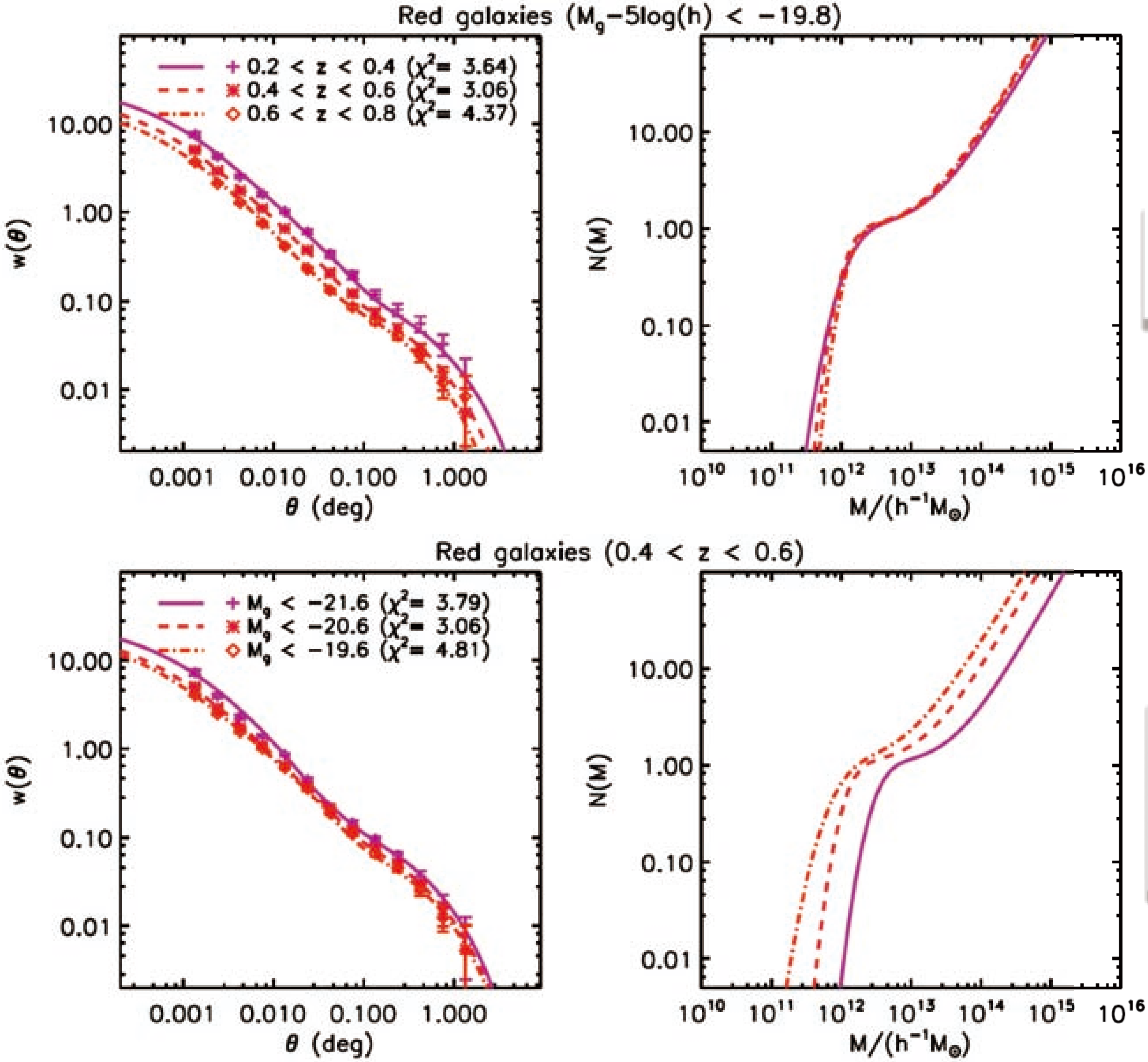
Clustering measurements - all galaxies



redshift evolution

Luminosity dependence

Clustering measurements - red galaxies



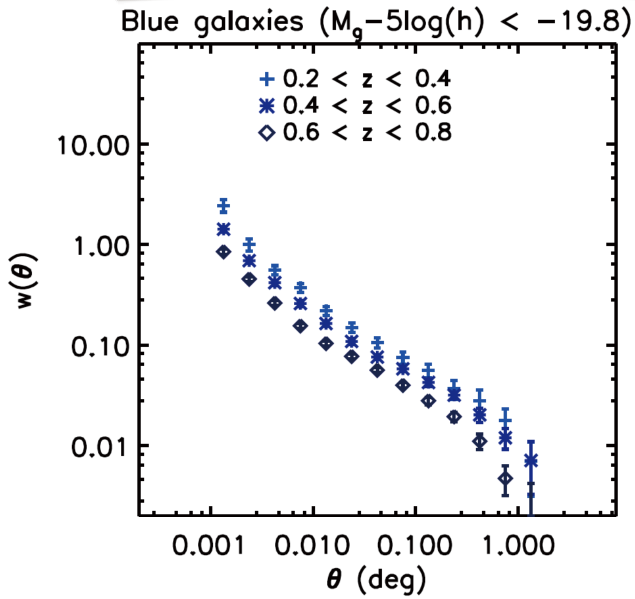
redshift evolution

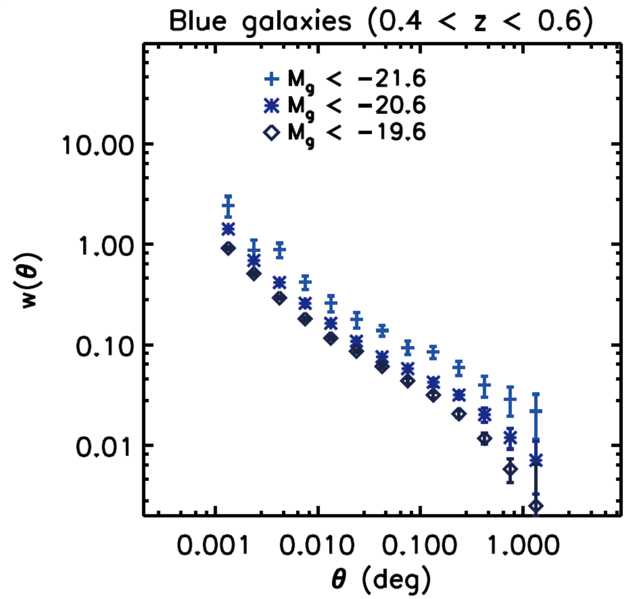
Luminosity dependence

Clustering measurements - blue galaxies

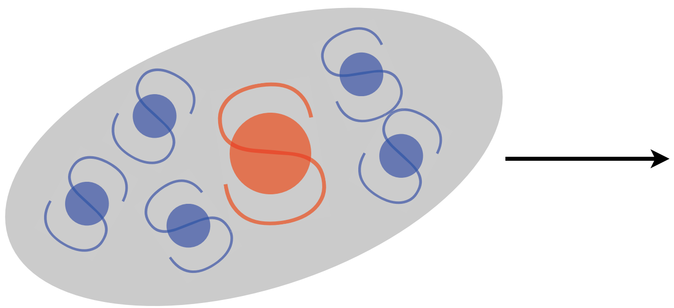
redshift evolution

Luminosity dependence



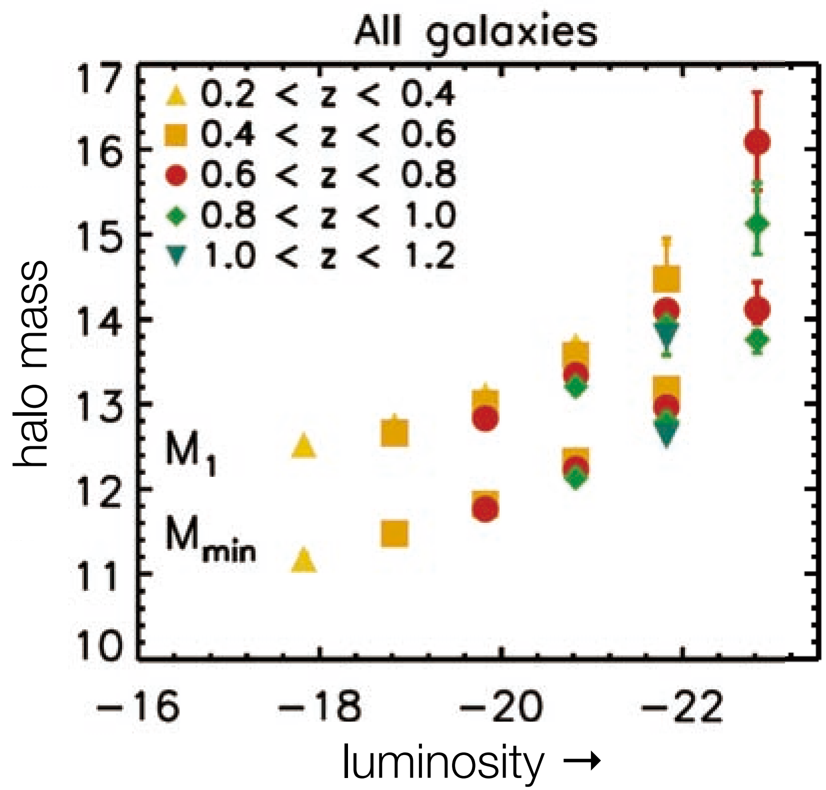


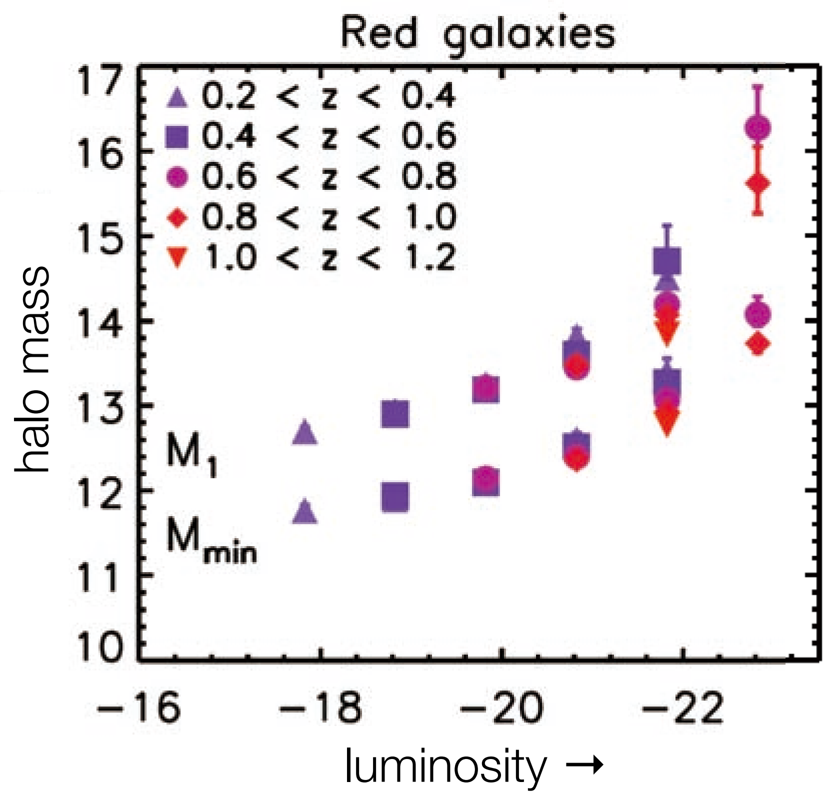
CAVEAT: blue satellite galaxies may belong to a red central galaxy.



The model based on the separation central/satellite breaks down. No HOD fitting for blue samples.

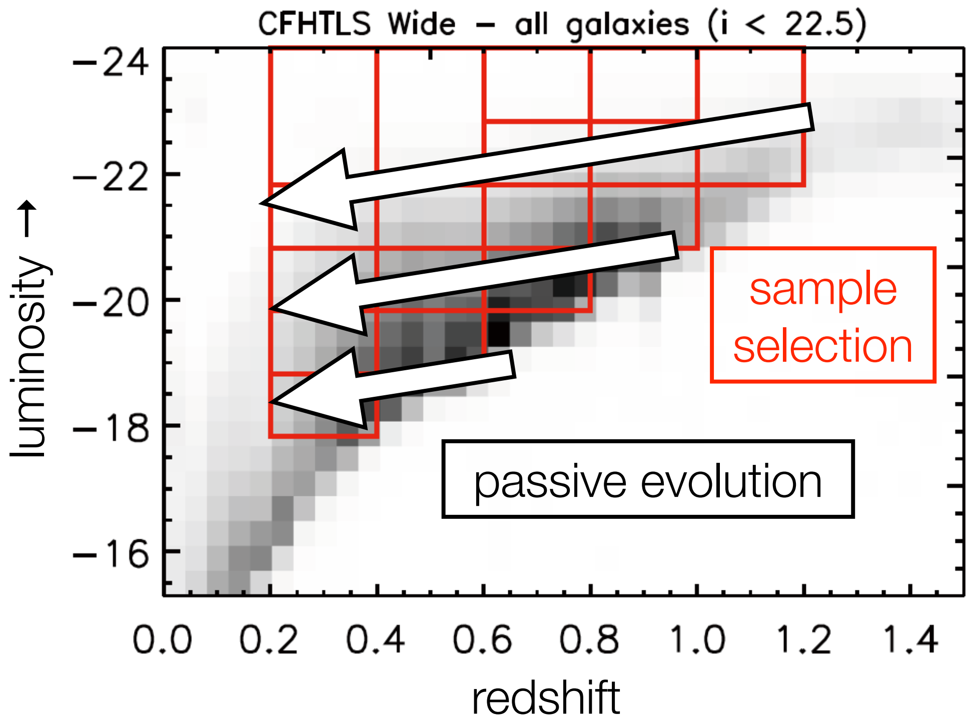
Halo mass vs galaxy luminosity





- brighter galaxies reside in more massive haloes
- halo masses decrease with redshift
- red galaxies reside in more massive haloes than blue ones

Redshift evolution of L/M_h?

