The dark side of supernova research

- Type I: dark energy
- Type II: dark matter
  - Type IIa: **axion**
  - Type IIb: **black holes**
  - Type IIn: sterile **neutrinos**
- How will we know the correct answer?

- Neutrino masses suggest **dark matter in the form of sterile (right-handed) neutrinos**
- **Pulsar velocities** explained by the same sterile neutrino with 2-20 keV mass (emission from a supernova is anisotropic!) Other astrophysical hints: reionization, star formation
- **X-ray bounds** and the future prospects, including Suzaku observations (together with Loewenstein, Biermann)

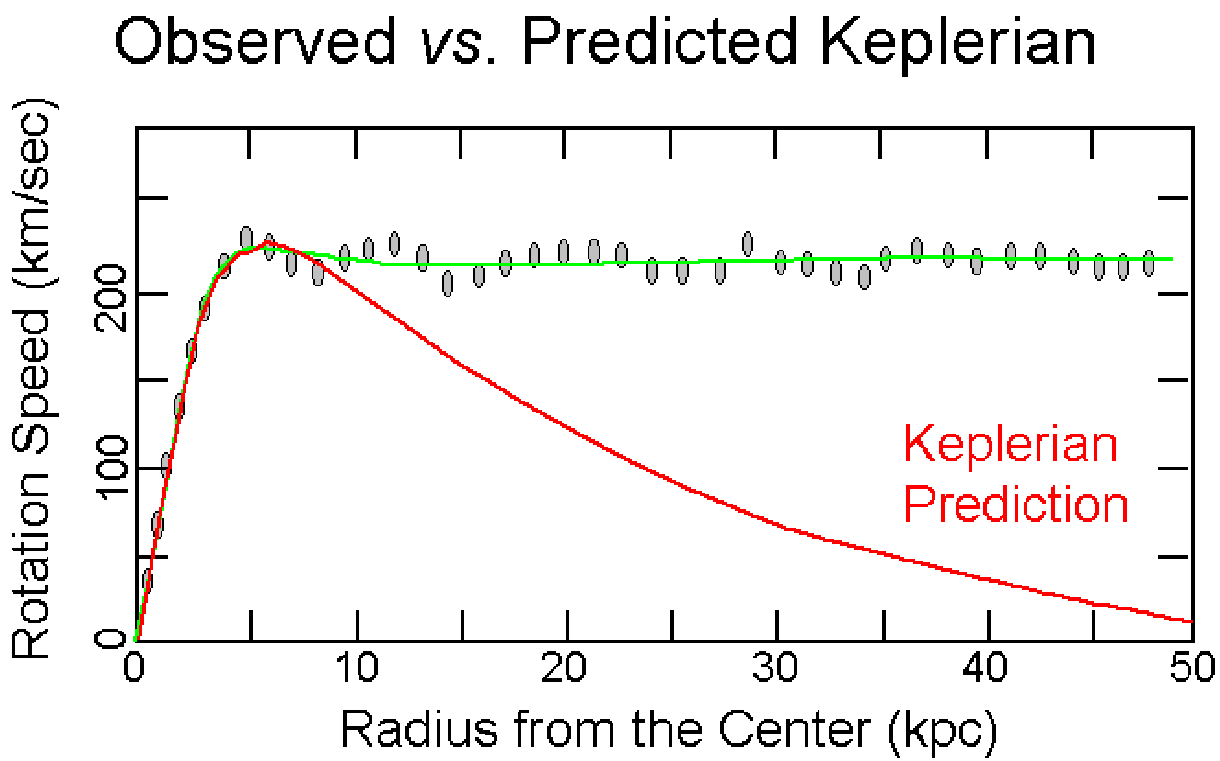
Dark matter

The only data at variance with the Standard Model

The evidence for dark matter is very strong:

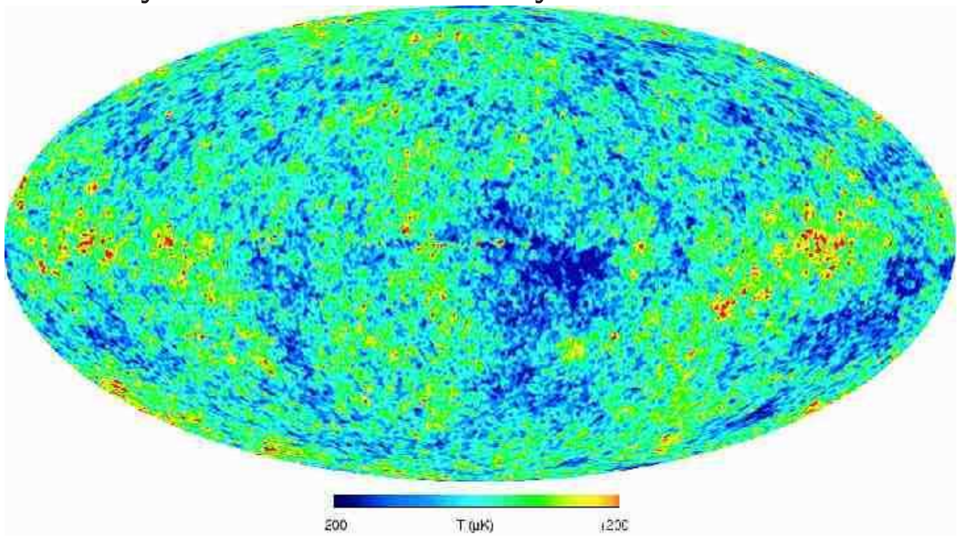
- galactic rotation curves cannot be explained by the disk alone
- cosmic microwave background radiation
- gravitational lensing of background galaxies by clusters is so strong that it requires a significant dark matter component.
- clusters are filled with hot X-ray emitting intergalactic gas (without dark matter, this gas would dissipate quickly).
- neat: 1E0657-56 shows separation of ordinary matter (gas) from dark matter

Galactic rotation curves

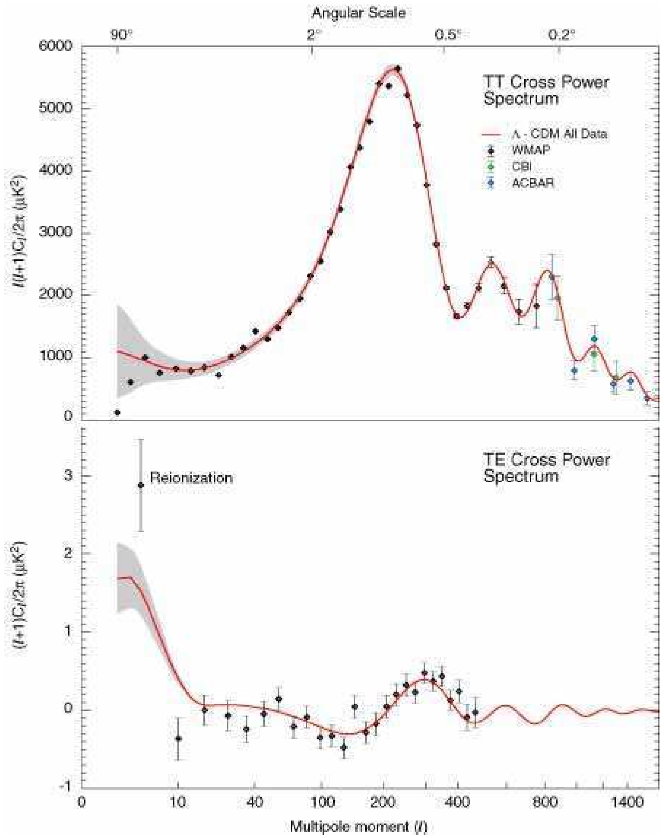


Cosmic microwave background radiation (CMBR)

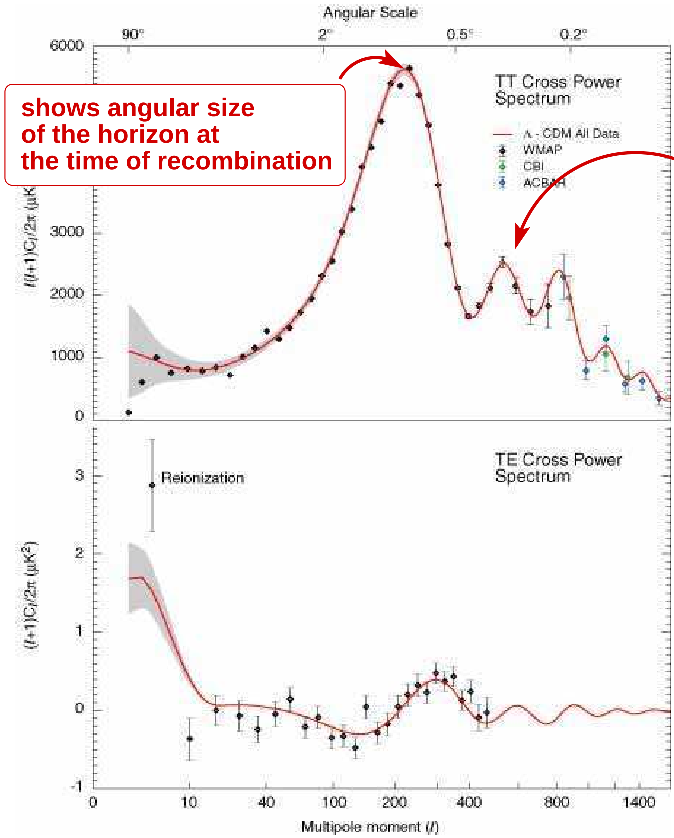
At *decoupling*, the atoms formed and the universe became transparent to radiation. Radiation emitted at that time has been red-shifted into the microwave range. Fluctuations have been measured first by COBE, and later by BOOMERANG, MAXIMA, ..., WMAP:



Power spectrum measured by WMAP



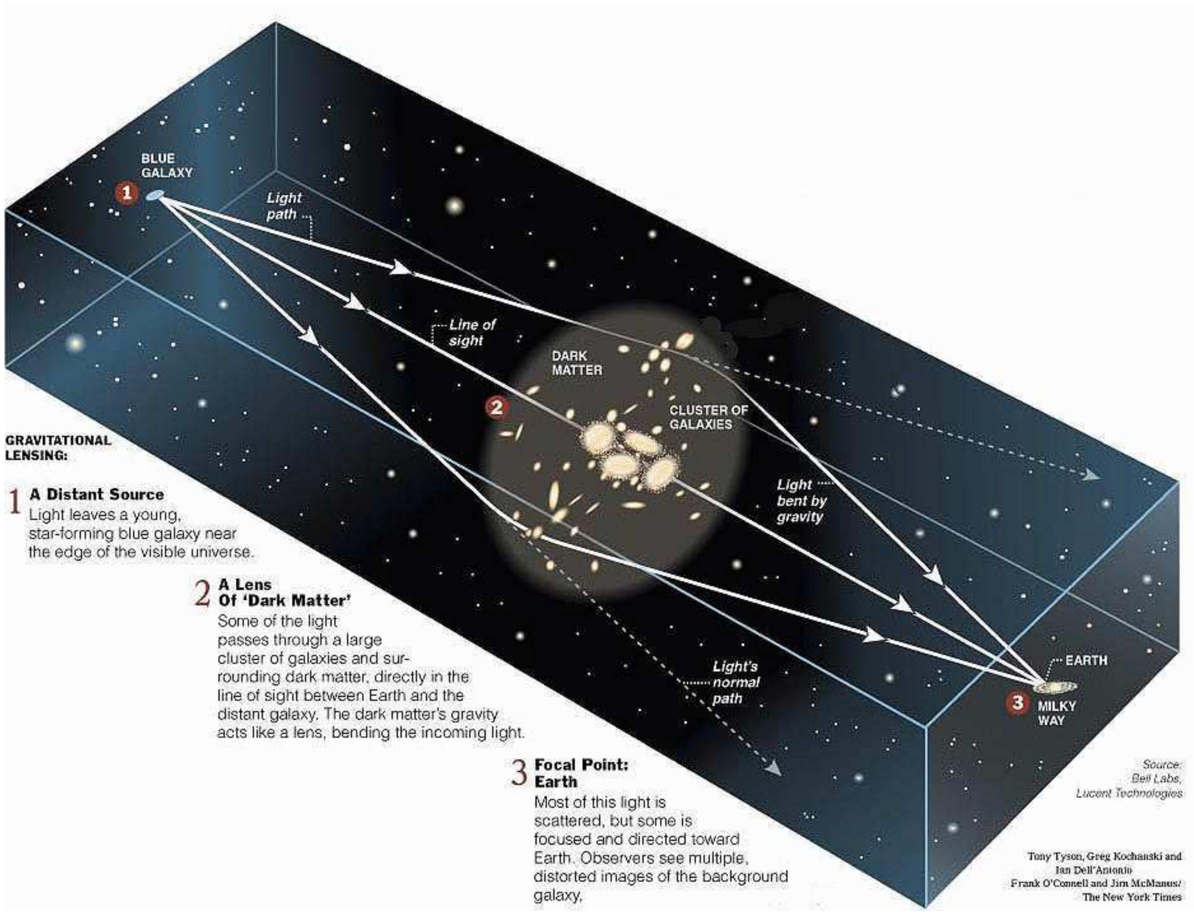
Power spectrum measured by WMAP

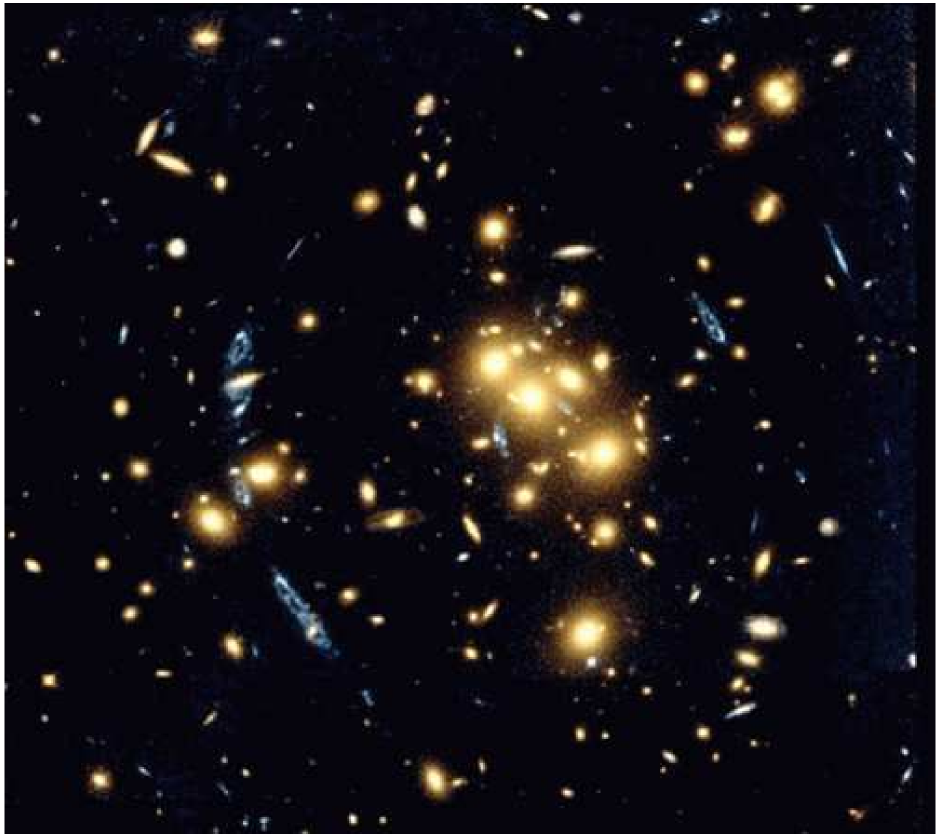


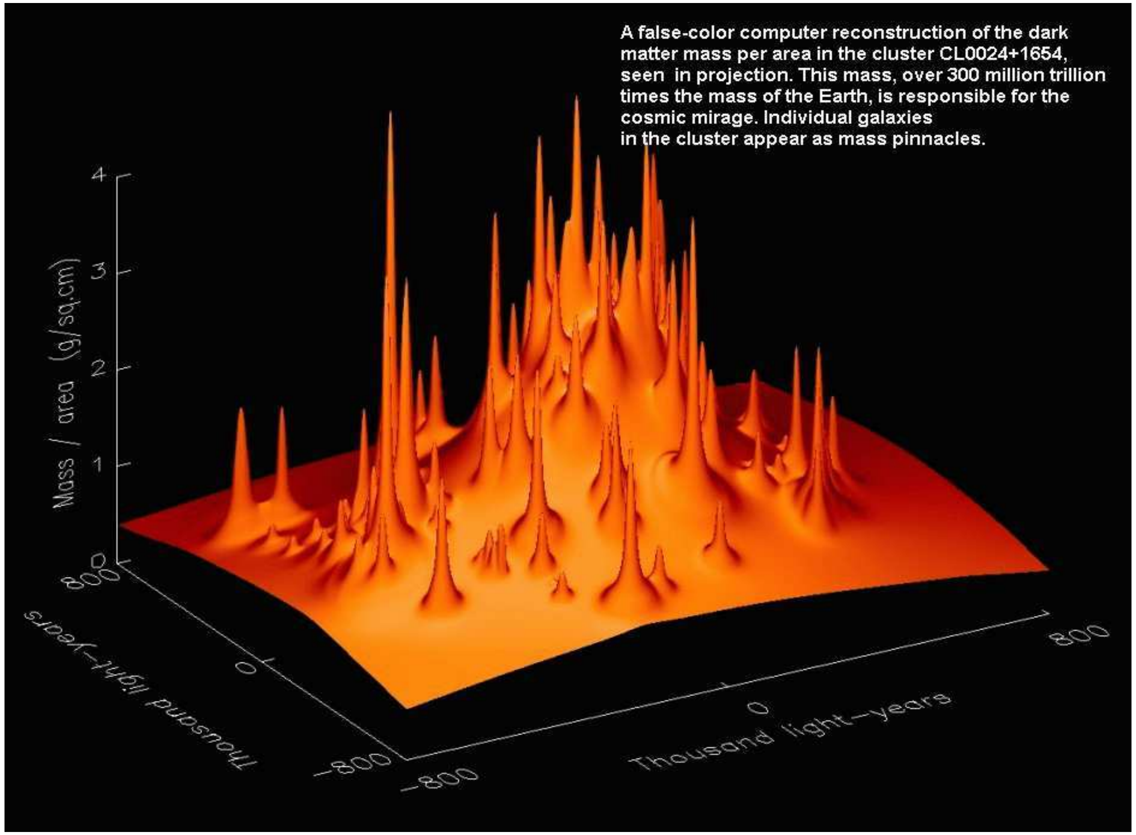
shows angular size of the horizon at the time of recombination

sensitive to the ratio of ordinary to dark matter densities

Gravitational lensing: seeing the invisible

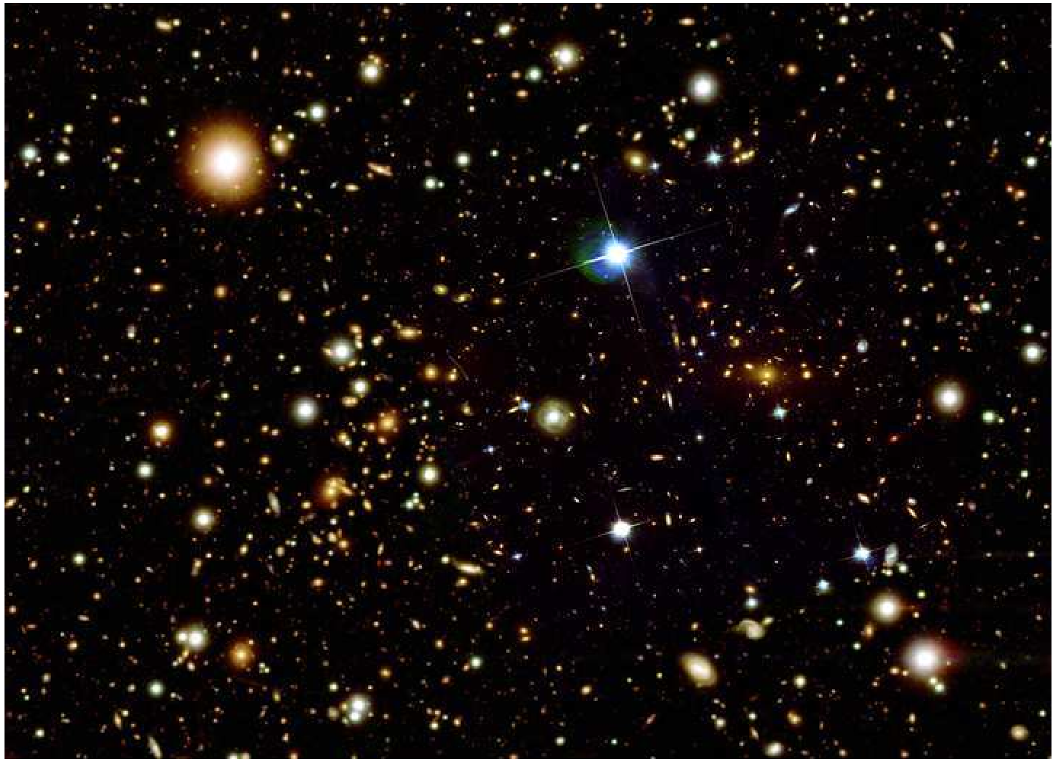




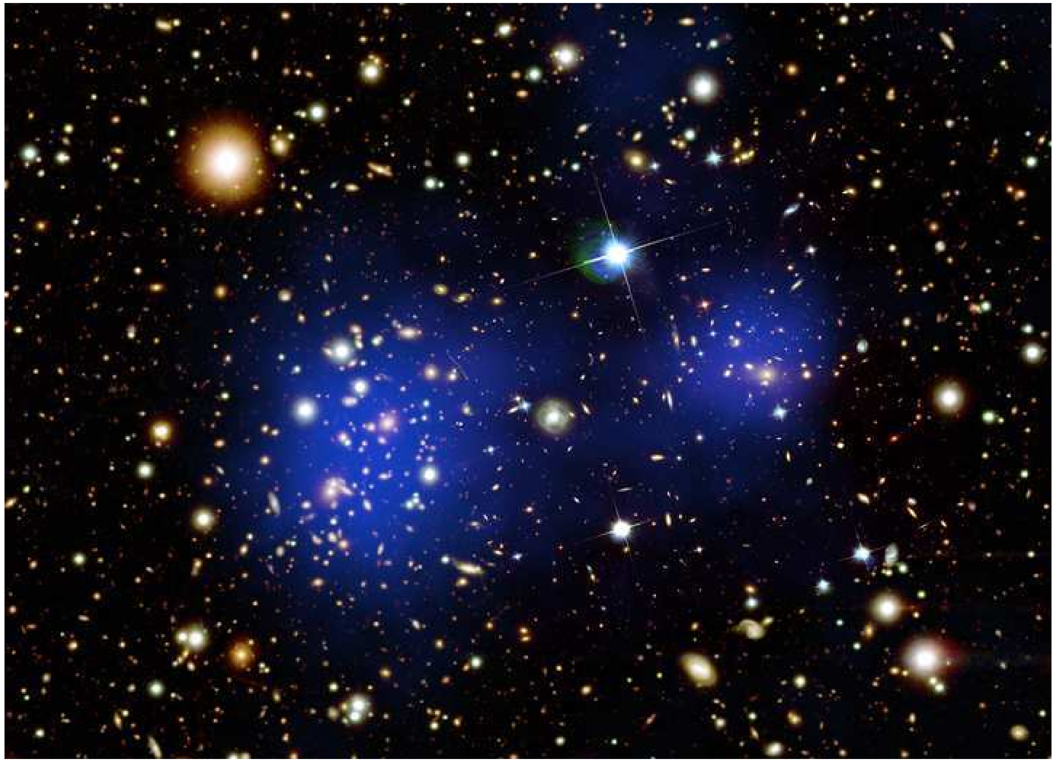


Foreground cluster CL0024+1654 produces multiple images of a blue background galaxy in the HST image (left). Mass reconstruction (right).