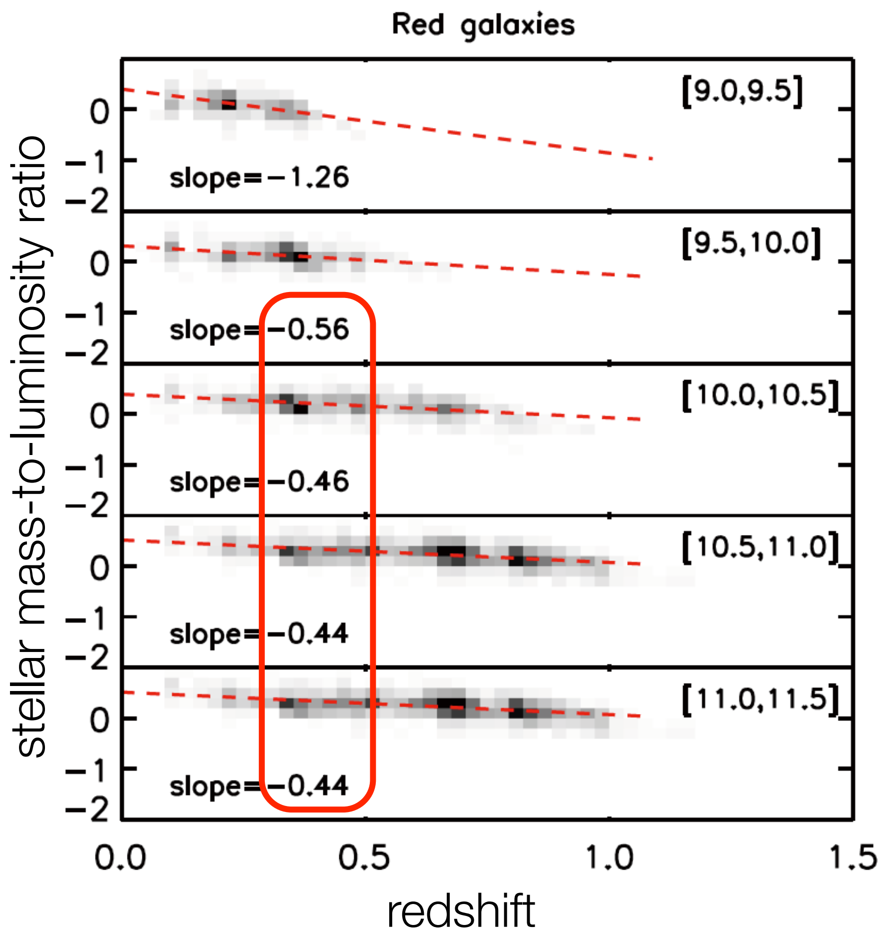


- **PB**: galaxies experience passive evolution (due to stellar population ageing)
- a constant luminosity selection "sees" different populations at different redshifts
- **consequence**: we observe less massive galaxies at higher redshift

decrease of M_h is partially due to this selection effect

Transforming luminosity into stellar mass

COSMOS 30-band stellar masses vs L_B:



From Ilbert et al. (2010)

No stellar masses (yet) in CFHTLS, but luminosity to stellar mass relation derived from COSMOS

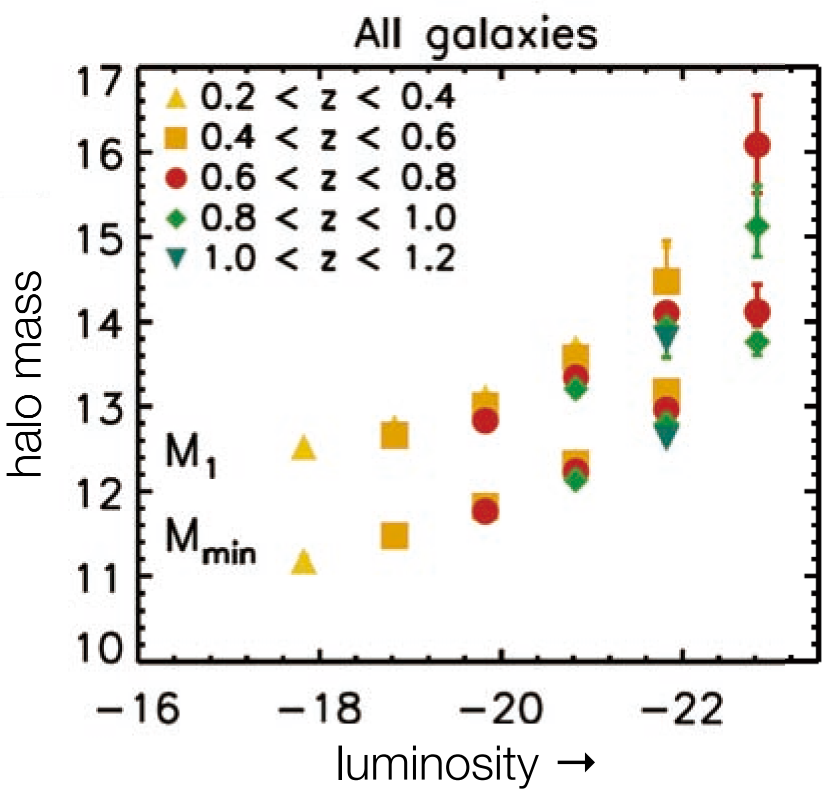
For red galaxies:

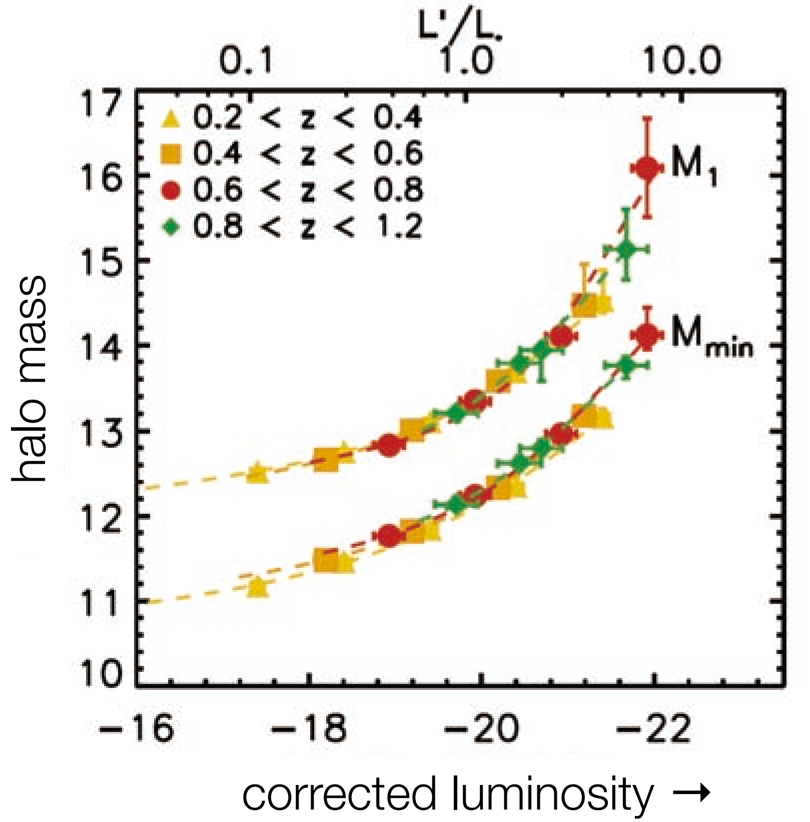
- stellar mass proportional to luminosity:

For all galaxies:

- non trivial relation due to the mixing of red and blue galaxies
- applied the “red” correction but probably underestimates the faint luminosity evolution

Stellar-to-halo mass relationship - all galaxies





stellar mass proxy

JC et al. (arXiv:1107.0616)

Stellar-to-halo mass relationship - all galaxies

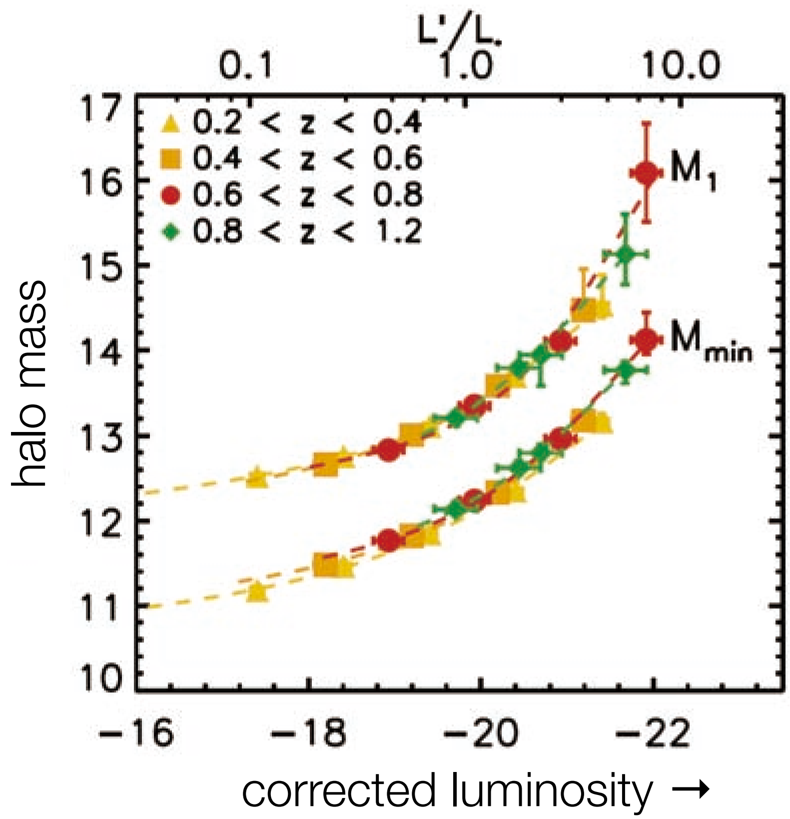
- relation between halo mass and central galaxy stellar mass
- redshift evolution of M_{star}/M_h
- but uncertainties at faint luminosity (where blue galaxies dominate)
- parameterised relation between luminosity and halo mass:

(Zehavi et al. 2010)

$$\frac{\;'\;}{\;*\;} \qquad \left(\frac{\qquad}{\qquad}\right)^{\alpha} \qquad \left(-\frac{\qquad}{\qquad} \qquad\right)$$

stellar mass proxy

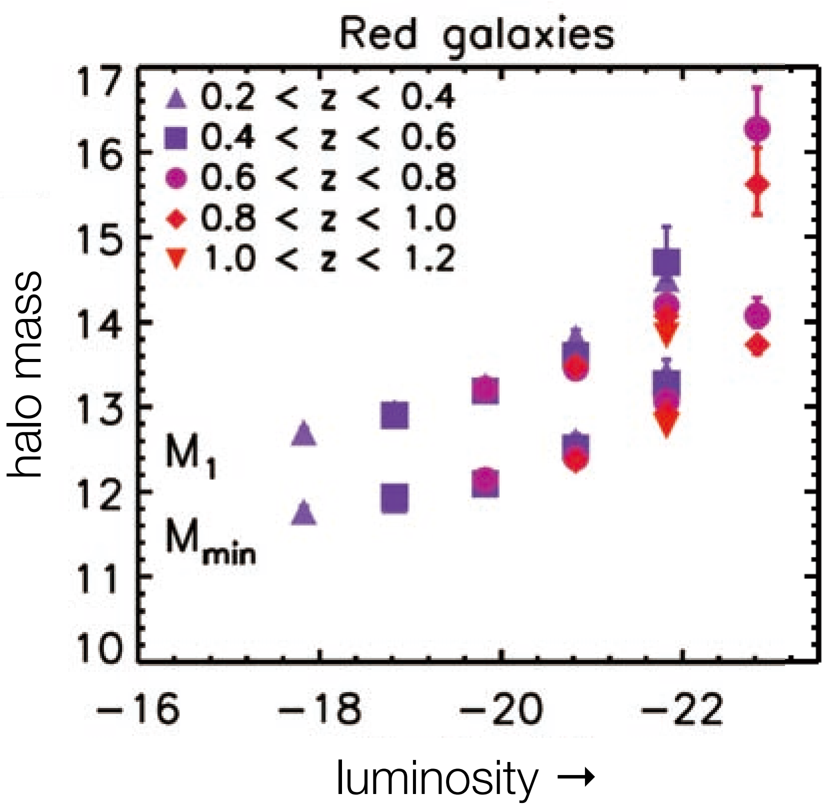
halo mass

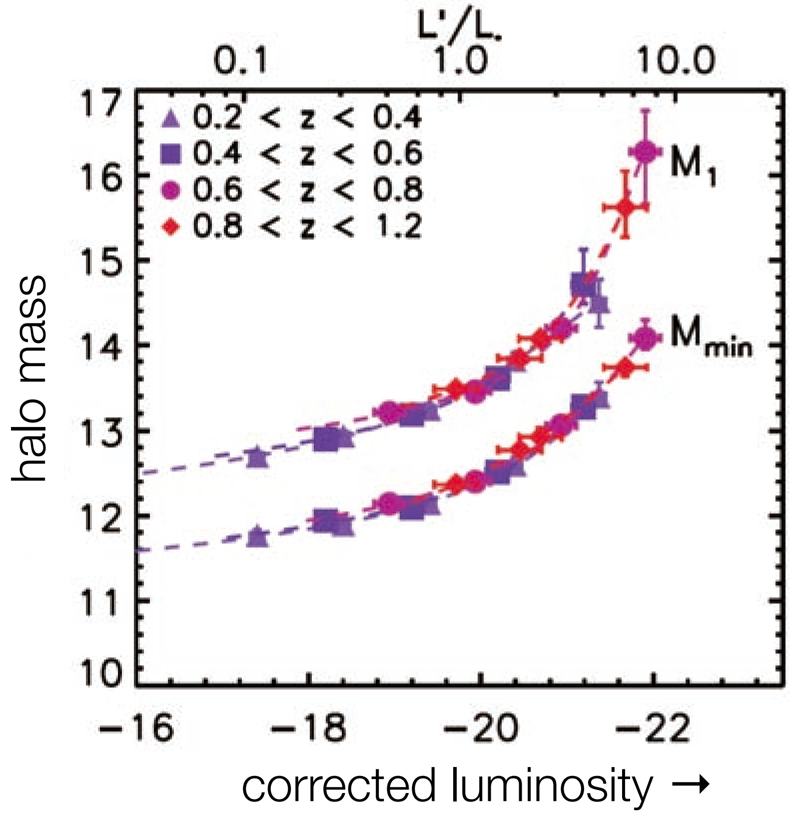


stellar mass proxy

JC et al. (arXiv:1107.0616)

Stellar-to-halo mass relationship - red galaxies



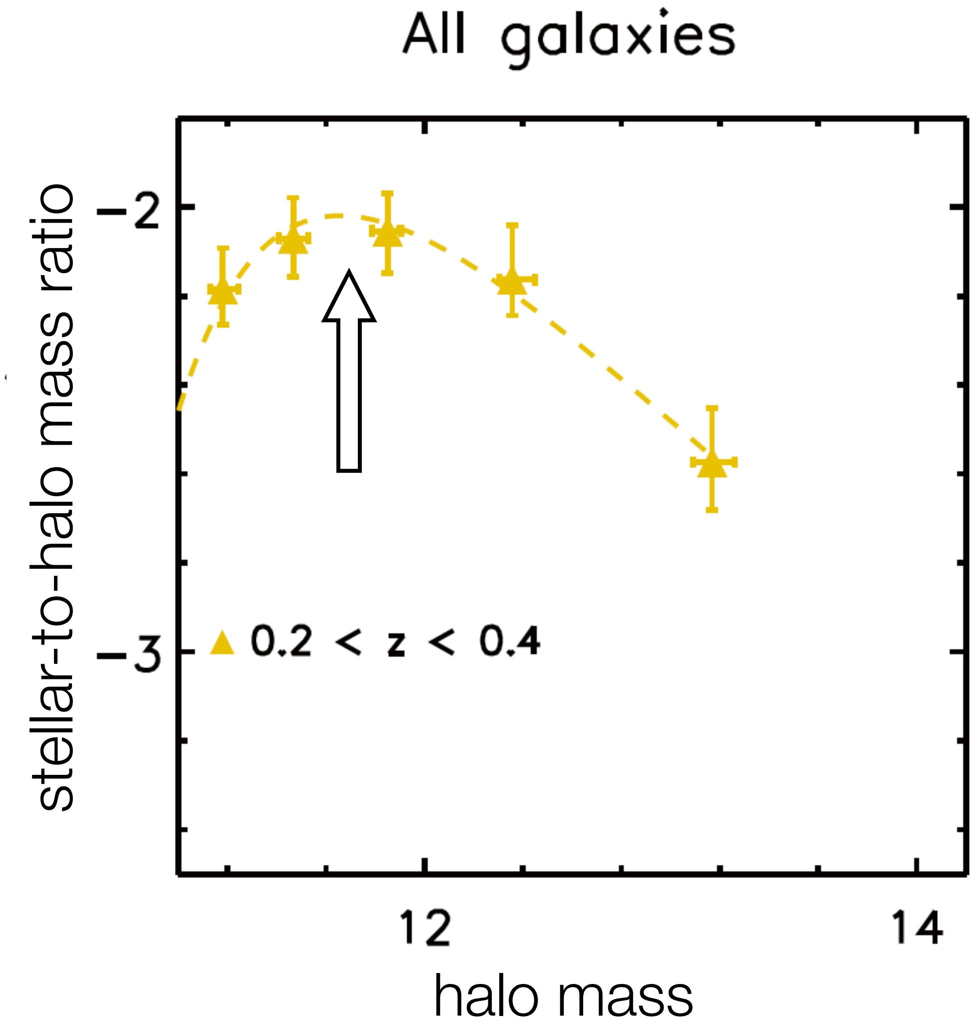


stellar mass proxy

- robust relation between halo mass and stellar mass
- weak redshift evolution of M_{star}/M_h

JC et al. (arXiv:1107.0616)

Stellar-to-halo mass ratio vs halo mass



We can re-write

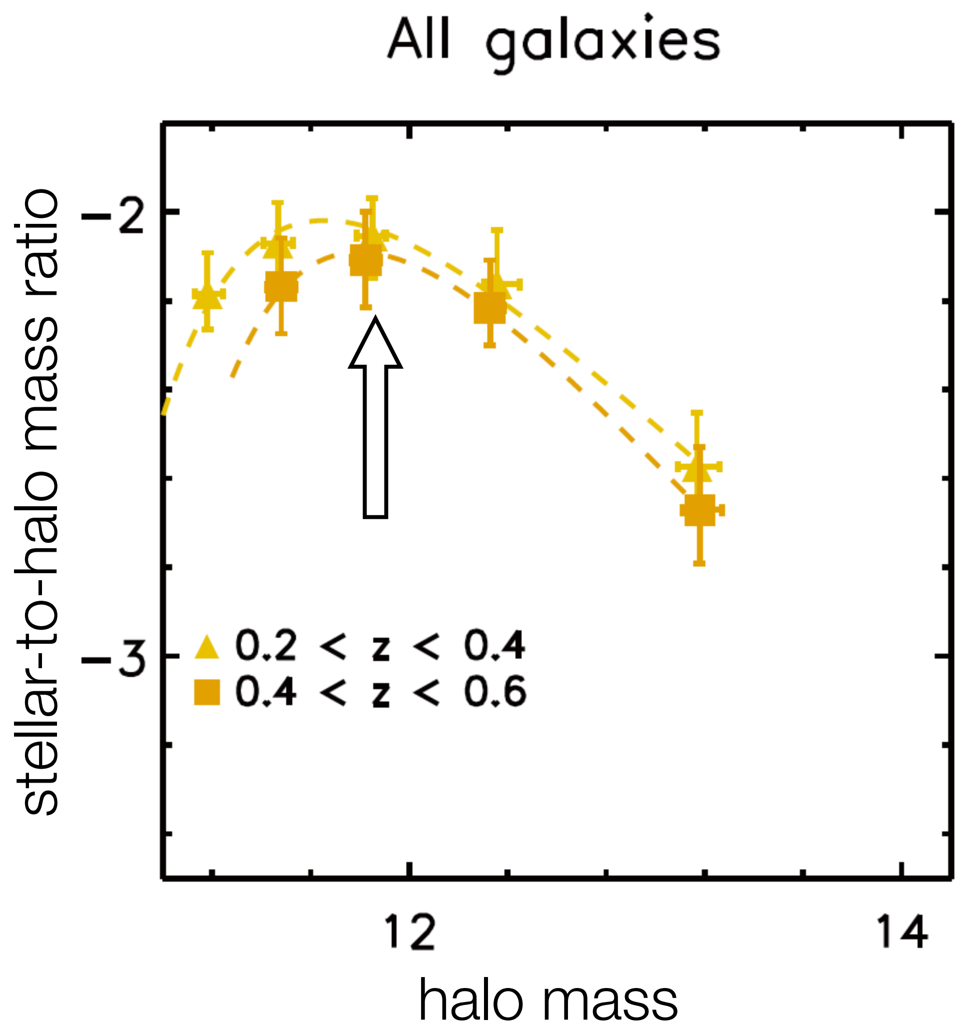
into

⇔ stellar-to-halo mass ratio

as function of halo mass

JC et al. (arXiv:1107.0616)

Stellar-to-halo mass ratio vs halo mass

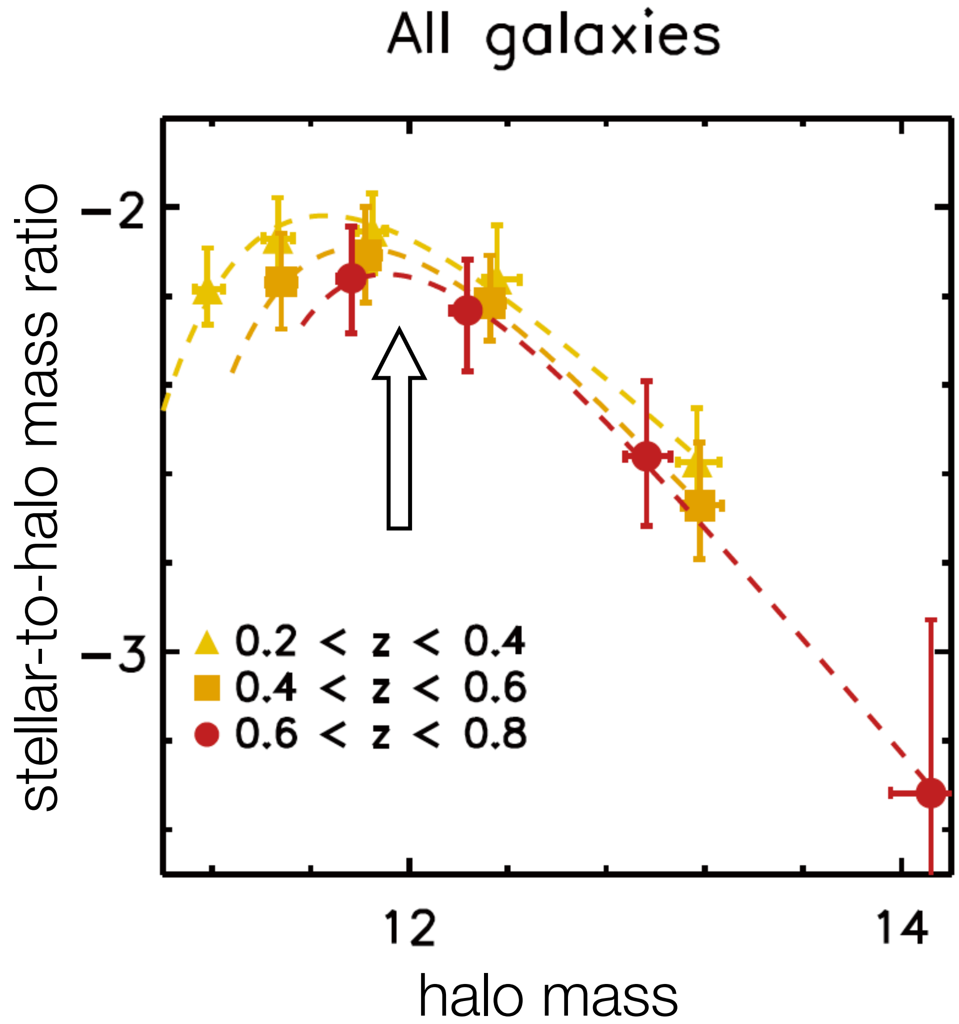


Redshift evolution:

- $M_{h,peak}$ shifts at higher mass
- M_{star}/M_h decreases with redshift

JC et al. (arXiv:1107.0616)

Stellar-to-halo mass ratio vs halo mass

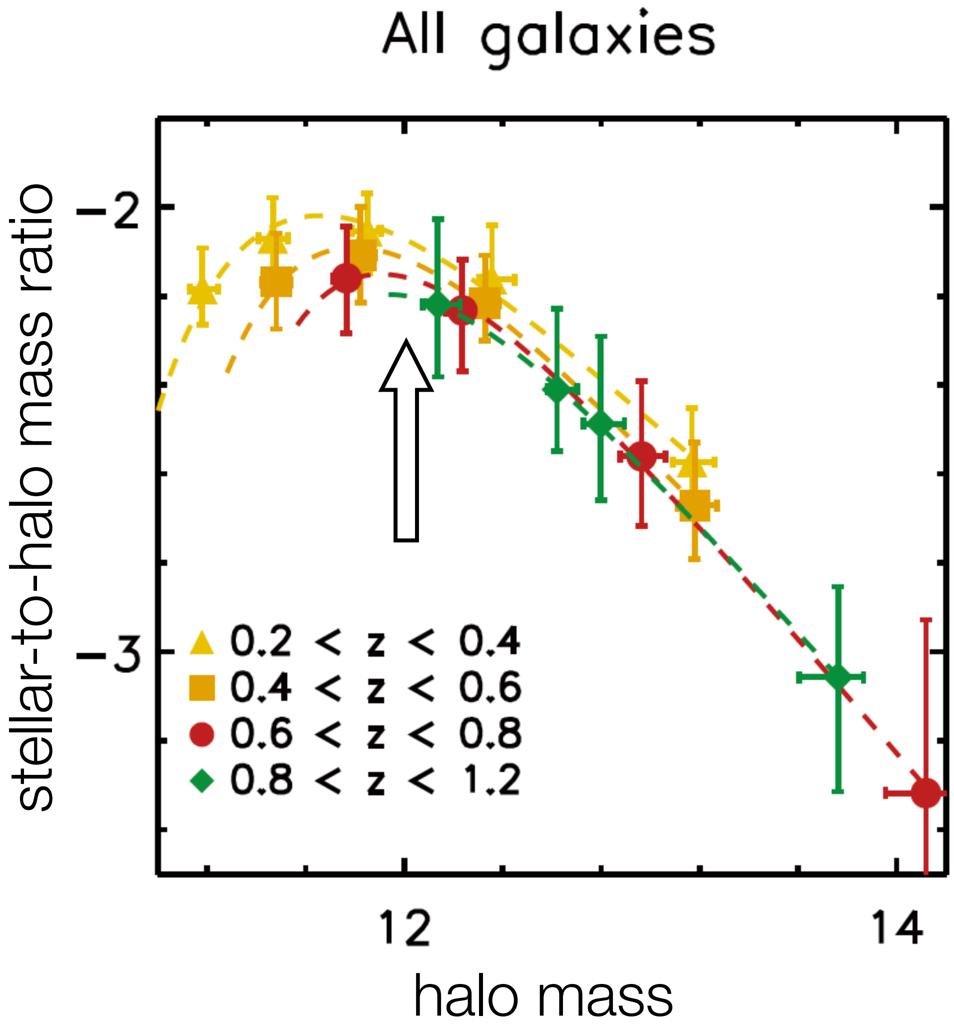


Redshift evolution:

- trends confirmed at higher redshift
- the position of the peak does not depend on L→M_{star}
- the amplitude variation depends on L→M_{star} (larger variation expected)

JC et al. (arXiv:1107.0616)

Stellar-to-halo mass ratio vs halo mass

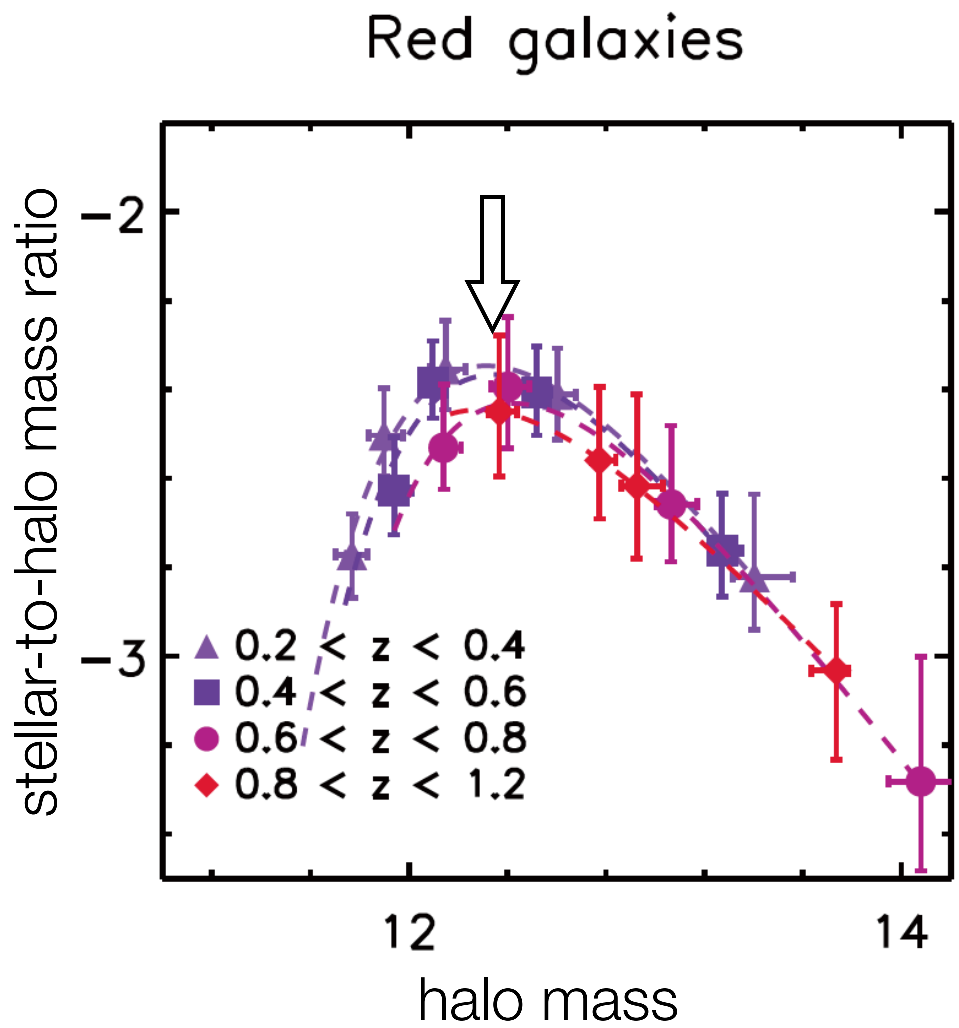


Redshift evolution:

- the peak is poorly constrained in the highest redshift bin
- stronger evolution seen in low-mass haloes than in high-mass ones

JC et al. (arXiv:1107.0616)

Stellar-to-halo mass ratio vs halo mass



Red galaxies

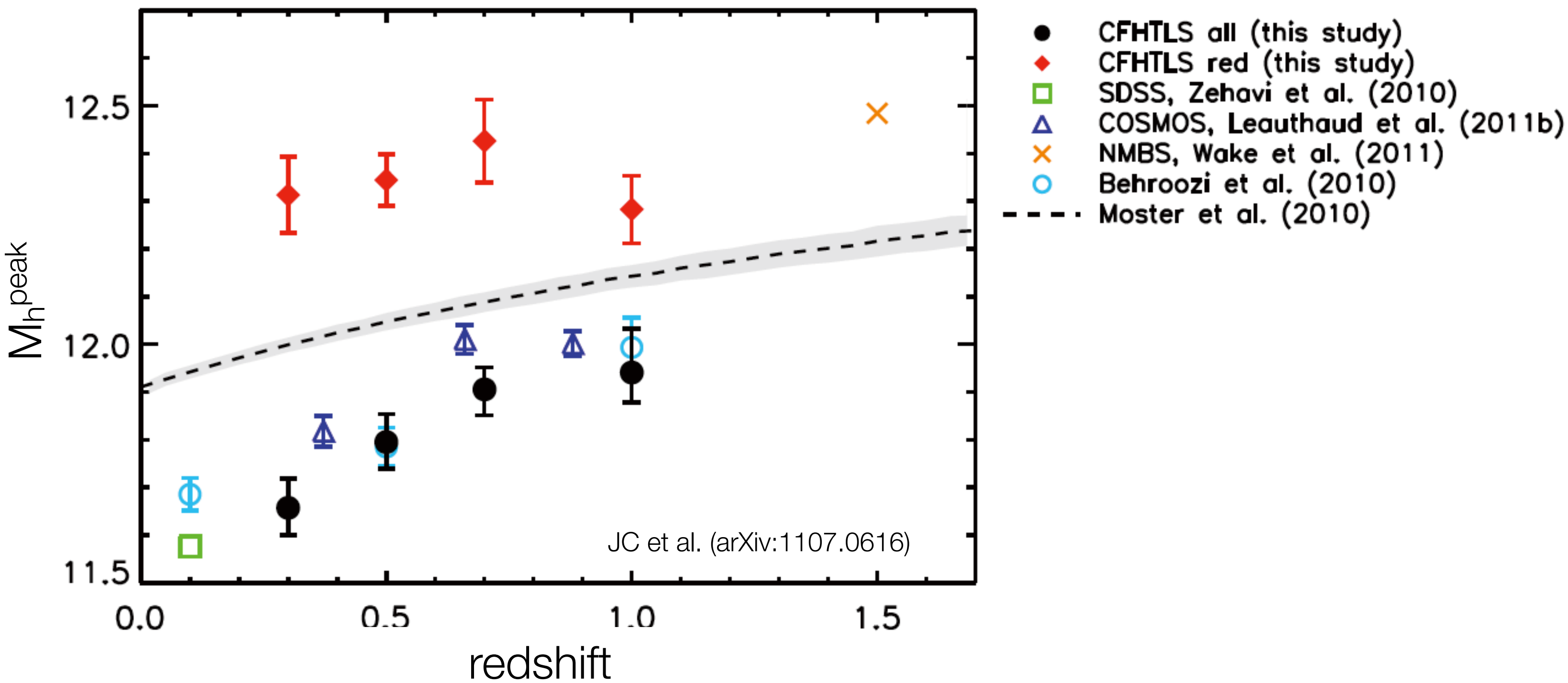
- the peak is at higher mass
- slower redshift evolution in low-mass haloes