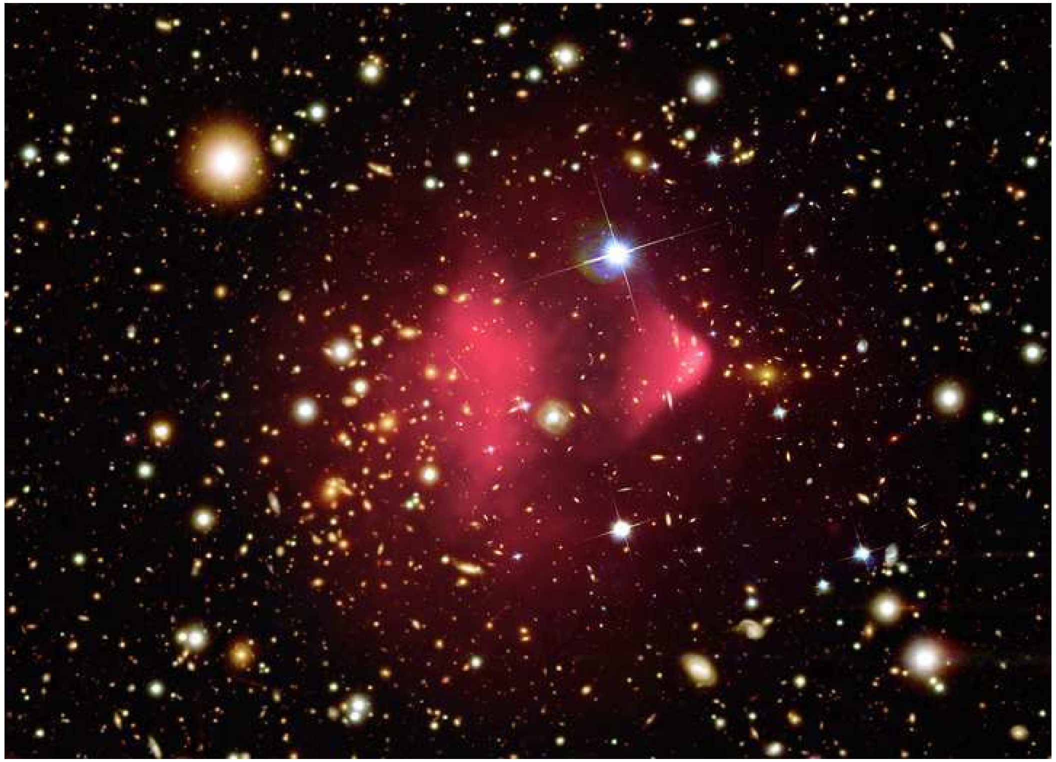
Merging clusters: optical image of 1E 0657-56



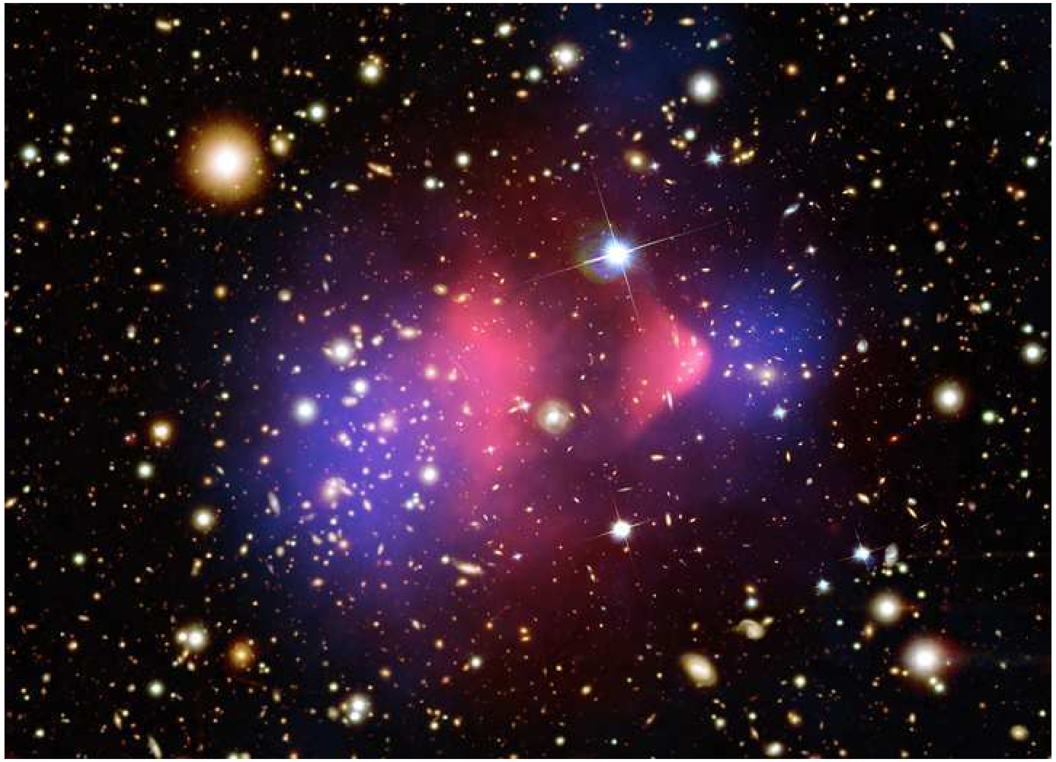
Merging clusters: grav. lensing image of 1E 0657-56



Merging clusters: Chandra x-ray image of 1E 0657-56

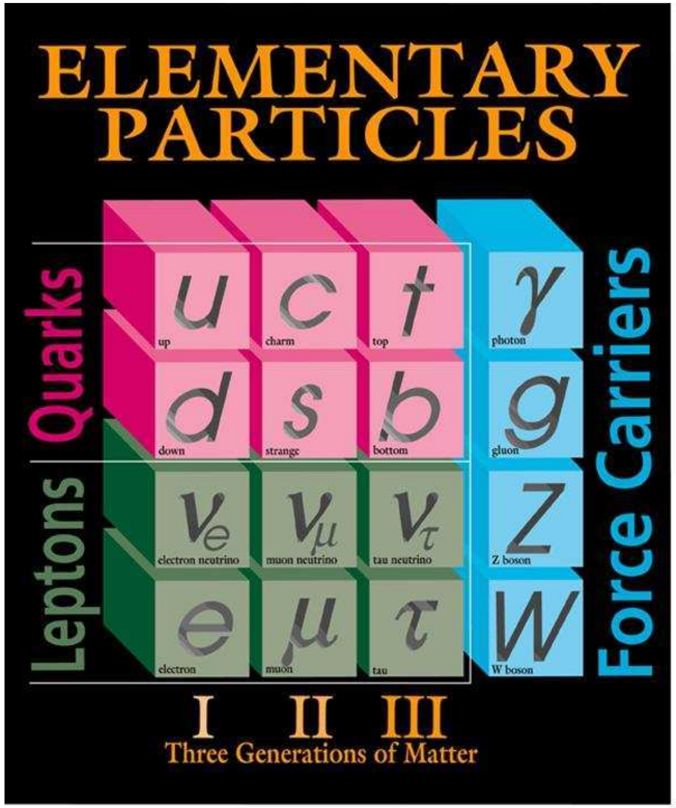


Merging clusters: image of 1E 0657-56



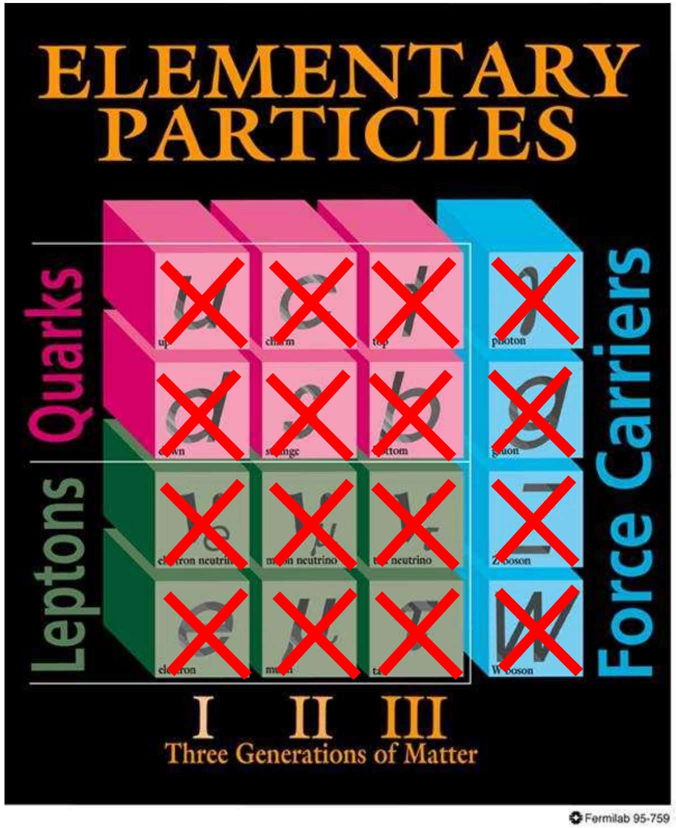
Gass, dark matter separated.

None of the known particles can be dark matter

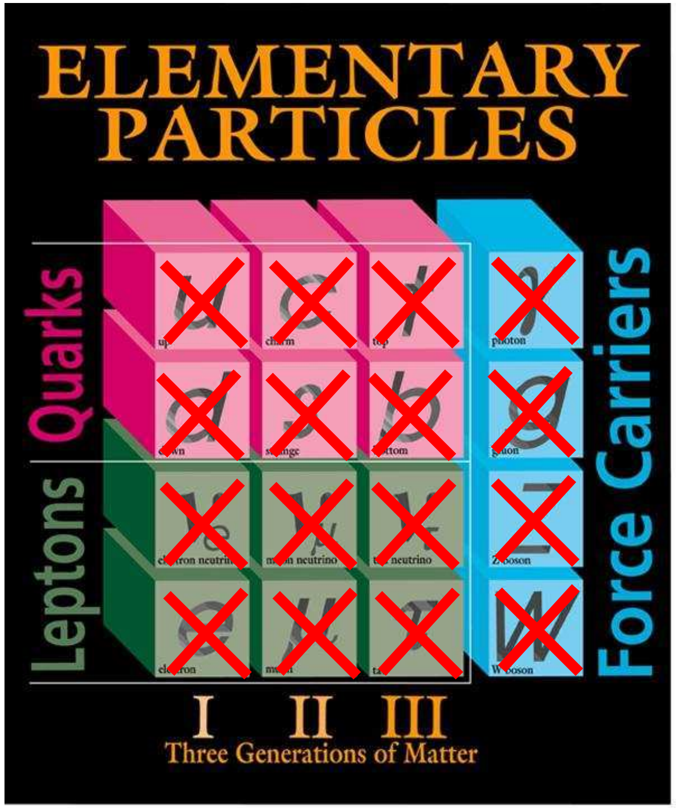


Fermilab 95-759

None of the known particles can be dark matter



Dark matter ⇒ new physics (at least one new particle)



Fermilab 95-759

The dark side: what is dark matter?

Can make guesses based on...

- ...compelling theoretical ideas
- ...simplicity
- ...observational clues

One has to guess the answer before one can make a discovery!

Neutrino masses

Neutrino mass ⇒ singlets probably exist!

The discovery of the neutrino mass is the discovery of new particles (at some scale).