JC et al. (arXiv:1107.0616)

Implications for galaxy evolution

- in the local Universe, the stellar content has been most efficiently accumulated in haloes of mass (M_{peak}) ~ 10^{12} M_{sun}
- star formation is "quenched" by feedback processes at both halo mass limits
- the shift of the peak with redshift is caused by a differential evolution in low- and high-mass haloes ("downsizing" effect)
- in the full sample the evolution is more rapid in low-mass haloes (samples dominated by blue galaxies)
- active star formation raises the SMHR with time (decreases with redshift)
- red galaxy SMHR do not show significant evolution, consistent with passive evolution

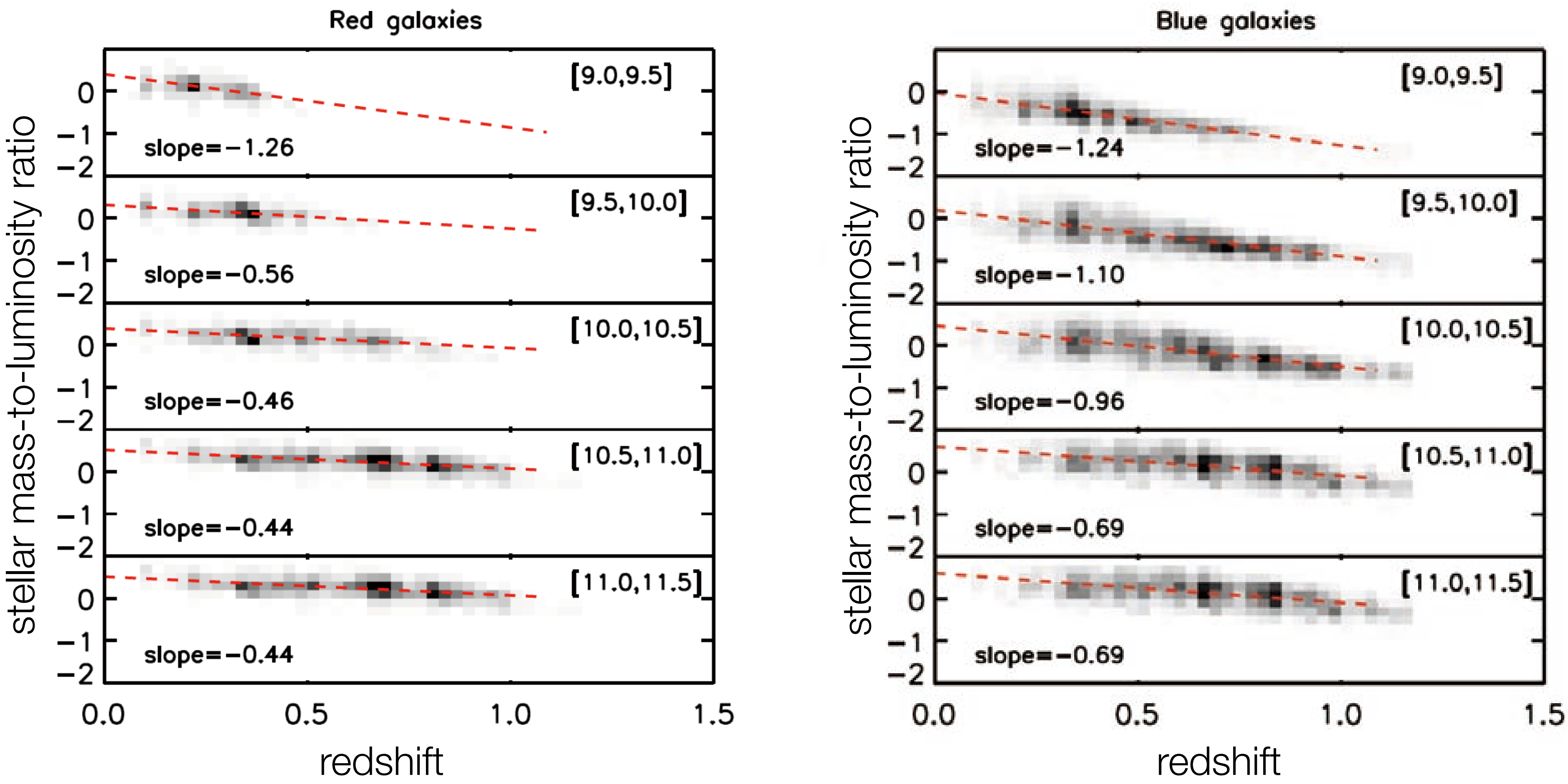
Comparison with the literature



- Similar trend for observations and semi-analytic simulations
- Excellent agreement with SDSS
- Lower value than in COSMOS (cosmic variance issue?)
- no significant evolution for red galaxies: passive evolution since z ~ 1.2?

Should we trust the L→M_{star} conversion?

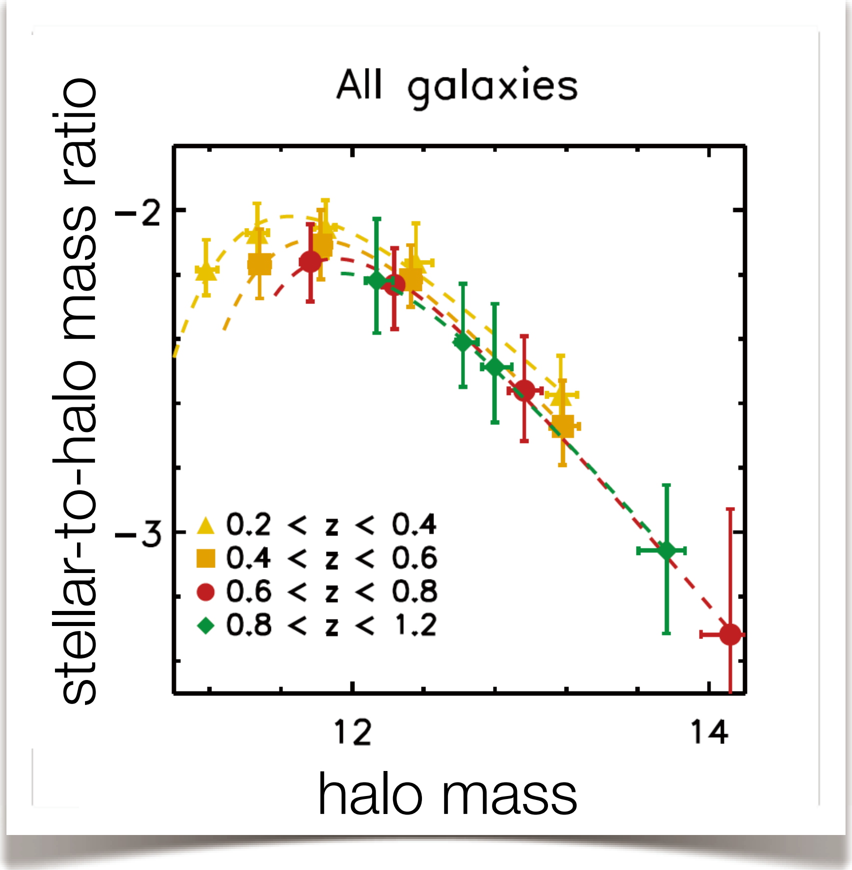
Red and blue galaxies have different M_{star}/L relations. Besides blue galaxy M_{star}/L depends on L. For the full sample M_{star}/L is fairly complex



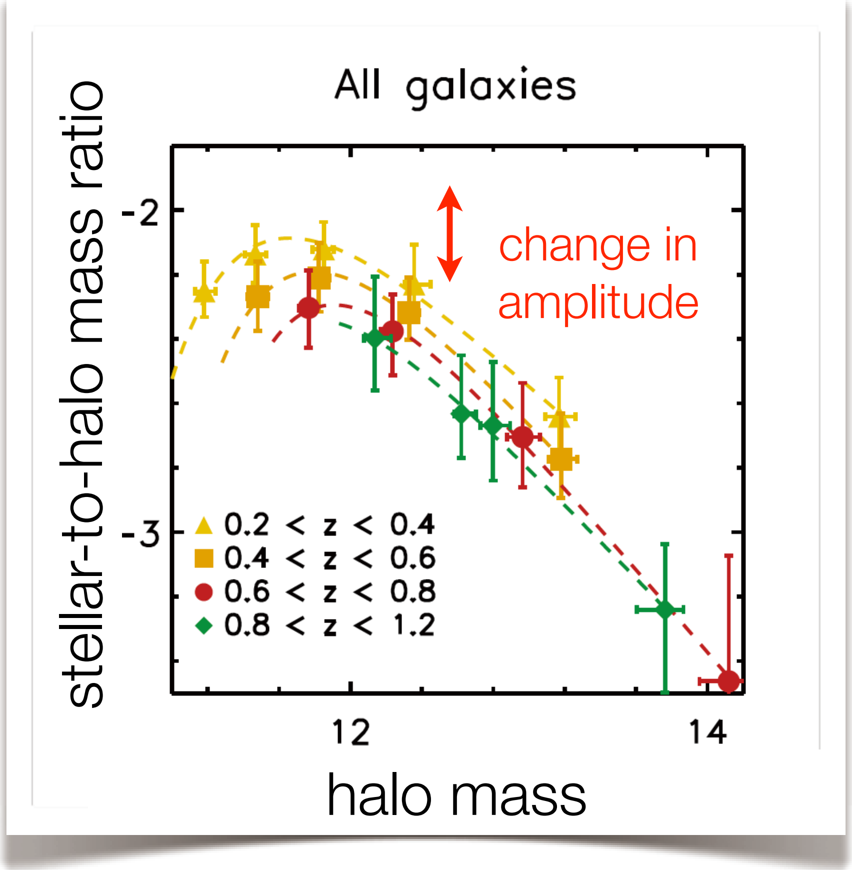
Should we trust the L→M_{star} conversion?

PB: "red" correction assumed for all samples. How do results change if we applied an "extreme" correction based in blue galaxy evolution?

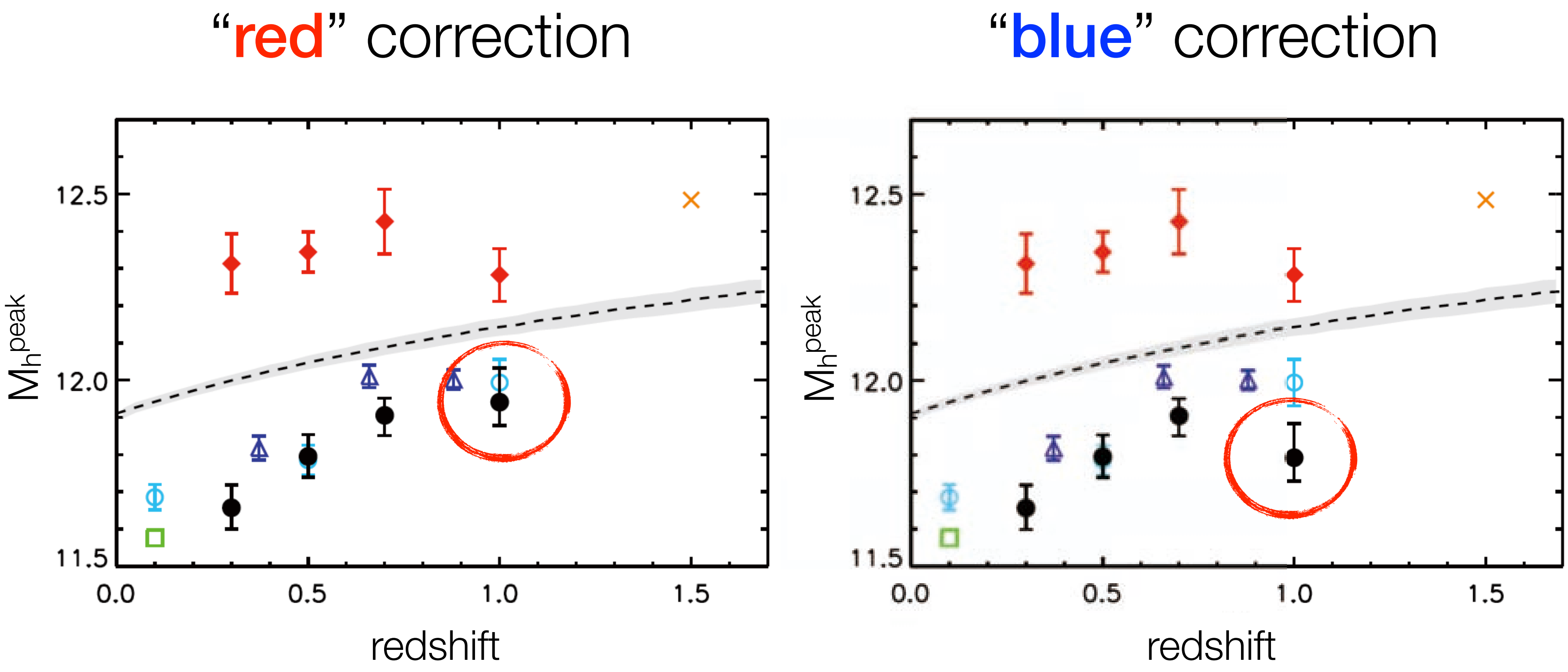
"**red**" correction



"**blue**" correction

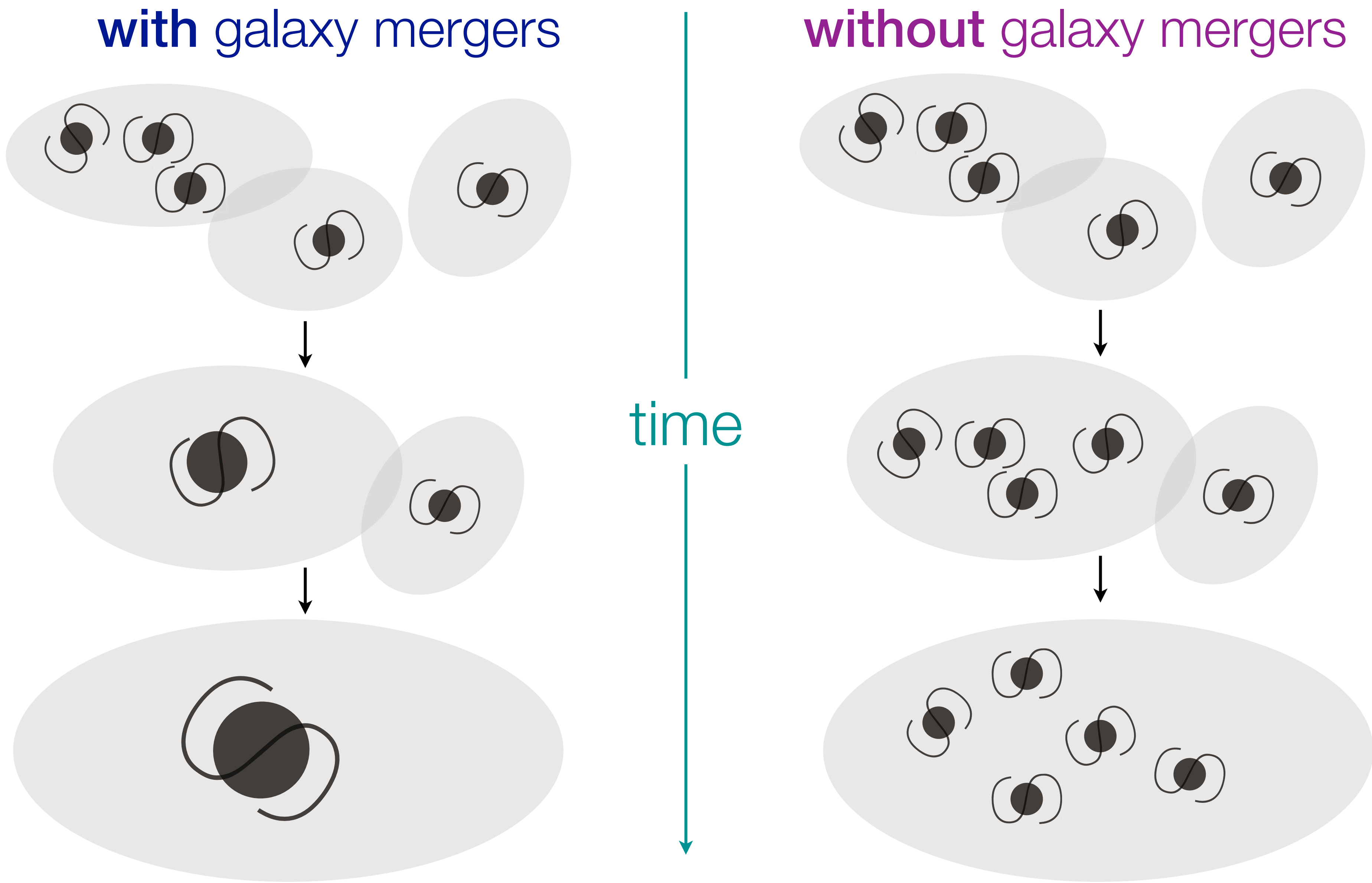


Should we trust the L→M_{star} conversion?

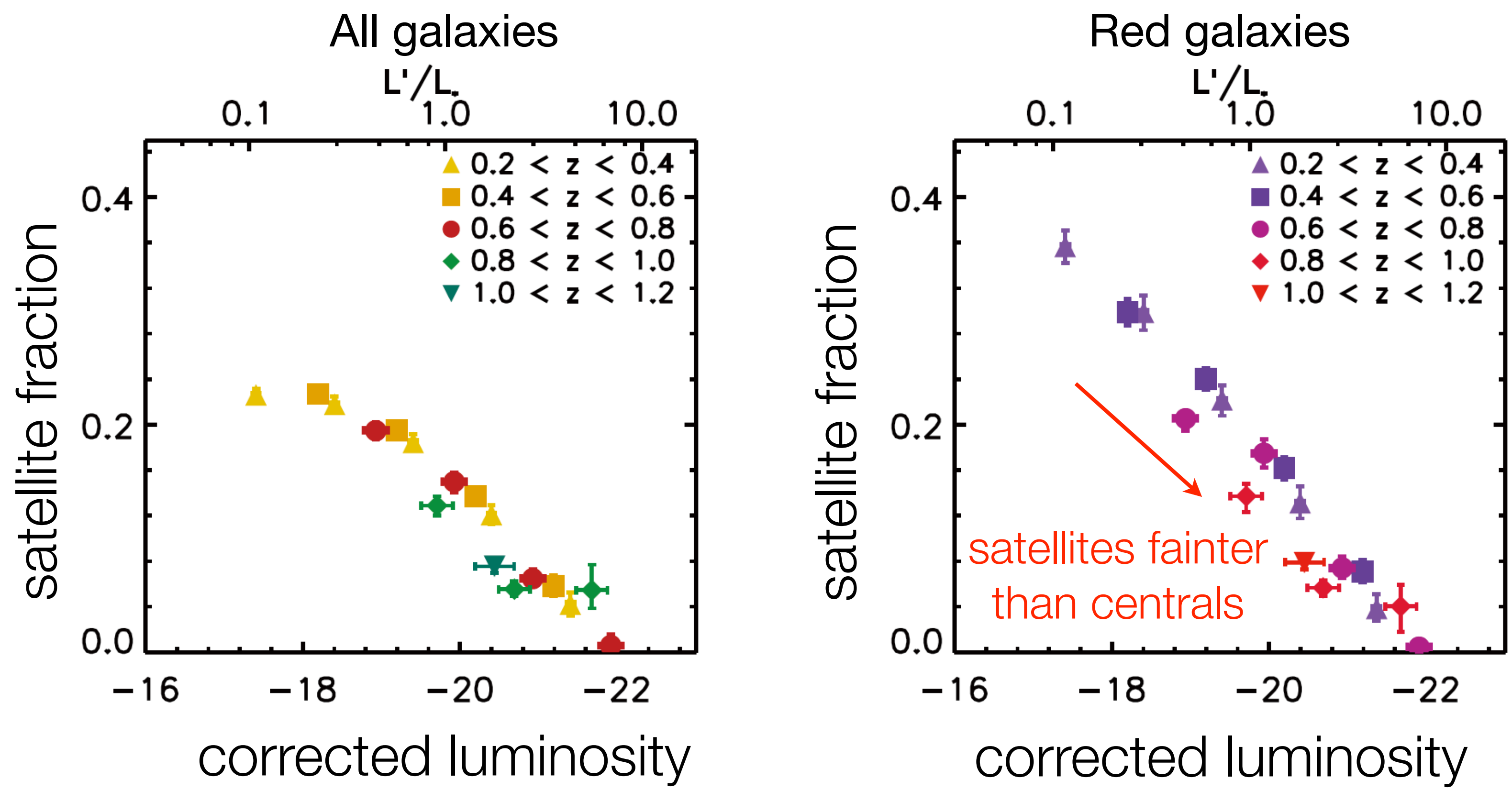


- amplitude of the peak strongly depends on the correction
- position of the peak is very robust in the range [0.2,0.8]
- but high redshift point is weakly constrained

The importance of mergers in galaxy evolution



Satellite galaxy fraction

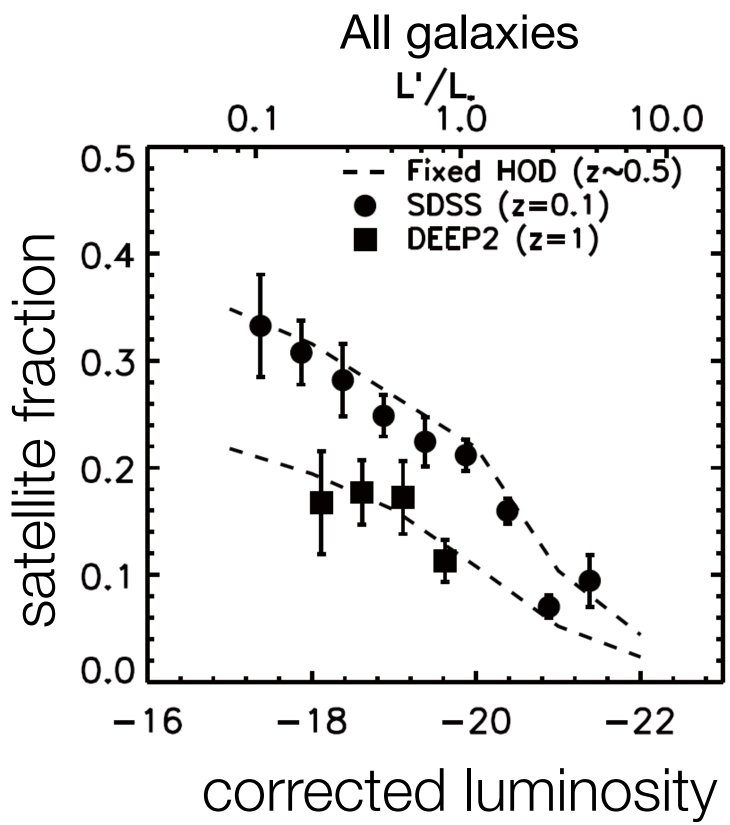


- number of red satellites is larger than in full sample

JC et al. (arXiv:1107.0616)

- increases (decreases) with time (redshift)
- flattens at high luminosity (larger number of small haloes with single galaxies)

Satellite galaxy fraction - redshift evolution



$\equiv$ $\otimes$

halo mass function (M,z)

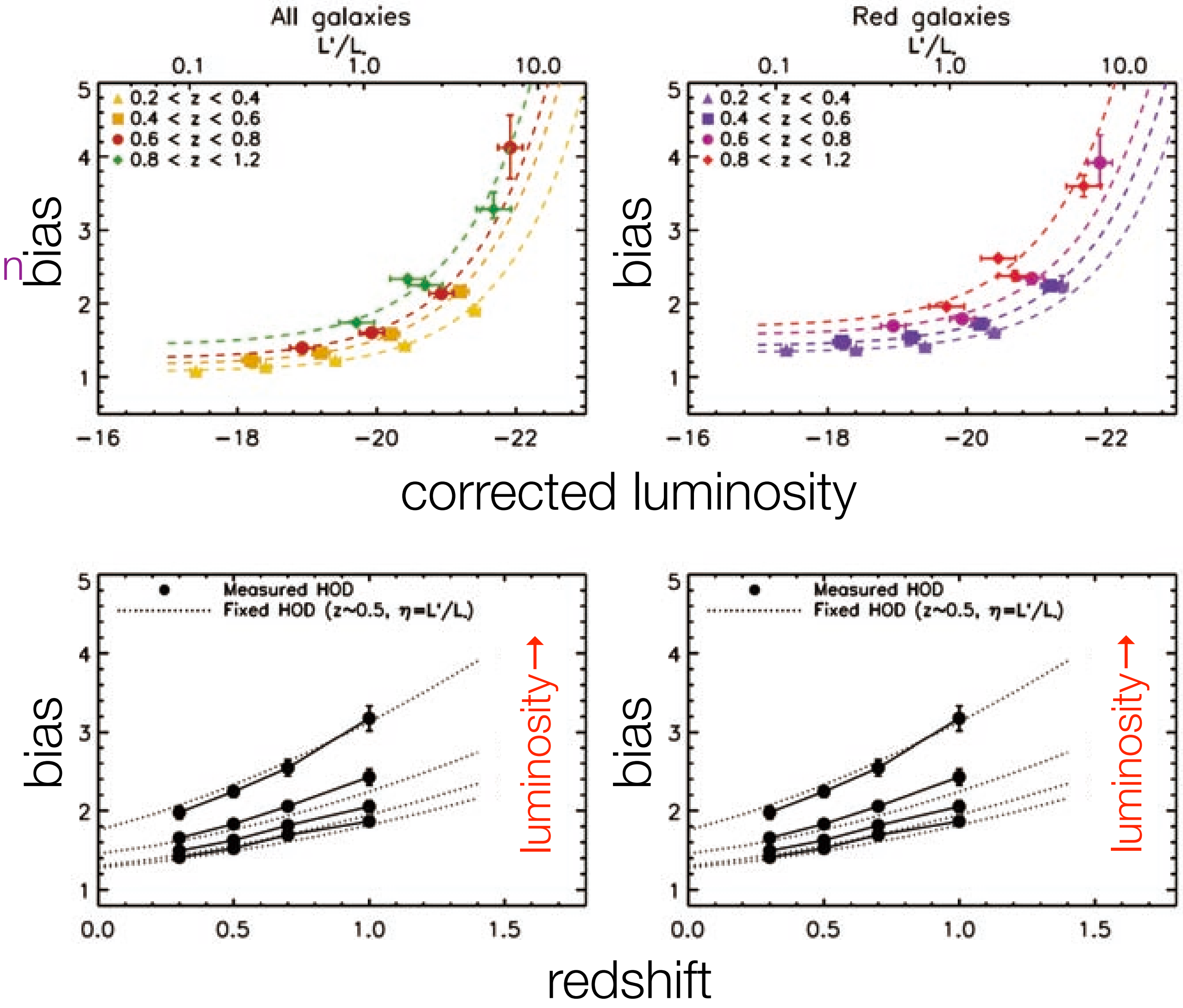
HOD [=f(z)?]

- keeping HOD fixed (z~0.5 values), we extrapolated f_{sat} in luminosity and redshift
- no significant departures with local Universe measurements

→ consistent with HOD(z) ~ cst. Minor role of galaxy mergers?

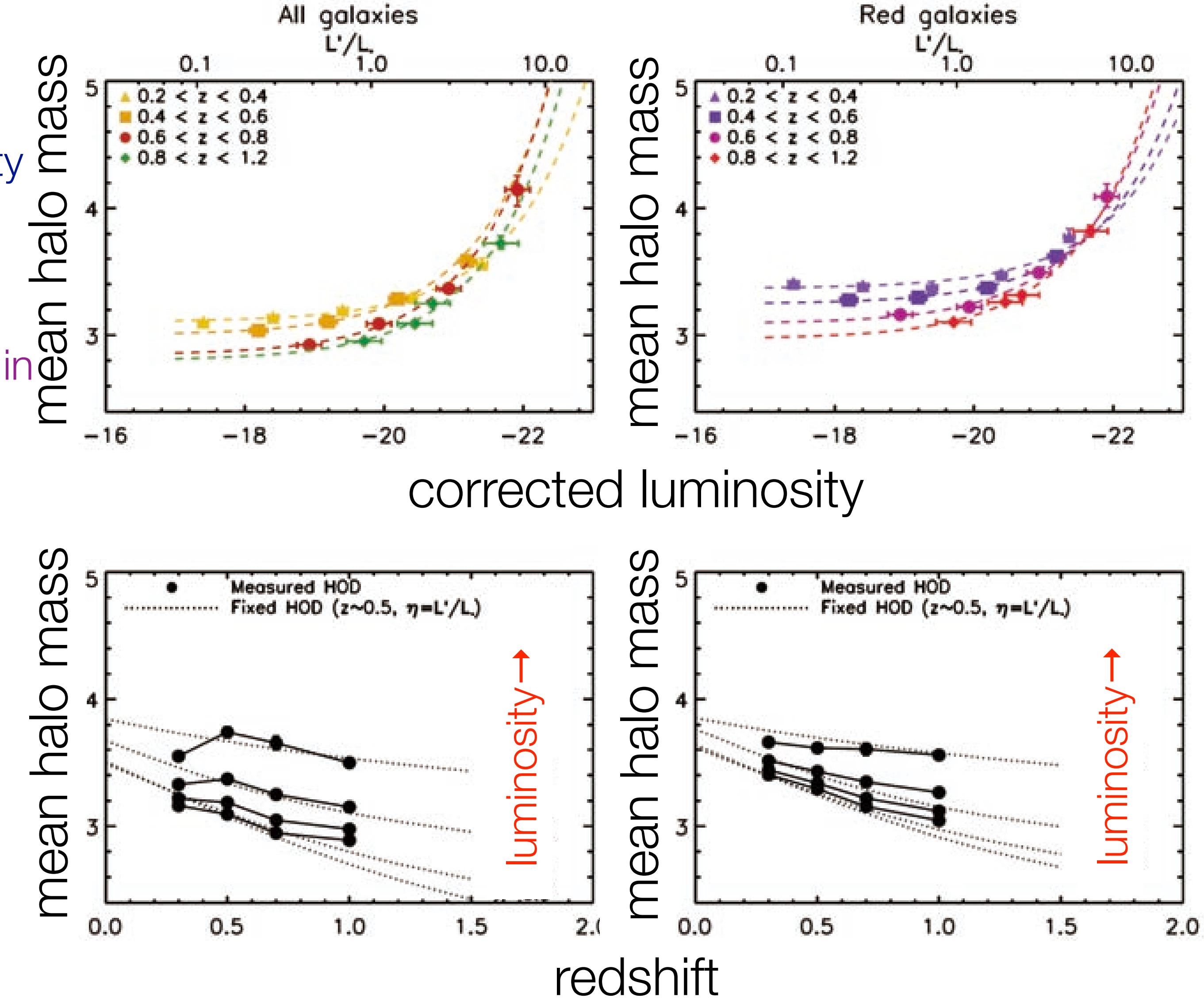
Galaxy bias

- bias increases with luminosity and redshift
- brighter galaxies reside in more massive haloes
- bright samples depart from “HOD (z) = cst” model



Mean halo mass

- mean halo mass increases with luminosity but decreases with redshift
- brighter galaxies reside in more massive haloes
- bright samples strongly depart from “HOD (z) = cst” model



V. Conclusions

Conclusions

- understanding galaxy formation and evolution is challenging
- the HOD formalism, a powerful combination of the halo model and simple assumptions, brings valuable hints on the relationship between galaxy and dark matter
- the CFHT Legacy Survey is a unique combination of depth, area and image quality
- we measured the galaxy clustering using advanced tools in the CFHTLS Wide and checked that no systematic would dominate our error budget
- when applying the HOD model to the CFHTLS, we were able to derive precise constraints on galaxy evolution

Conclusions

- for the full sample, $M_{h,\ peak}$ = 4.5 10^11 M_{sun} and moves towards higher halo masses at higher redshifts, suggesting that the bulk of star-formation activity migrates from higher halo mass at high redshifts (z~1.2) to lower halo mass haloes at lower redshifts (z~0)

- red galaxies do not evolve significantly but experience passive evolution

- for galaxies in haloes < 10^12 M_{sun}, we observed a increase in satellite fraction of about 2, which is consistent with a pictures where galaxy mergers do not play a significant role for galaxy evolution

- an important step further is to better understand the physical processes responsible for the evolution of galaxies

- future observations with accurate stellar masses and a better model for blue galaxy clustering will be necessary to complete this study

Additional content

Abstract

New insights on galaxy evolution since z ~ 1.2 from the CFHT Legacy Survey

Abstract: In the last few years, it is has become increasingly apparent that the mass of dark matter haloes in which galaxies reside is a key factor in regulating their formation and evolution. It is now evident that galaxies in low- and high-mass haloes experience very distinct fate. In this presentation, I will first explain why studying the relationship between stellar mass and halo mass brings valuable clues about physical processes involved in galaxy evolution. In the context of the halo model, the simple - but powerful - assumption that the number of galaxies only depends on halo mass, the halo occupation distribution (HOD) model, leads to an accurate analytic prediction of the galaxy distribution. Reciprocally, interpreting galaxy clustering using the HOD model allows to make a direct comparison between galaxy properties and halo mass. By using accurate galaxy clustering measurements over 133 deg2 of the "Wide" component of the Canada-France-Hawaii Telescope Legacy Survey (CFHTLS), we performed a detailed investigation of the changing relationship between galaxies and the dark matter haloes they inhabit from z ~ 1.2. I will then pursue my talk with a presentation of this unique data set combining depth, large area and high image quality, and I will finally present our results and their implications for galaxy evolution.