Neutrino masses

Discovery of neutrino masses implies a plausible existence of right-handed (sterile) neutrinos. Most models of neutrino masses introduce sterile states

$$\{\nu_e, \nu_\mu, \nu_\tau, \nu_{s,1}, \nu_{s,2}, ..., \nu_{s,N}\}$$

and consider the following lagrangian:

$$\mathcal{L} = \mathcal{L}_{\rm SM} + \bar{\nu}_{s,a} \left(i\partial_\mu \gamma^\mu \right) \nu_{s,a} - y_{\alpha a} H \, \bar{L}_\alpha \nu_{s,a} - \frac{M_{ab}}{2} \, \bar{\nu}^c_{s,a} \nu_{s,b} + h.c. \, ,$$

where H is the Higgs boson and L_α $(\alpha = e, \mu, \tau)$ are the lepton doublets. The mass matrix:

$$M = \left(\begin{array}{cc} \tilde{m}_{3\times 3} & D_{3\times N} \\ D^T_{N\times 3} & M_{N\times N} \end{array} \right)$$

Neutrino masses

Discovery of neutrino masses implies a plausible existence of right-handed (sterile) neutrinos. Most models of neutrino masses introduce sterile states

$$\{\nu_e, \nu_\mu, \nu_\tau, \nu_{s,1}, \nu_{s,2}, ..., \nu_{s,N}\}$$

and consider the following lagrangian:

$$\mathcal{L} = \mathcal{L}_{\rm SM} + \bar{\nu}_{s,a} \left(i\partial_\mu \gamma^\mu \right) \nu_{s,a} - y_{\alpha a} H \, \bar{L}_\alpha \nu_{s,a} - \frac{M_{ab}}{2} \, \bar{\nu}^c_{s,a} \nu_{s,b} + h.c. \, ,$$

where H is the Higgs boson and L_α $(\alpha = e, \mu, \tau)$ are the lepton doublets. The mass matrix:

$$M = \begin{pmatrix} 0 & D_{3\times N} \\ D^T_{N\times 3} & M_{N\times N} \end{pmatrix}$$

What is the *natural* scale of M?

Seesaw mechanism [Yanagida et al.]

In the Standard Model, the matrix D arises from the Higgs mechanism:

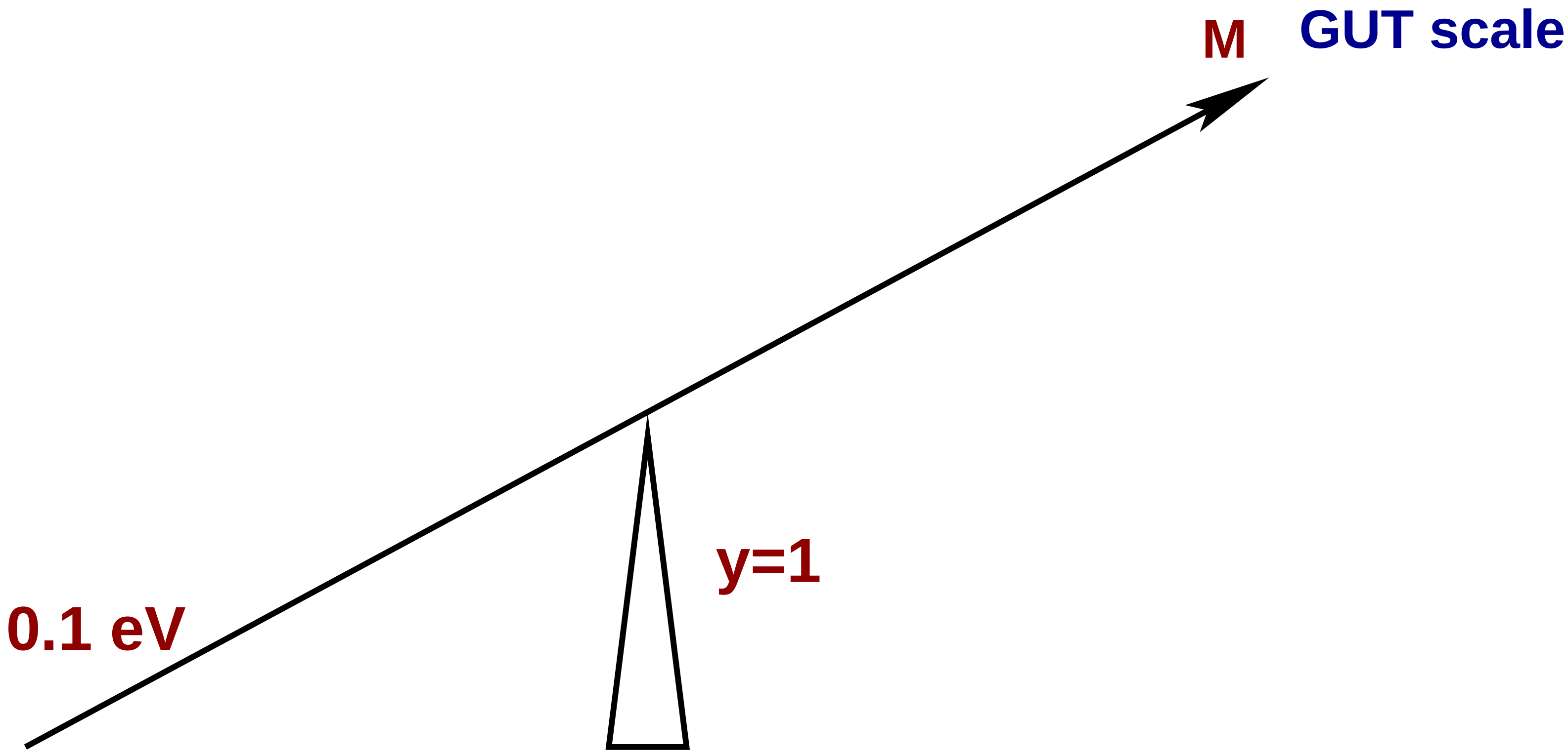
$$D_{ij} = y_{ij} \langle H \rangle$$

Smallness of neutrino masses **does not** imply the smallness of Yukawa couplings. For large M,

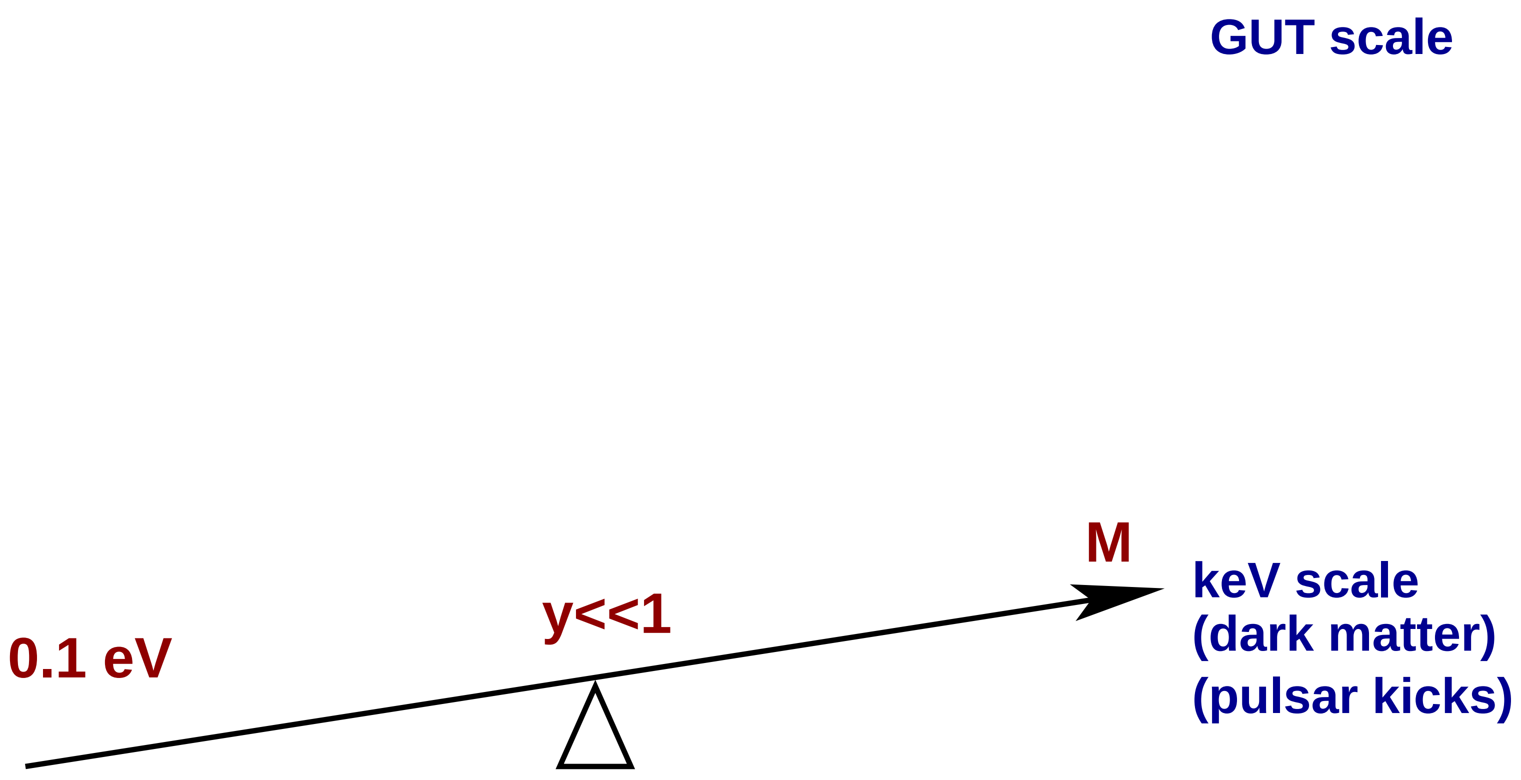
$$m_\nu \sim \frac{y^2 \langle H \rangle^2}{M}$$

One can understand the smallness of neutrino masses even if the Yukawa couplings are $y \sim 1$ [Yanagida; Gell-Mann, Ramond, Slansky; Glashow; Mohapatra, Senjanović].

Seesaw mechanism



Seesaw mechanism



Is $y \sim 1$ better than $y \ll 1$?

Depends on the model.

- If $y \approx$ some intersection number in string theory, then $y \sim 1$ is natural
- If y comes from wave function overlap of fermions living on different branes in a model with extra-dimensions, then it can be exponentially suppressed, hence, $y \ll 1$ is natural.

In the absence of theory of the Yukawa couplings, one evokes some naturalness arguments.

't Hooft's naturalness criterion

Small number is natural if setting it to zero increases the symmetry

Small breaking of the symmetry $\Rightarrow$ small number

- Pion masses are small because the massless pions correspond to exact chiral symmetry **natural**
- Gauge hierarchy problem: small $M_{\rm Higgs}/m_{\rm Planck}$ is **not natural in the Standard Model** because setting $M_{\rm Higgs} = 0$ does not increase the symmetry. In a supersymmetric extension, $M_{\rm Higgs} \approx M_{\rm Higgsino}$, and setting $M_{\rm Higgsino} = 0$ increases the overall (chiral) symmetry. Hence, a light Higgs is **natural in SUSY models**.
- Cosmological constant problem: $\Lambda \to 0$ does not increase the symmetry. Hence, **not natural**.

What if one apples this criterion to sterile neutrinos? Symmetry increases for $M \to 0$, namely, the chiral symmetry of right-handed fields.

Small M is technically **natural**.

Clues from cosmology?

Baryon asymmetry of the universe could be generated by **leptogenesis** [Fukugita,Yanagida]

However, leptogenesis can work for both $M \gg 100$ GeV and $M < 100$ GeV:

- For $M \gg 100$ GeV, heavy sterile neutrino decays can produce the lepton asymmetry, which is converted to baryon asymmetry by sphalerons [Fukugita,Yanagida]
- For $M < 100$ GeV, neutrino oscillations can produce the lepton asymmetry, which is converted to baryon asymmetry by sphalerons [Akhmedov, Rubakov, Smirnov; Asaka, Shaposhnikov]

Over the years, neutrino physics has shown many theoretical prejudices to be wrong: neutrinos were expected to be massless, neutrinos were expected to have small mixing angles, etc.