Photometric redshifts

Weaknesses of spectroscopic redshifts:

- requires hours of observations
- limited to i<24
- fails in the “redshift desert”

Photometric redshift is the most efficient way to measure numerous redshifts

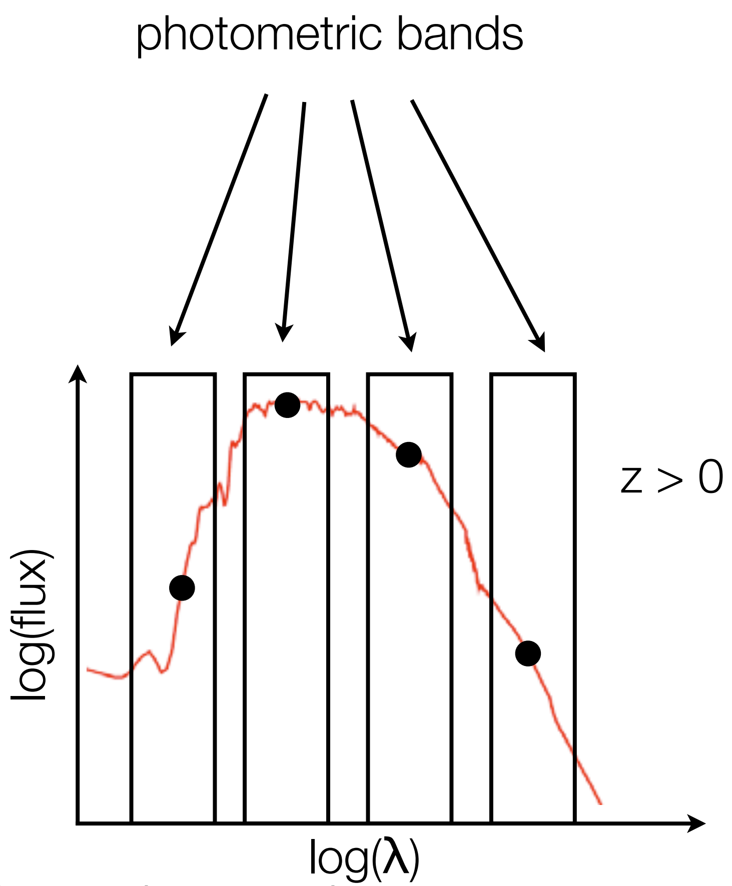
- relies on global spectral features
- efficient
- statistical purpose only

Weaknesses of photometric redshifts:

- suffers from large uncertainties
- and important degeneracies

Several methods are currently used:

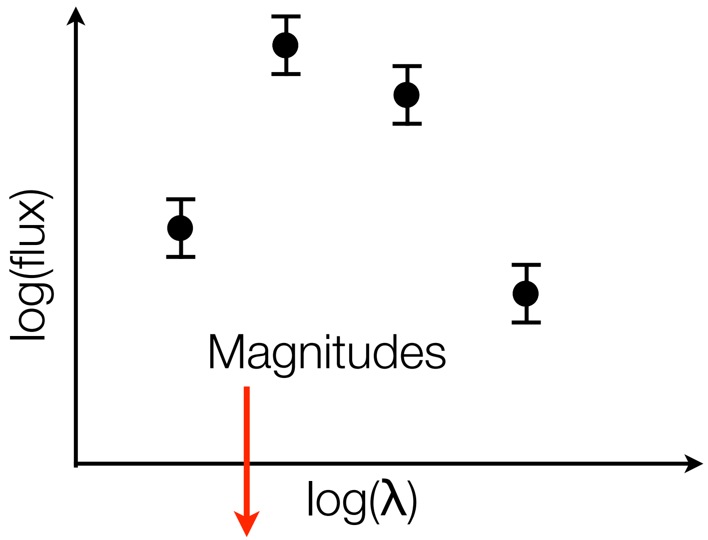
1. Artificial Neural Networks
2. Template fitting
3. Bayesian methods

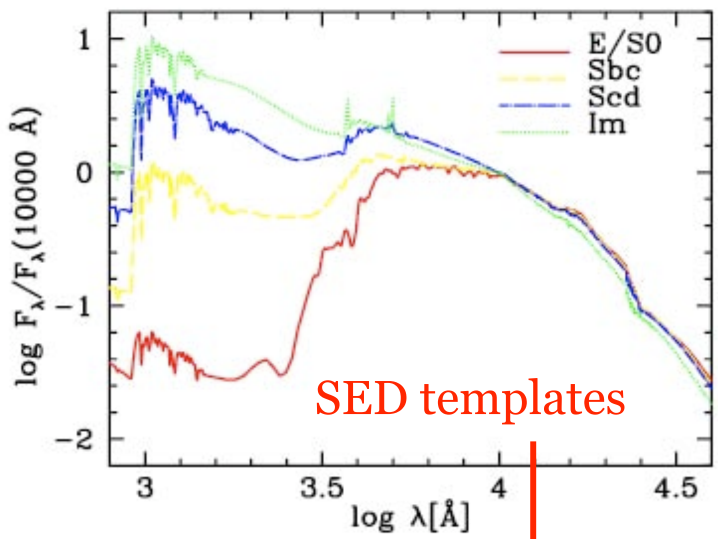


Some photometric surveys:
HDF, SDSS, CFHTLS, COSMOS
COMBO 17, GOODS, GEMS, etc.

Template fitting

- For e.g. : Hyper-z (Bolzonella et al. 2000), Le Phare (S. Arnouts & O. Ilbert)
- Uses SED templates and minimizes the χ^2 to find the redshift





Coleman, Wu & Weedman (1980)

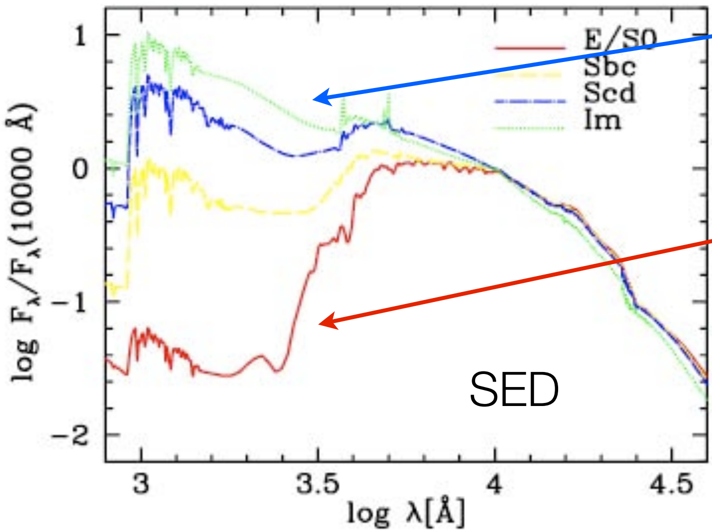
$$\chi^2(z,T,A) = \sum_{f=1}^{N_f} \left[\frac{F^f_{obs} - A \times F^f_{pred}(z,T)}{\sigma^f_{obs}} \right]^2$$

predicted magnitudes

Best χ^2 gives photo-z, model and normalization and allows to extract more physics (but strong dependence to templates)

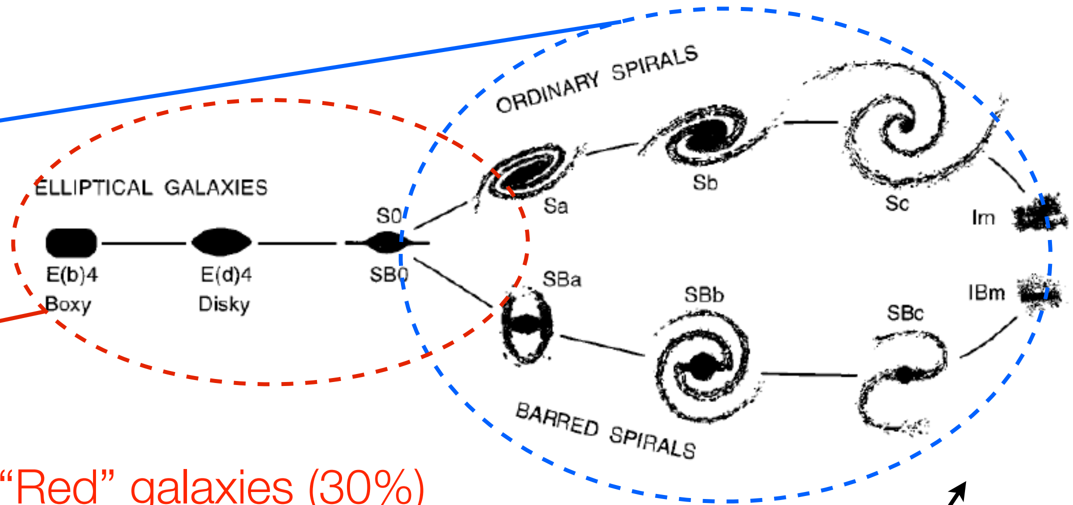
2. Template fitting: galaxy populations

Hubble's sequence improved by
Kormendy & bender (1996)



SED

(Hyper z's user
manual)



"Red" galaxies (30%)

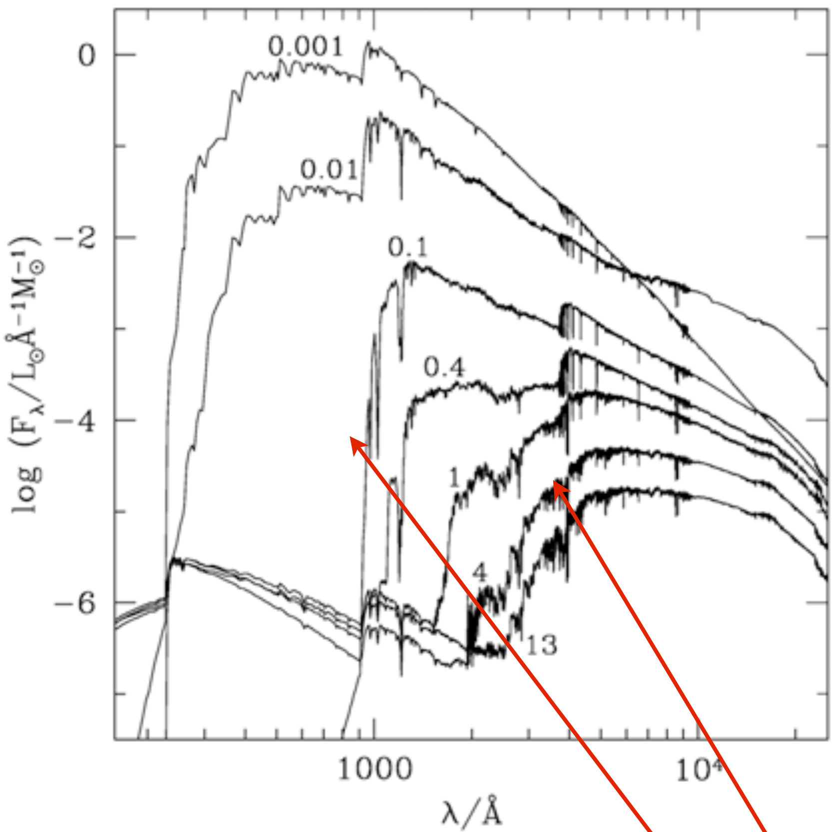
"Blue" galaxies (70%)

Irregular galaxies

+ morphology? (not resolved from ground-based telescopes)

Stellar evolution

Evolution of a stellar population



Bruzual & Charlot (2003)

Balmer break

Lyman break

Red galaxies show older stars and stronger Breaks -> leads to best photo-z estimates

Blue galaxies form stars and contain a lot of gas -> needs to take the internal extinction into account

Degeneracies come from

- confusion z/type
- z evolution
- SED uncertainties

Bayesian approach

With observed or synthetic models of SED + Bayes' theorem: $P(z|C) = \dfrac{P(C|z)P(z)}{P(C)}$

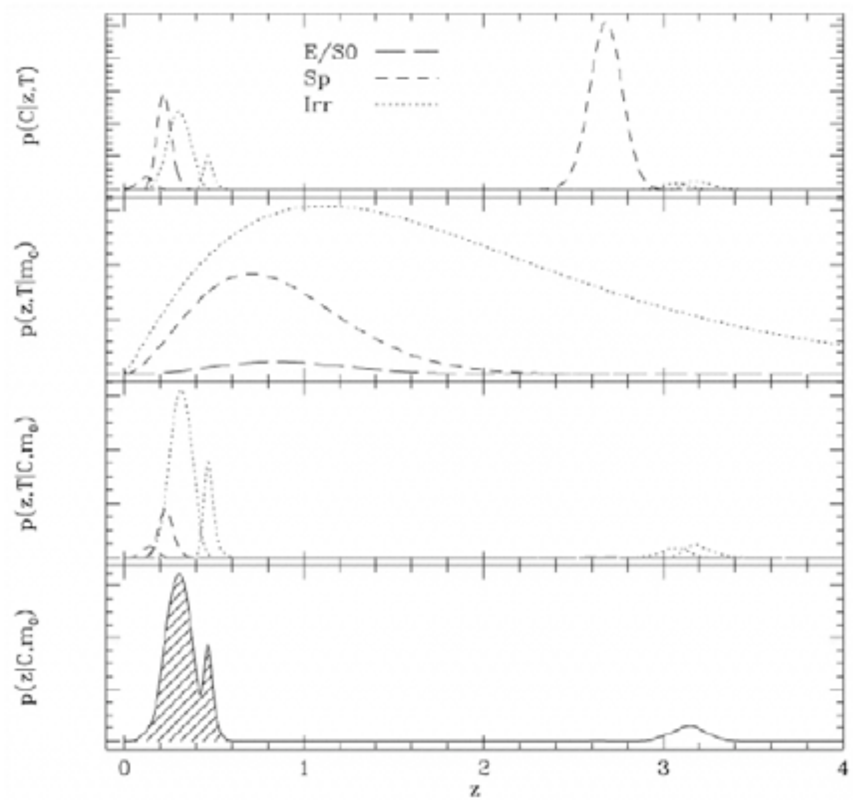
$\mathcal{L}(z) = \exp(-\frac{\chi^2}{2})$ likelihood →

n(z,T,m) derived from spectro-z's prior →

posterior →

Probability Distribution Function ≡ PDF(z)

marginalization →



BPz, Benitez (2000)

Very useful to break degeneracies but beware of abusive use of priors !!

Le Phare

PHotometric Analysis for Redshift Estimations (S. Arnouts & O. Ilbert)

Code for photometric redshift estimation based on template fitting and bayesian approach

magnitudes
+ errors

$\chi^2(z, T, A)$

SED libraries:
galaxies/stars/QSO
+ filter set
+ extinction, IGM

(Lots of assumptions here)

- Best redshift, second solution
- Type, E(B-V)
- Returns a χ^2 for galaxy, star and QSO library
- Uncertainties, PDF(z)
- Absolute magnitudes
- Physical parameters: mass,mean age,SFR,etc.

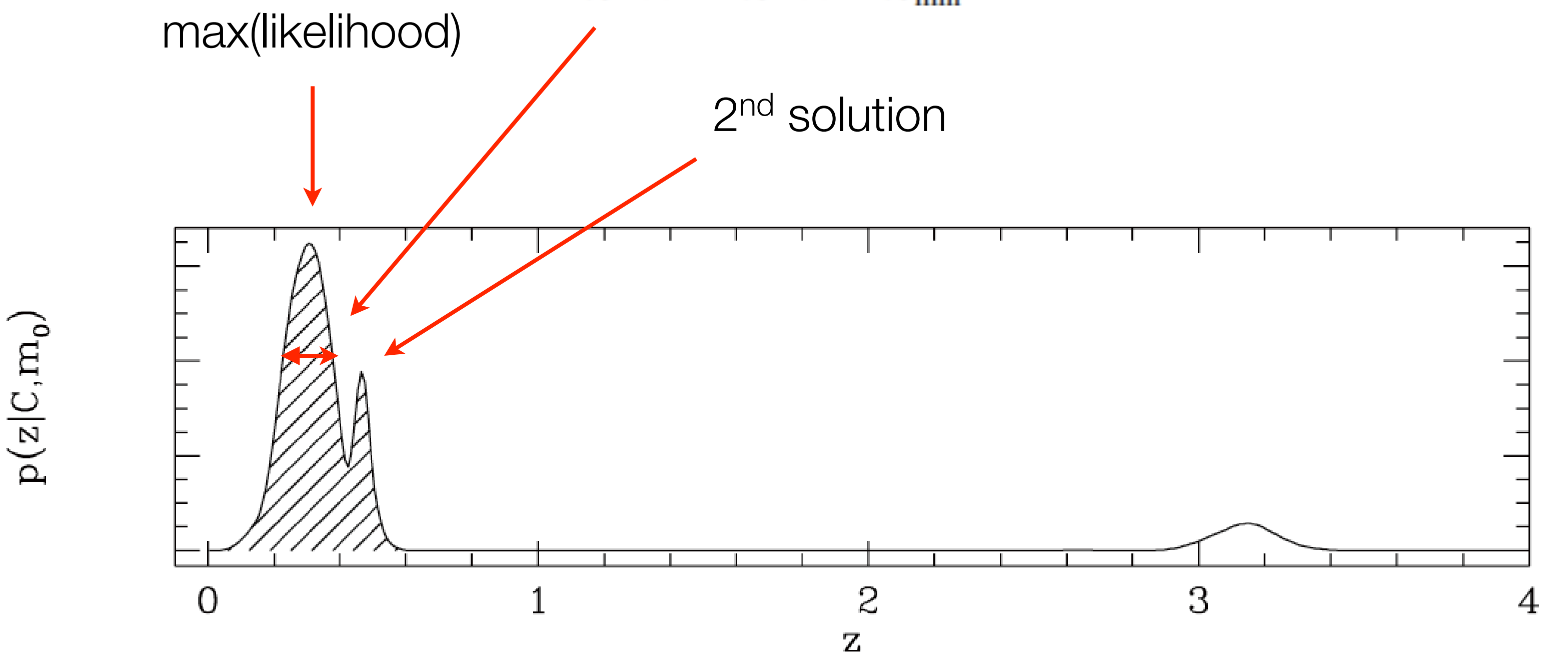
advanced features

But: can be time consuming...

Probability Distribution Function

68% confidence interval:

$$\Delta\chi^2(z) = \chi^2(z) - \chi^2_{\min} = 1$$



- redshift estimate: max(likelihood) or median(PDF)
- error estimate: 68% confidence limits
- odds parameter (Benitez et al. 2000)= “peakiness” of the PDF
- second peak information provided
- real PDF? Probably not...
- How to use all the information?

Photo-z errors

Comparison with spectro-z's
= "real errors"

photo-z dispersion, **σ**, robust estimator of the standard deviation:

$$\sigma_{\Delta z/(1+z_s)} = 1.48 \times \mathrm{median}(|\Delta z|/(1+z_s))$$

outlier rate, **η** (number of catastrophic errors), as the number of objects with:

$$|\Delta z| \geq 0.15 \times (1 + z_s)$$

Computed by Le Phare
= "estimated errors"

From the PDF(z), 68% confidence limits are defined as:

$$\Delta\chi^2(z) = \chi^2(z) - \chi^2_{\mathrm{min}} = 1$$

and our error estimate becomes:

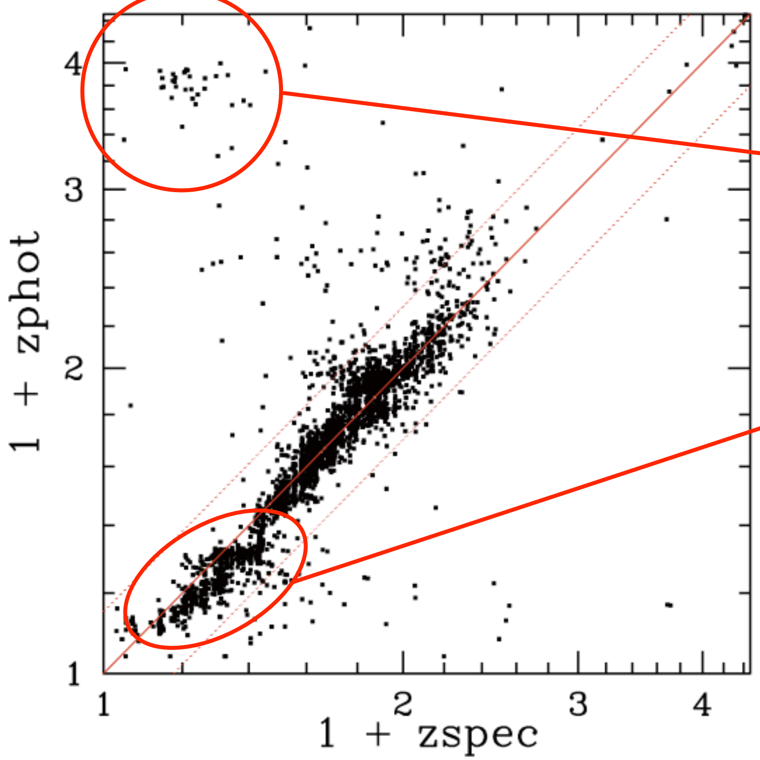
$$\sigma_{z_p} = \frac{|z_{\mathrm{left}}(68\%) - z_{\mathrm{right}}(68\%)|}{2}$$

$$\sigma_{\Delta z/(1+z_s)} = \sigma_{z_p}$$

Results with LePhare

Raw results with template fitting
(no spectroscopic calibration)

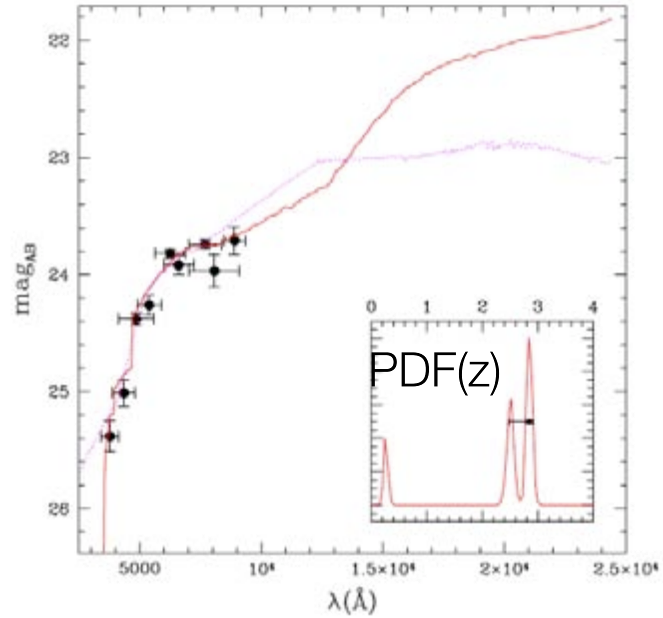
Ilbert et al. (2006), $i'_{AB} < 22.5$



CFHTLS deep vs VVDS deep (le Fevre et al. 2005)

Catastrophic errors
= “Outliers”

Systematic trends



Idea: calibrate photo-z’s with spectro-z’s.
Ilbert et al. (2006):
1. correction of systematic offsets
2. template optimization
3. galaxy distribution priors

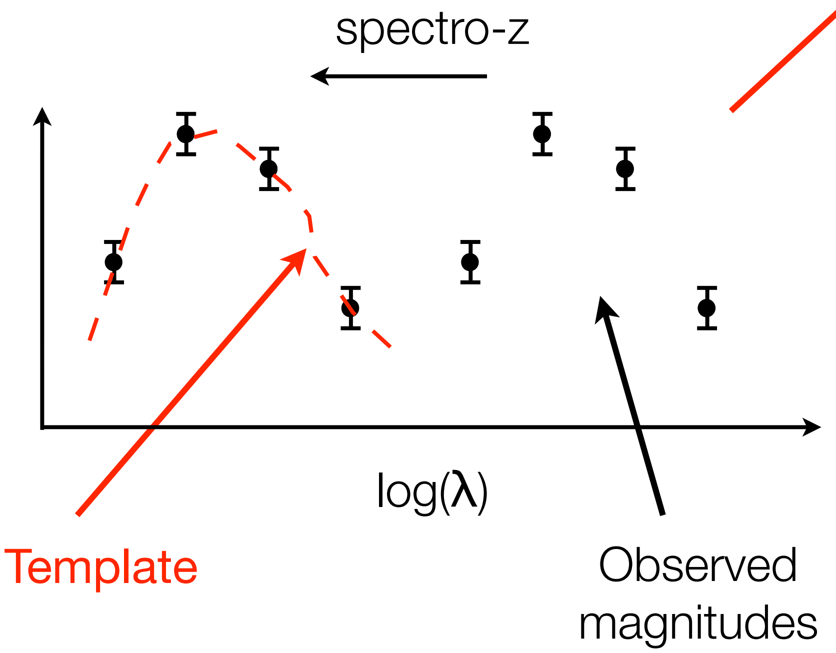
1. Correction of systematic offsets

Mainly comes from poor filter knowledge and calibration uncertainties

Using spectro-z's, one can put back the photometry in the rest frame and compute the sum:

$$\psi^2 = \sum_{N_{gal}} \left[\left(A \times F^f_{pred} - F^f_{obs} + s^f \right) / \sigma^f_{obs} \right]^2$$

All objects with spectro-z



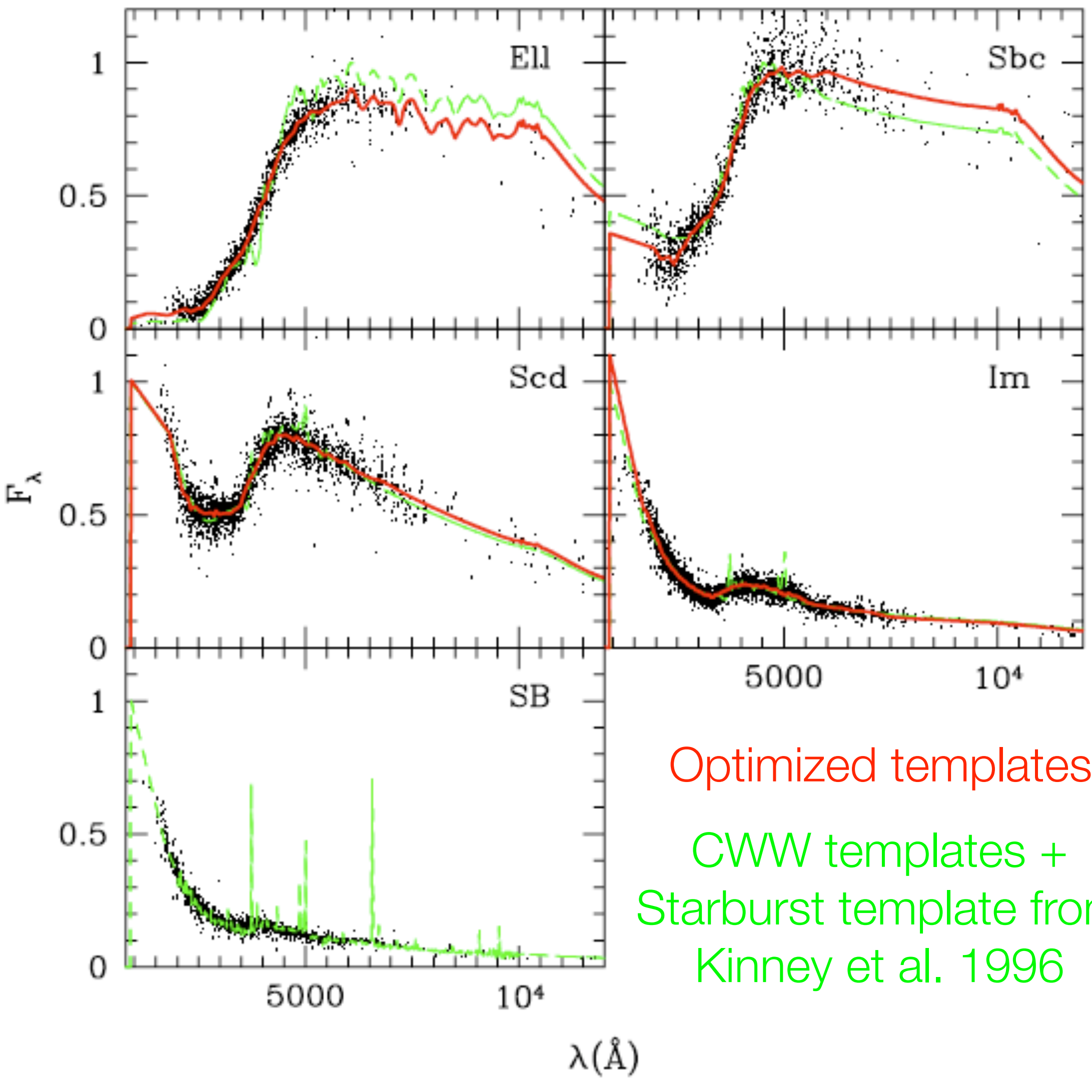
we notice photometric systematic offsets but with small dependence on magnitude and SED type:

filter	CWW $i'_{AB} < 21.5$	CWW $i'_{AB} < 22.5$	CWW $i'_{AB} < 23.5$	PEGASE $i'_{AB} < 22.5$
B	+0.068	+0.071	+0.067	+0.078
V	–0.037	–0.043	–0.046	–0.038
R	+0.089	+0.090	+0.093	+0.102
I	–0.002	+0.001	+0.004	–0.008
$u*$	+0.020	+0.019	+0.008	+0.076
g'	–0.071	–0.079	–0.080	–0.080
r'	+0.002	–0.002	–0.005	+0.000
i'	+0.000	+0.000	+0.000	+0.000
z'	–0.006	–0.008	–0.006	–0.027

Ilbert et al. (2006) with VVDS

2. Template optimisation

Ilbert et al. (2006)