Since the fundamental theory of neutrino masses is lacking, one should

consider all allowed values
for the sterile neutrino masses

in the following lagrangian:

$$\mathcal{L} = \mathcal{L}_{\mathrm{SM}} + \bar{\nu}_{s,a} \left(i\partial_\mu \gamma^\mu\right) \nu_{s,a} - y_{\alpha a} H \, \bar{L}_\alpha \nu_{s,a} - \frac{M_{aa}}{2} \, \bar{\nu}^c_{s,a} \nu_{s,a} + h.c. \, ,$$

where M is can be small or large

Astrophysical clues: dark matter

Dark matter – a simple solution:

侘寂

use one of the particles already introduced to give the neutrino masses

$\Rightarrow$ **sterile neutrino**

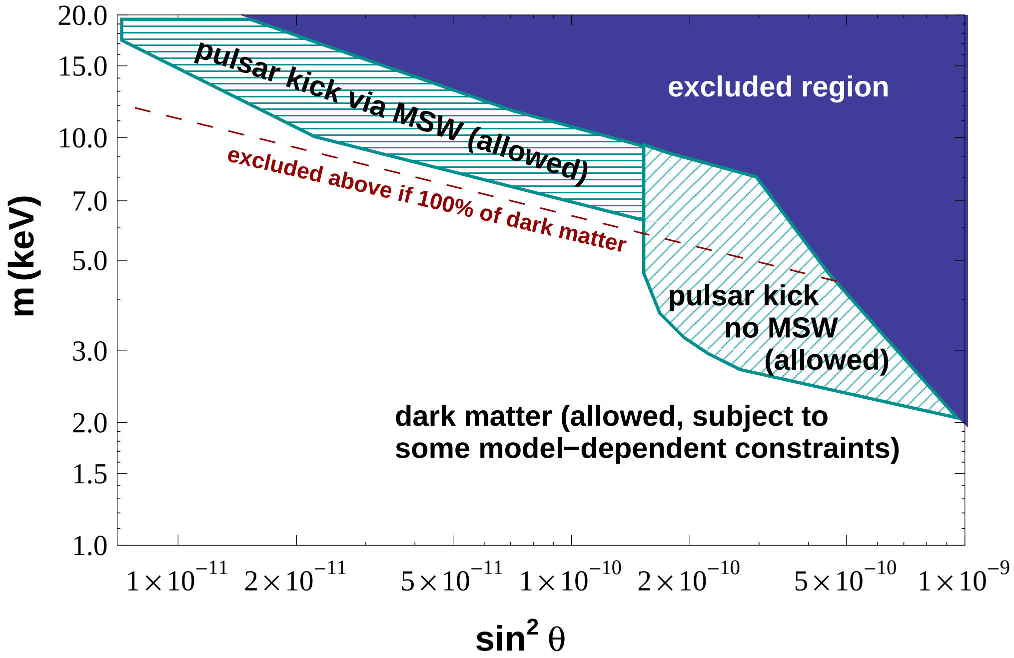
side benefit: **explanation of the pulsar kicks, supernova asymmetries**

Sterile neutrinos in the early universe

Sterile neutrinos are produced in primordial plasma through

- off-resonance oscillations. [Dodelson, Widrow; Abazajian, Fuller; Dolgov, Hansen; Asaka, Laine, Shaposhnikov et al.]
- oscillations on resonance, if the lepton asymmetry is non-negligible [Fuller, Shi]
- production mechanisms which do not involve oscillations
  - inflaton decays directly into sterile neutrinos [Shaposhnikov, Tkachev]
  - Higgs physics: both mass and production [AK, Petraki]

X-ray and Lyman-α bounds on sterile neutrinos



[AK, PRL **97**, 241301 (2006)]

[A. Palazzo, D. Cumberbatch, A. Slosar, J. Silk Phys.Rev. **D76**, 103511 (2007)]

Dark matter and the Lyman-α forest.

The bounds depend on the production mechanism.

$$\lambda_{FS} \approx 1\,\mathrm{Mpc} \left(\frac{\mathrm{keV}}{m_s}\right) \left(\frac{\langle p_s \rangle}{3.15\,T}\right)_{T\approx 1\,\mathrm{keV}}$$

The ratio

$$\left(\frac{\langle p_s \rangle}{3.15\,T}\right)_{T\approx 1\,\mathrm{keV}} = \begin{cases} 0.9 & \text{for production off} - \text{resonance} \\ 0.6 & \text{for MSW resonance (depends on L)} \\ 0.2 & \text{for production at T} > 100\ \text{GeV} \end{cases}$$

Neutrino masses: new scale or new Higgs physics?

$$\mathcal{L} = \mathcal{L}_{\rm SM} + \bar{N}_a \left(i \partial_\mu \gamma^\mu \right) N_a - y_{\alpha a} H \, \bar{L}_\alpha N_a - \frac{M_a}{2} \, \bar{N}_a^c N_a + h.c. \, ,$$

To explain the pulsar kicks and dark matter, one needs $M \sim$ **keV**. Is this a new fundamental scale? Perhaps. Alternatively, it could arise from the Higgs mechanism:

$$\mathcal{L} = \mathcal{L}_{\rm SM} + \bar{N}_a \left(i \partial_\mu \gamma^\mu \right) N_a - y_{\alpha a} H \, \bar{L}_\alpha N_a - h_a \, S \, \bar{N}_a^c N_a + V(H, S)$$

$$M = h \langle S \rangle$$

Now $S \rightarrow NN$ decays can produce sterile neutrinos.

For small h, the sterile neutrinos are out of equilibrium in the early universe, but S is in equilibrium. There is a new mechanism to produce sterile dark matter at $T \sim m_S$ from decays $S \longrightarrow NN$:

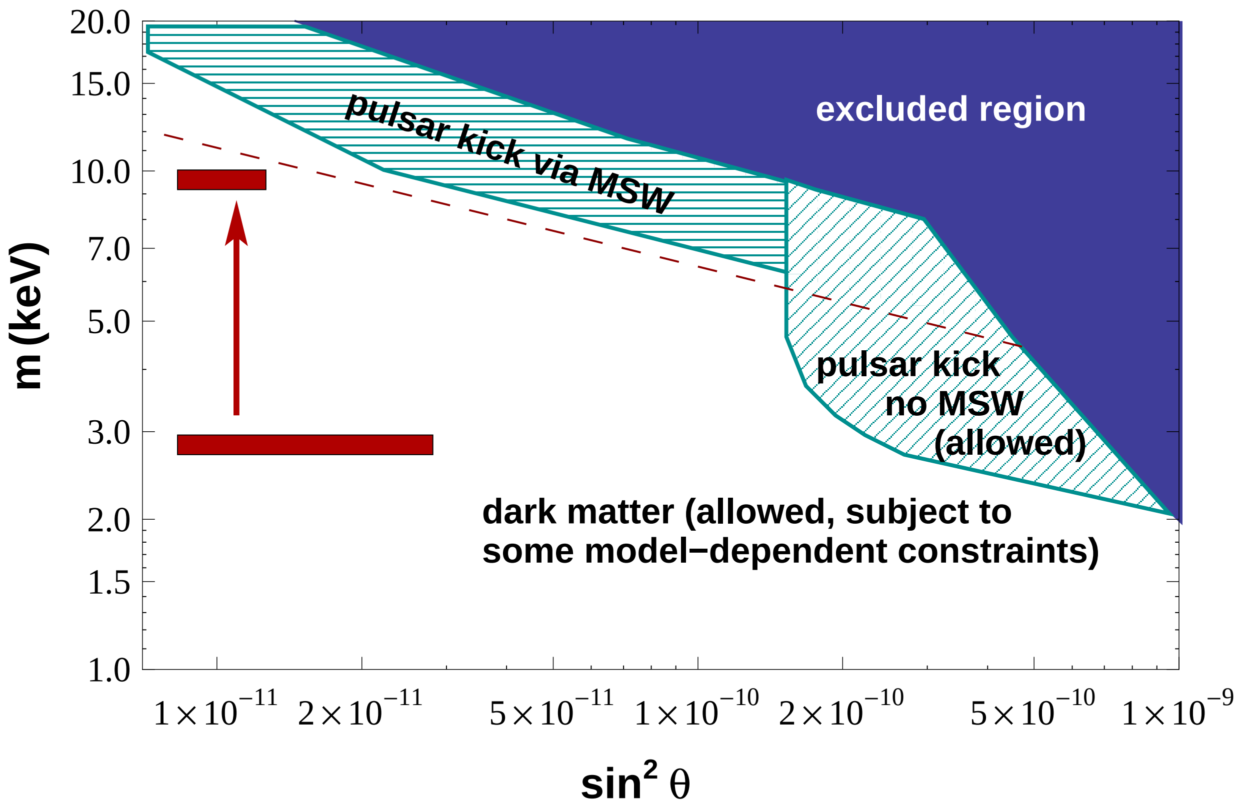
$$\Omega_s = 0.2 \left(\frac{33}{\xi}\right) \left(\frac{h}{1.4 \times 10^{-8}}\right)^3 \left(\frac{\langle S \rangle}{\tilde{m}_S}\right)$$

Here ζ is the dilution factor due to the change in effective numbers of degrees of freedom.

$$\langle S \rangle = \frac{M_s}{h} \sim \frac{\text{few keV}}{1.4 \times 10^{-8}} \sim 10^2\,\text{GeV}$$

The sterile neutrino momenta are red-shifted by factor $\zeta^{1/3} > 3.2$. [AK, Petraki]

Cooling changes the Lyman-α bounds

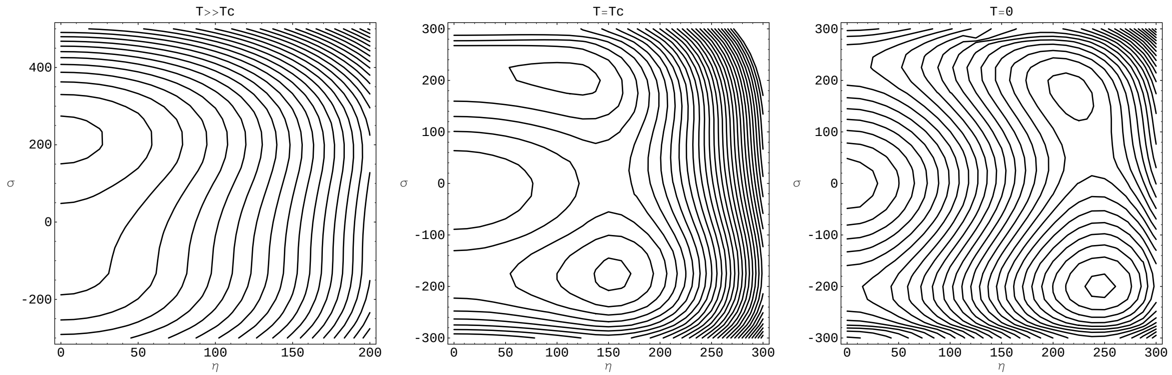


[AK, PRL **97**:241301 (2006); Petraki, AK, PRD 77, 065014 (2008); Petraki, PRD 77, 105004 (2008)]

Implications for the EW phase transition and the LHC

One may be able to discover the *singlet Higgs* at the LHC [Profumo, Ramsey-Musolf, G. Shaughnessy; Davoudiasl et al.; O'Connell et al.; Ramsey-Musolf, Wise]

The presence of S in the Higgs sector changes the nature of the electroweak phase transition [AK, Petraki]



First-order transition, CP in the Higgs sector $\Longrightarrow$ **electroweak baryogenesis**

Astrophysical clues: supernova

- Sterile neutrino emission from a supernova is anisotropic due to
  1. asymmetries in the urca cross sections
  2. magnetic effects on neutrino oscillations
- Sterile neutrinos with masses and mixing angles consistent with dark matter can explain the pulsar velocities

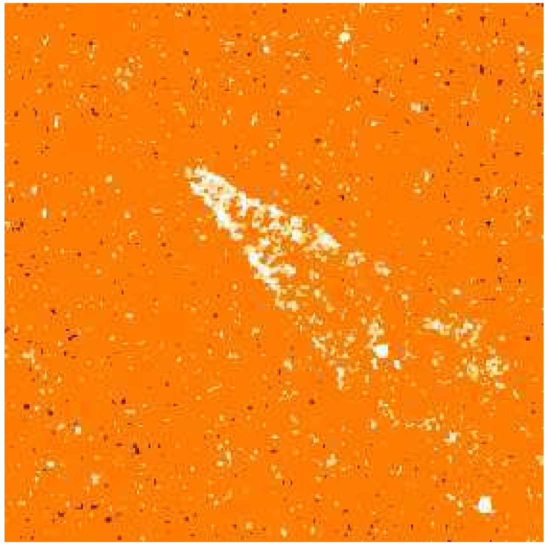
[AK, Segrè; Fuller, AK, Mocioiu, Pascoli; Barkovich, D'Olivo, Montemayor]

The pulsar velocities.

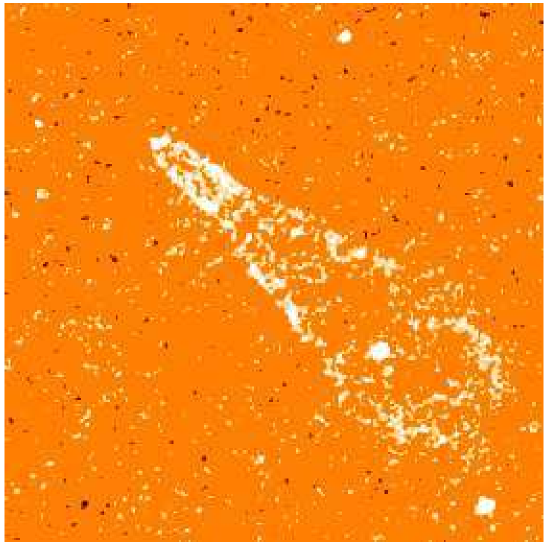
Pulsars have large velocities, $\langle v \rangle \approx 250 - 450$ km/s.
[Cordes *et al.*; Hansen, Phinney; Kulkarni *et al.*; Lyne *et al.*]

A significant population with $v > 700$ km/s,
about **15 %** have $v > 1000$ km/s, up to 1600 km/s.
[Arzoumanian *et al.*; Thorsett *et al.*]

A very fast pulsar in Guitar Nebula

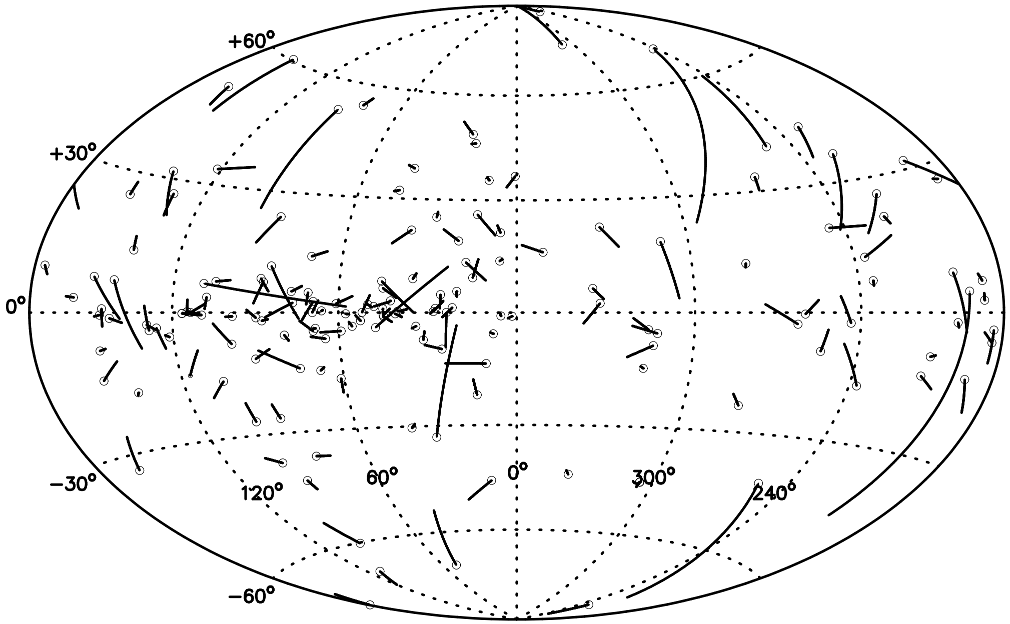


HST, December 1994



HST, December 2001

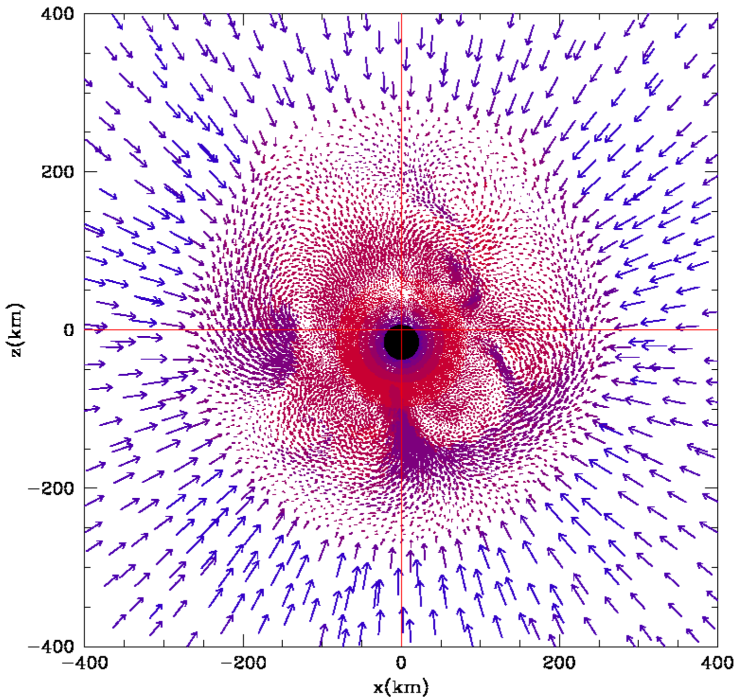
Map of pulsar velocities

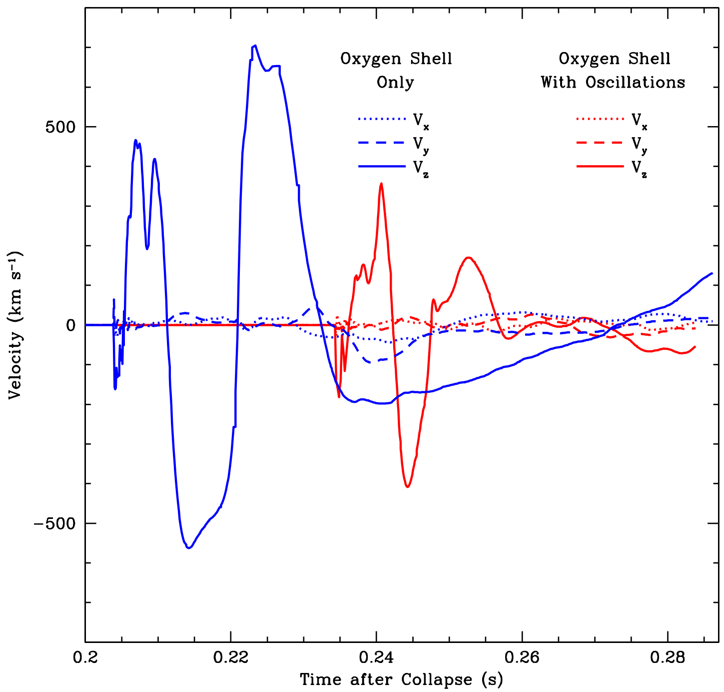


Proposed explanations:

- asymmetric collapse [Shklovskii] (small kick)
- evolution of close binaries [Gott, Gunn, Ostriker] (not enough)
- acceleration by EM radiation [Harrison, Tademaru] (kick small, predicted polarization not observed)
- asymmetry in EW processes that produce neutrinos [Chugai; Dorofeev, Rodinov, Ternov] (asymmetry washed out)
- "cumulative" parity violation [Lai, Qian; Janka] (it's *not* cumulative)
- various exotic explanations
- explanations that were "not even wrong"...

Asymmetric collapse





"...the most extreme asymmetric collapses
do not produce final neutron star velocities above 200km/s" [Fryer]

Supernova neutrinos

Nuclear reactions in stars lead to a formation of a heavy iron core. When it reaches $M \approx 1.4 M_{\odot}$, the pressure can no longer support gravity. $\Rightarrow$ collapse.

Energy released:

$$\Delta E \sim \frac{G_N M_{\rm Fe\ core}^2}{R} \sim 10^{53} {\rm erg}$$

99% of this energy is emitted in neutrinos

Pulsar kicks from neutrino emission?

Pulsar with $v \sim 500$ km/s has momentum

$$M_{\odot} v \sim 10^{41} \text{ g cm/s}$$

SN energy released: 10^{53} erg $\Rightarrow$ in neutrinos. Thus, the total neutrino momentum is

$$P_{\nu;\,\text{total}} \sim 10^{43} \text{ g cm/s}$$

a **1% asymmetry** in the distribution of **neutrinos**

is sufficient to explain the pulsar kick velocities

But what can cause the asymmetry??

Magnetic field?

Neutron stars have large magnetic fields. A typical pulsar has surface magnetic field $B \sim 10^{12} - 10^{13}$ **G**.

Recent discovery of *soft gamma repeaters* and their identification as *magnetars*

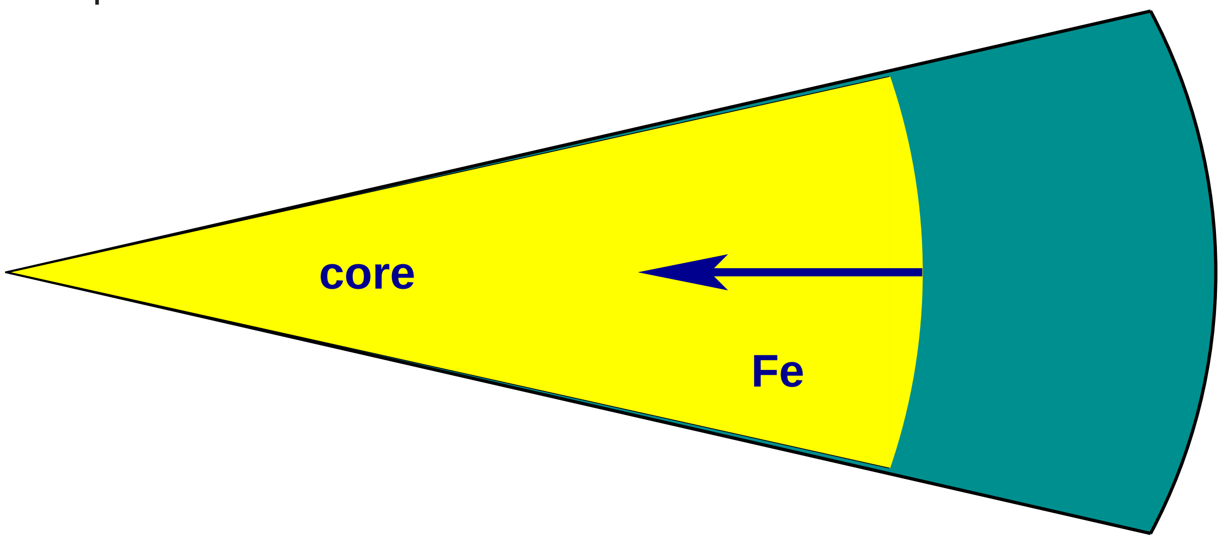
$\Rightarrow$ some neutron stars have surface magnetic fields as high as $10^{15} - 10^{16}$ **G**.

$\Rightarrow$ **magnetic fields inside can be** $10^{15} - 10^{16}$ **G.**

Neutrino magnetic moments are negligible, but the **scattering of neutrinos off polarized electrons and nucleons** is affected by the magnetic field.