Optimized templates

CWW templates + Starburst template from Kinney et al. 1996

Problem: original templates (CWW) have been observed with different instruments and on a small sample of galaxies

might not represent real SEDs and create systematic trends

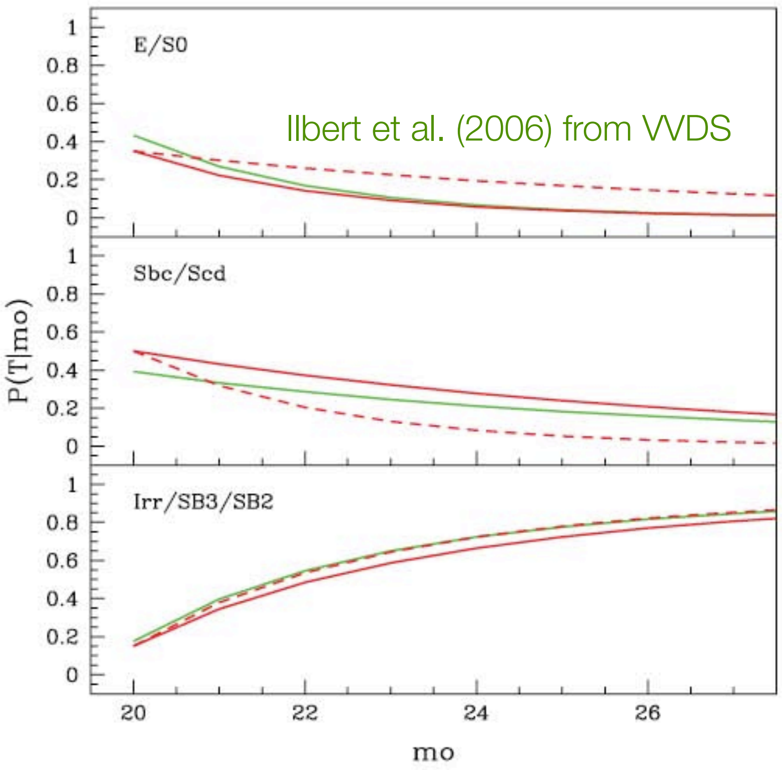
solution: adapting the templates to the data. Again with spectro-z, put back the photometry in the rest frame

3. Galaxy distribution priors

Priors constructed with spectro-z galaxy distributions (here VVDS)

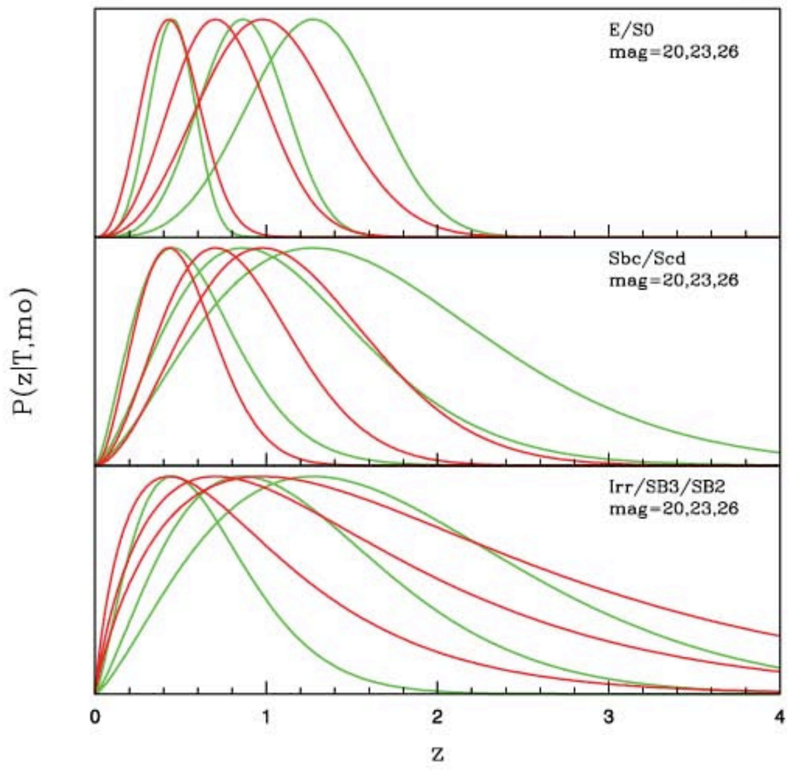
SED type probability distribution

$$p(T|i'_{\mathrm{AB}}) = f_t e^{-k_t(i'_{\mathrm{AB}}-20)}$$



redshift distribution per type

$$p(z|T, i'_{\mathrm{AB}}) \propto z^{\alpha_t} \exp\left(-\left[\frac{z}{z_{0t} + k_{mt}(i'_{\mathrm{AB}} - 20)}\right]^{\alpha_t}\right)$$



Marganiner & Wittman (2008)

$$p(z, T|i'_{\mathrm{AB}}) = p(T|i'_{\mathrm{AB}})p(z|T, i'_{\mathrm{AB}})$$

Improved results with LePhare

CFHTLS photo-z's vs VVDS spectro-z's, i'$_{AB}$ < 24

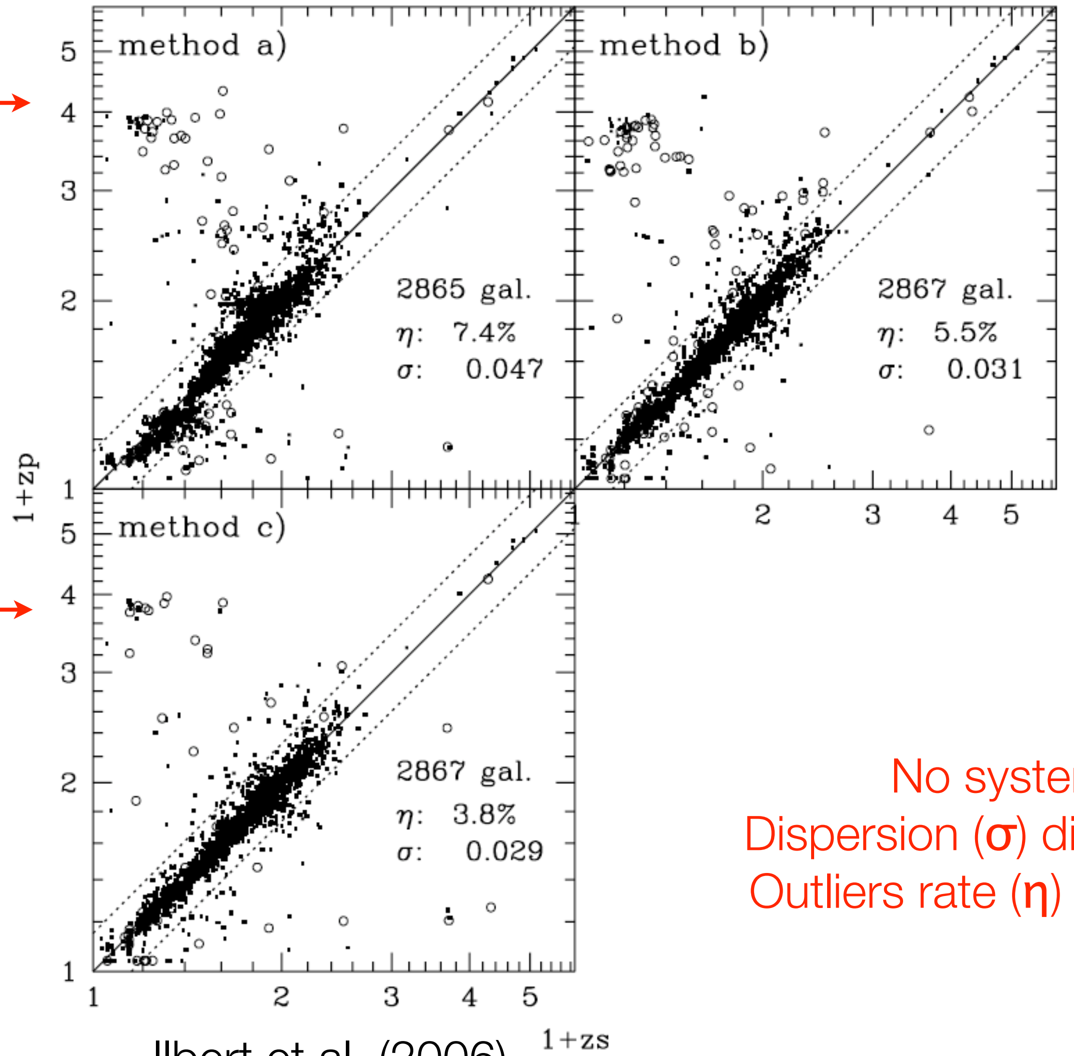
Raw template fitting - no improvement

Offset correction + template optimization

Removes systematics

priors

Reduces catastrophic errors



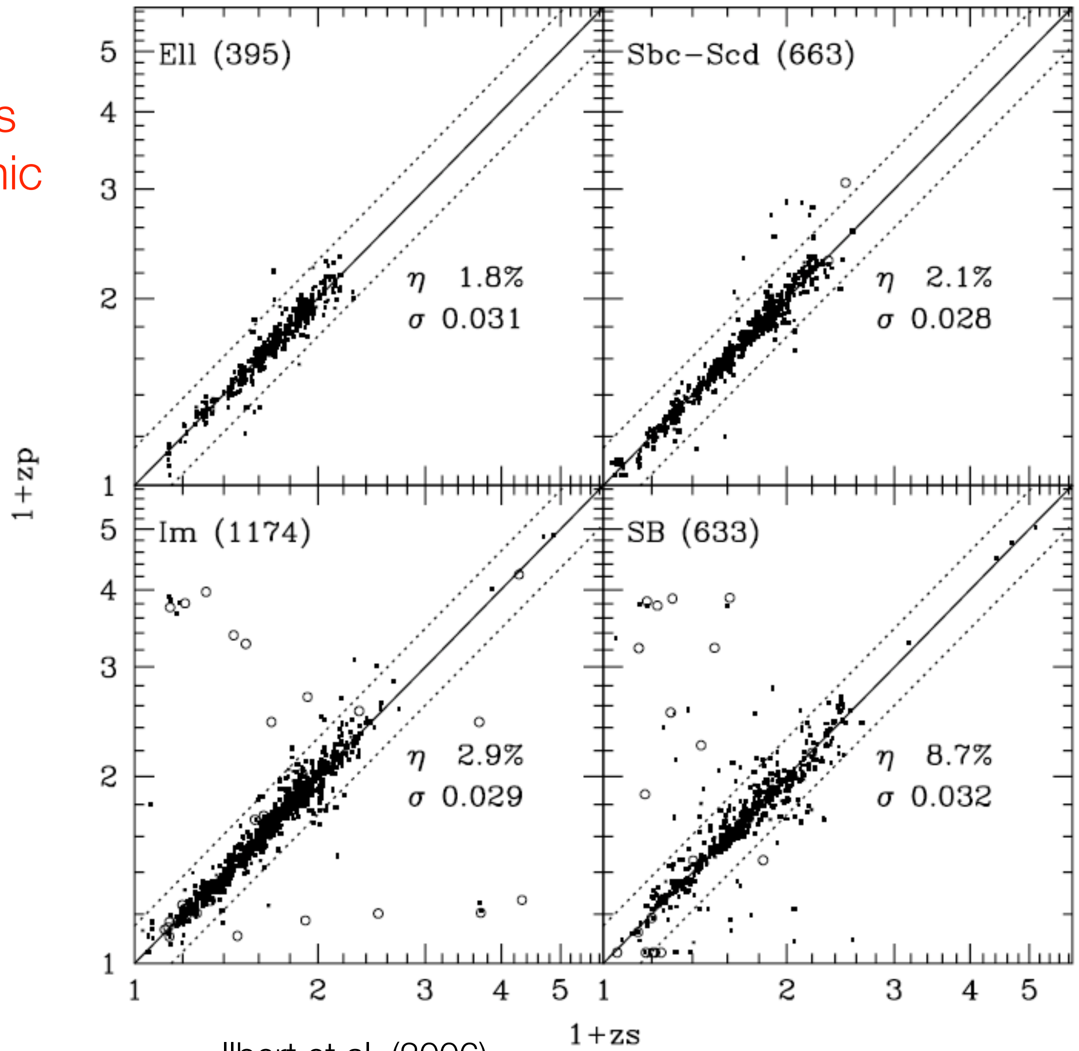
Ilbert et al. (2006)

No systematic
Dispersion (σ) divided by 1.5
Outliers rate (η) divided by 2

Improved results with LePhare

CFHTLS photo-z's vs VVDS spectro-z's, i'$_{AB}$ < 24

Elliptic/red galaxies show no catastrophic errors

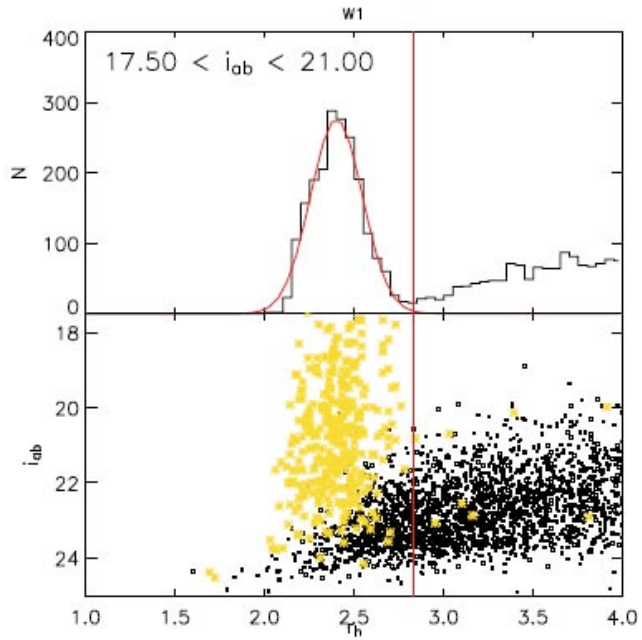


Ilbert et al. (2006)

Starburst galaxies need more improvement (emission lines?)

Star/galaxy separation

stars can be an important source of contamination: up to 50% Size + magnitude is sometimes not enough to separate stars from galaxies

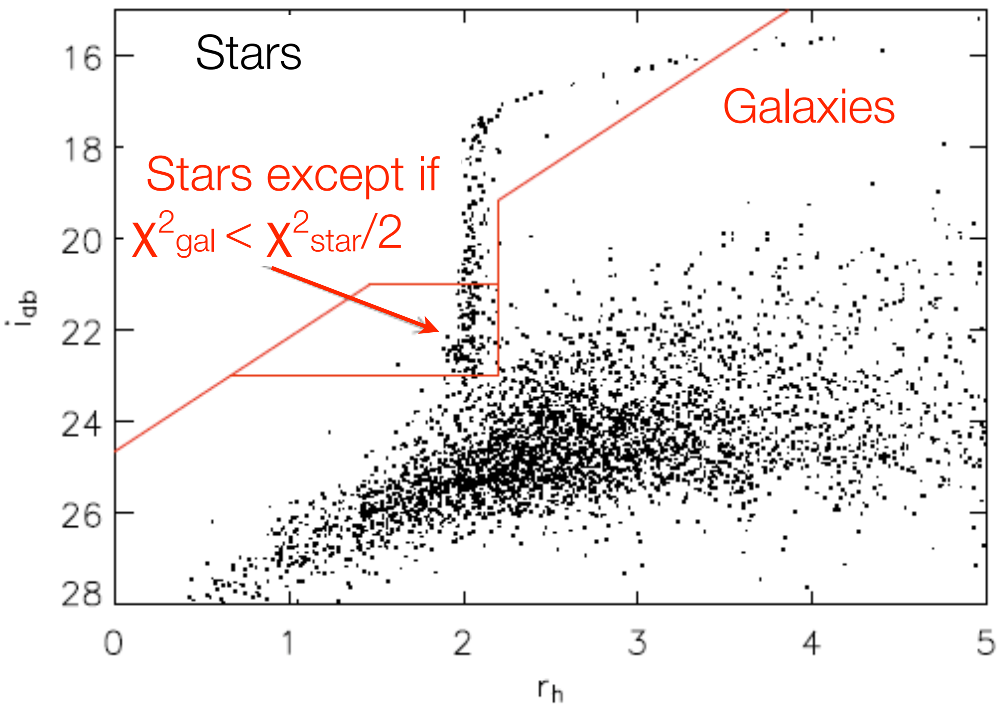


Idea: adding the colour information to improve the selection:

Star <=> $r_h < r_{hlimit}$ AND $\chi^2_{gal} > \chi^2_{star}/2$

Galaxy <=> $r_h > r_{hlimit}$ OR $\chi^2_{gal} < \chi^2_{star}/2$

calibrated with VVDS spectra



Results: in W4, star contamination reduced from 50% to 8%, keeping 99% of galaxies

Method used in the source selection for VIPERS