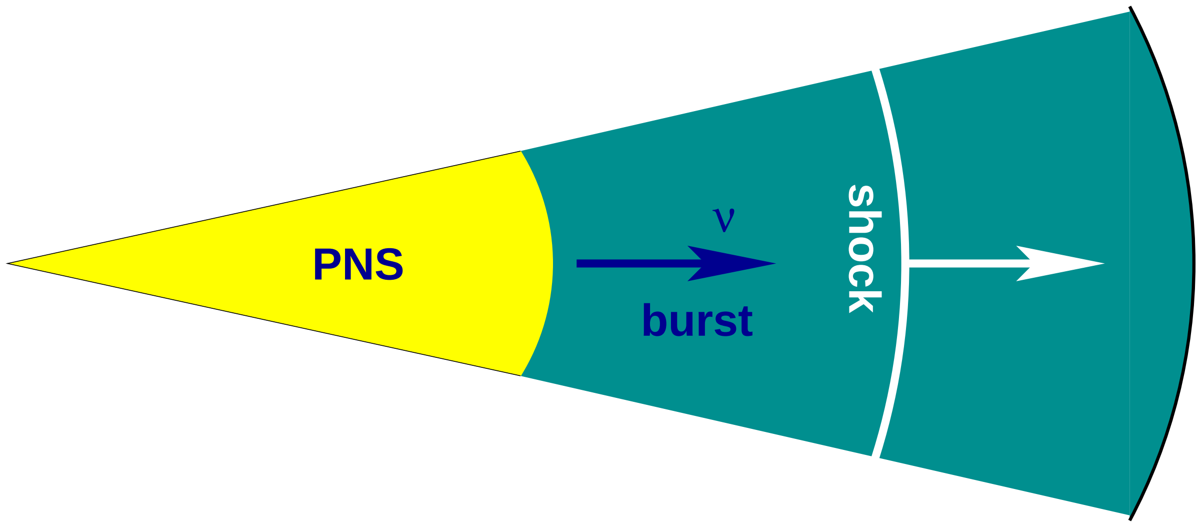
Core collapse supernova

Onset of the collapse: $t = 0$



Core collapse supernova

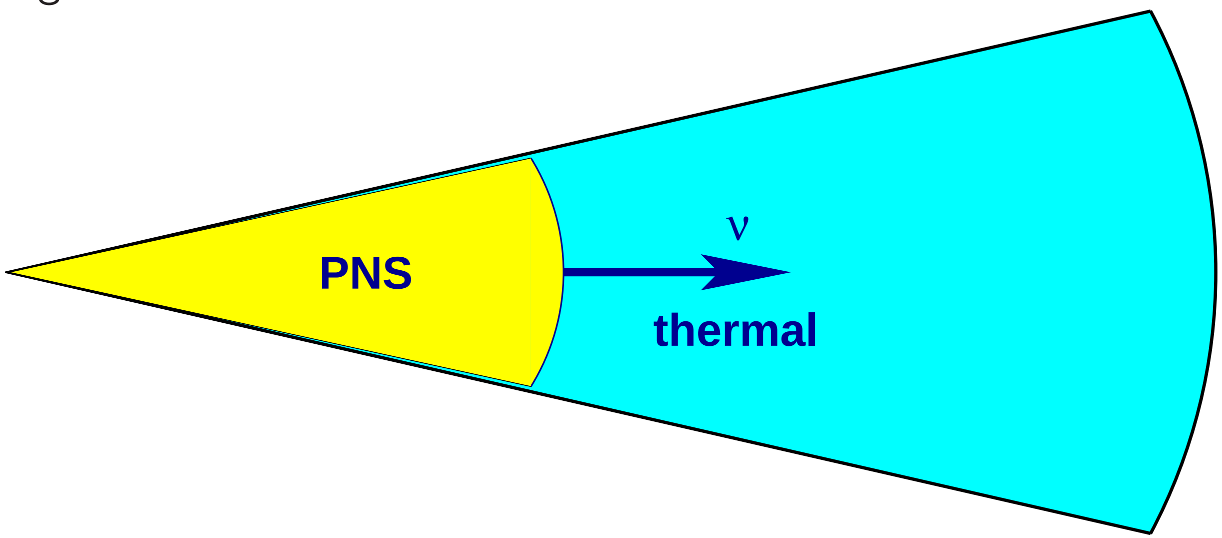
Shock formation and "neutronization burst": $t = 1 - 10$ **ms**



Protoneutron star formed. Neutrinos are trapped. The shock wave breaks up nuclei, and the initial neutrino come out (a few %).

Core collapse supernova

Thermal cooling: $t = 10 - 15$ s



Most of the neutrinos emitted during the cooling stage.

Electroweak processes producing neutrinos (urca),

$$p + e^- \rightleftharpoons n + \nu_e \text{ and } n + e^+ \rightleftharpoons p + \bar{\nu}_e$$

have an asymmetry in the production cross section, depending on the spin orientation.

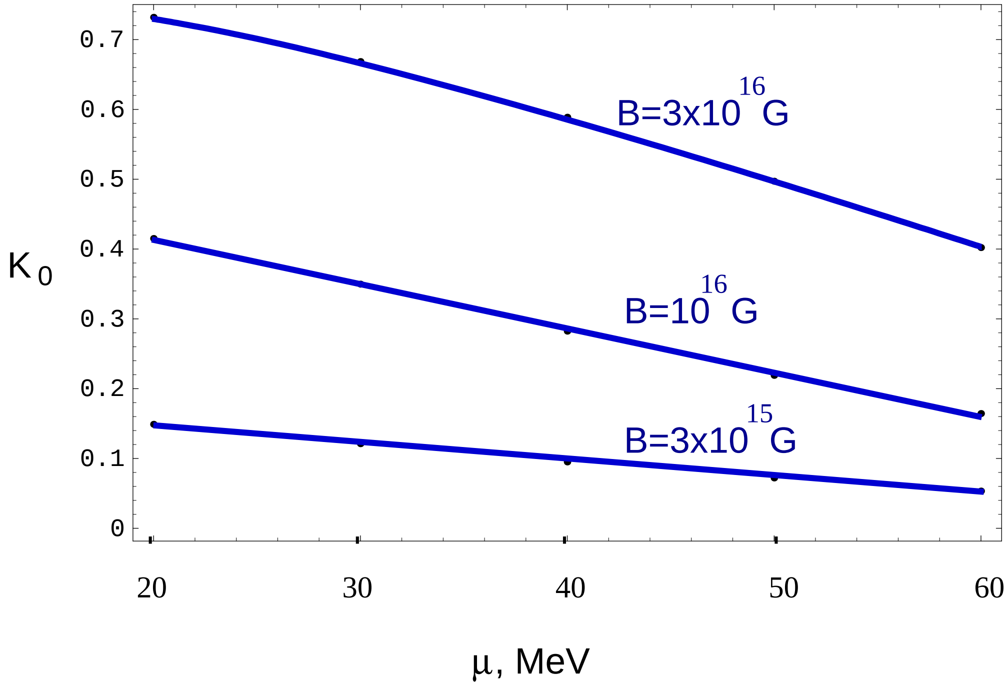
$$\sigma(\uparrow \quad e^-, \uparrow \quad \nu) \neq \sigma(\uparrow \quad e^-, \downarrow \nu)$$

The asymmetry:

$$\tilde{\epsilon} = \frac{g_V^2 - g_A^2}{g_V^2 + 3g_A^2} k_0 \approx 0.4\, k_0,$$

where k_0 is the fraction of electrons in the lowest Landau level.

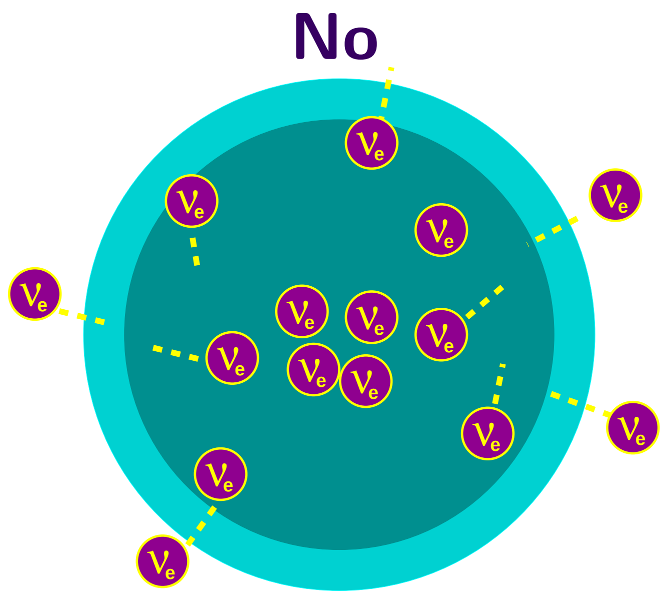
In a strong magnetic field,



k_0 is the fraction of electrons in the lowest Landau level.

Pulsar kicks from the asymmetric production of neutrinos?
[Chugai; Dorofeev, Rodionov, Ternov]

Can the weak interactions asymmetry cause an anisotropy in the flux of neutrinos due to a large magnetic field?



Neutrinos are trapped at high density.

Can the weak interactions asymmetry cause an anisotropy in the flux of neutrinos due to a large magnetic field?

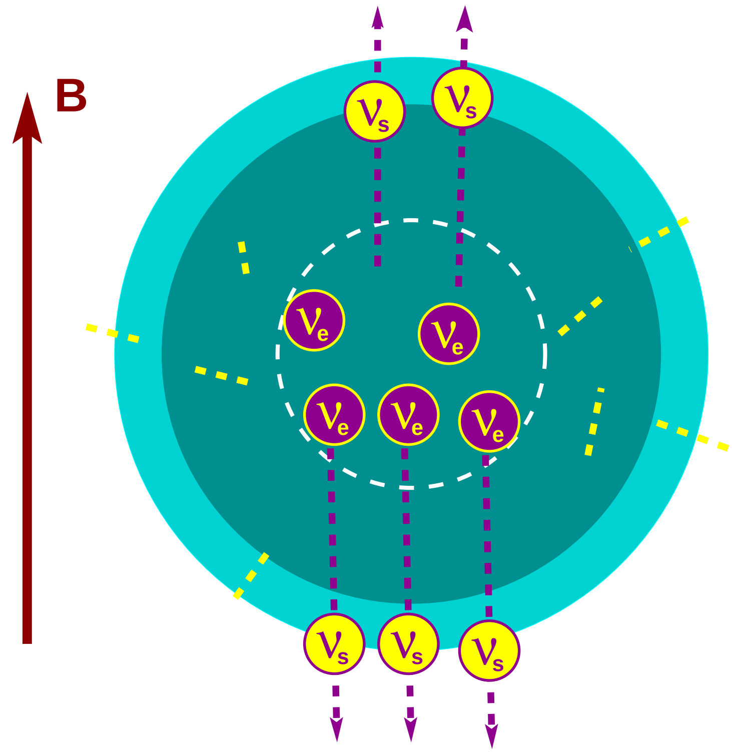
No

Rescattering washes out the asymmetry

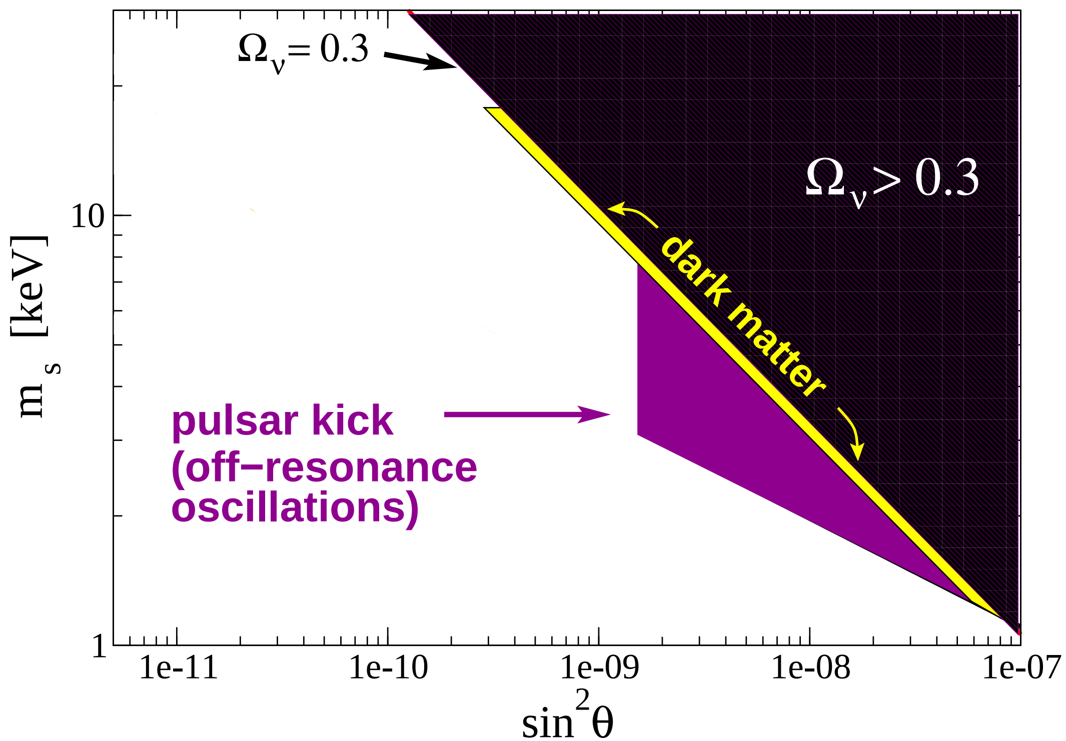
In approximate thermal equilibrium the asymmetries in scattering amplitudes do not lead to an anisotropic emission [Vilenkin,AK, Segrè]. Only the outer regions, near neutrinospheres, contribute, but the kick would require a mass difference of $\sim 10^2$ eV [AK,Segrè].

However, if a weaker-interacting sterile neutrino was produced in these processes, the asymmetry would, indeed, result in a pulsar kick!

[AK, Segrè; Fuller, AK, Mocioiu, Pascoli]


B
νs
νs
νe
νe
νe
νe
νe
νs
νs
νs

Allowed range of parameters (time scales, fraction of total energy emitted):



[Fuller, AK, Mocioiu, Pascoli]

Resonance in the magnetic field

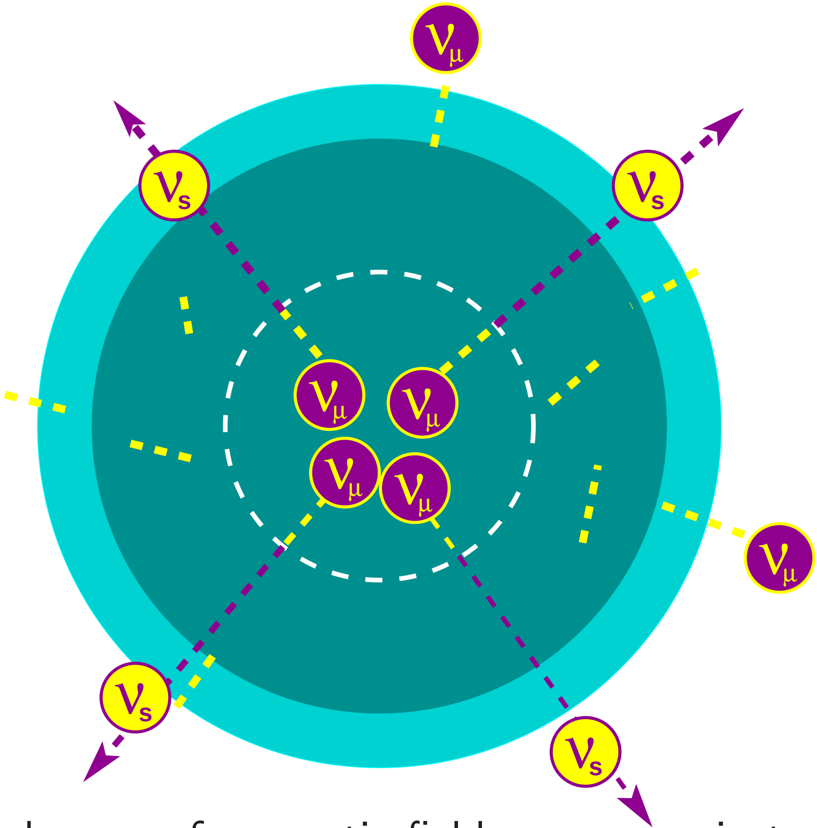
Matter potential:

$$\begin{aligned} V(\nu_s) &= 0 \\ V(\nu_e) &= -V(\bar{\nu}_e) = V_0\,(3\,Y_e - 1 + 4\,Y_{\nu e}) \\ V(\nu_{\mu,\tau}) &= -V(\bar{\nu}_{\mu,\tau}) = V_0\,(Y_e - 1 + 2\,Y_{\nu e}) \;+\; c_L^Z\,\frac{\vec{k}\cdot\vec{B}}{k} \end{aligned}$$

$$c_L^Z = \frac{eG_F}{\sqrt{2}}\left(\frac{3N_e}{\pi^4}\right)^{1/3}$$

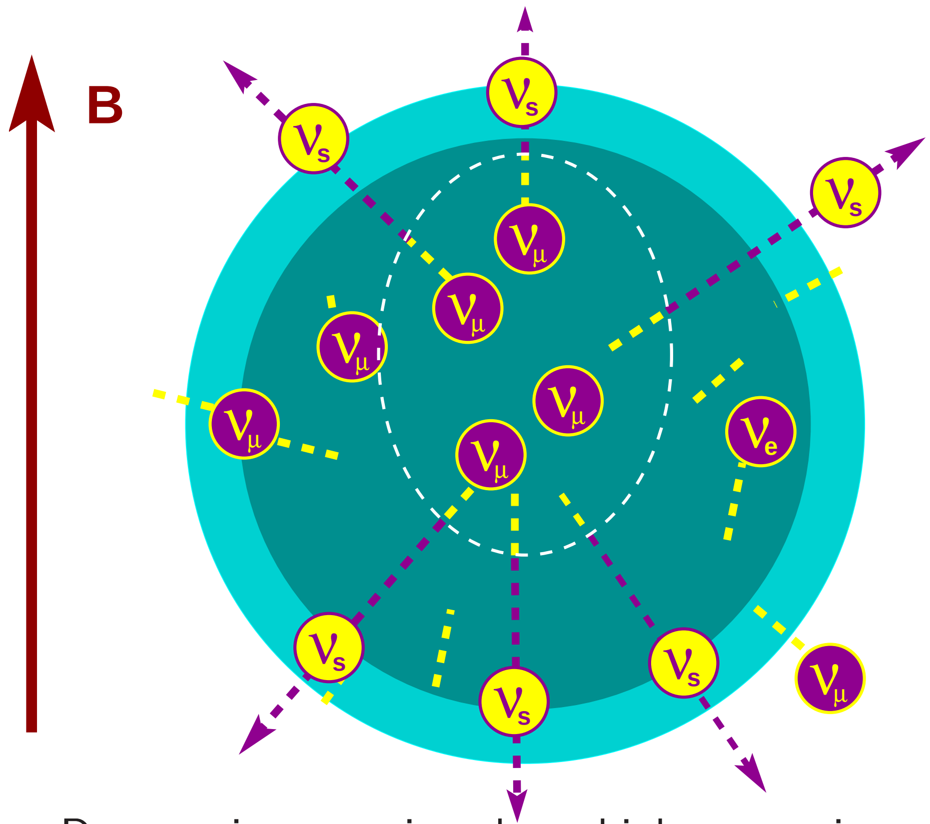
[D'Olivo, Nieves, Pal; Semikoz]

The magnetic field shifts the position of the resonance because of the $\frac{\vec{k}\cdot\vec{B}}{k}$ term in the potential:



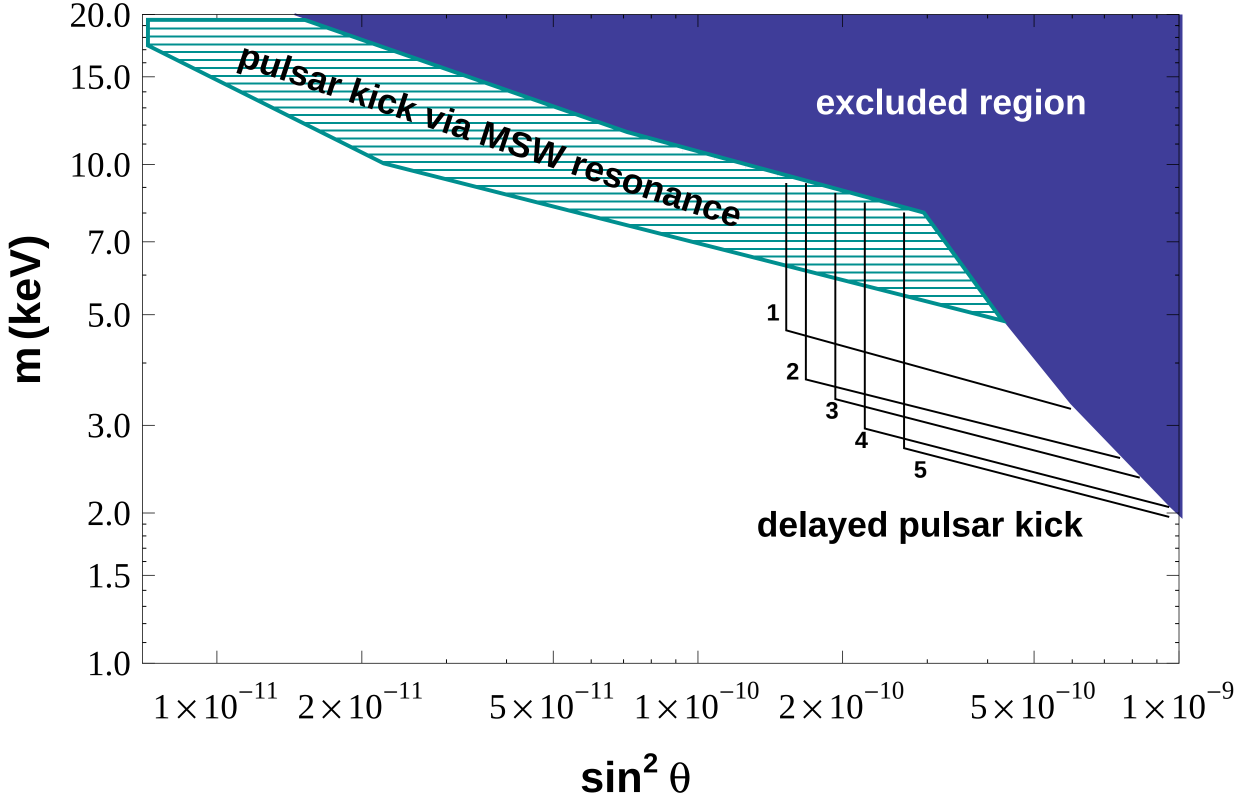
In the absence of magnetic field, ν_s escape isotropically

The magnetic field shifts the position of the resonance because of the $\frac{\vec{k}\cdot\vec{B}}{k}$ term in the potential:



Down going neutrinos have higher energies

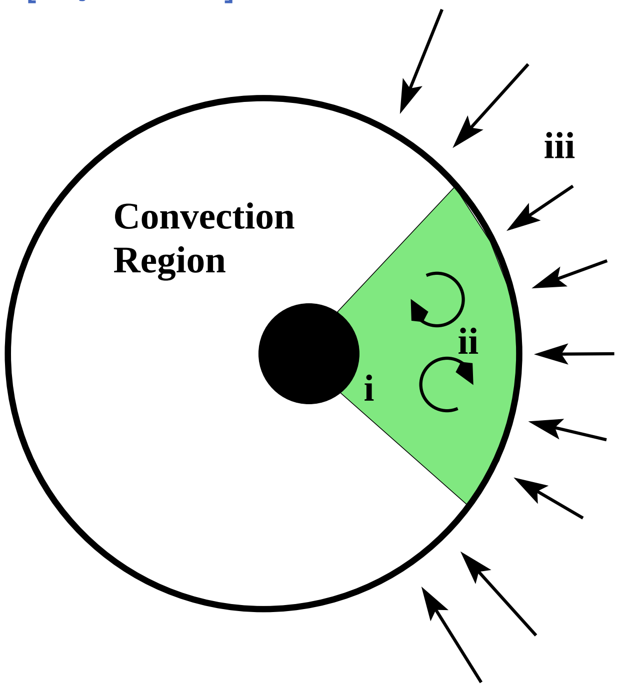
Allowed range of masses and mixing angles



[A.K., Segrè; Fuller, A.K.,Mocioiu,Pascoli; Barkovich, D'Ollivo, Montemayor; AK et al.]

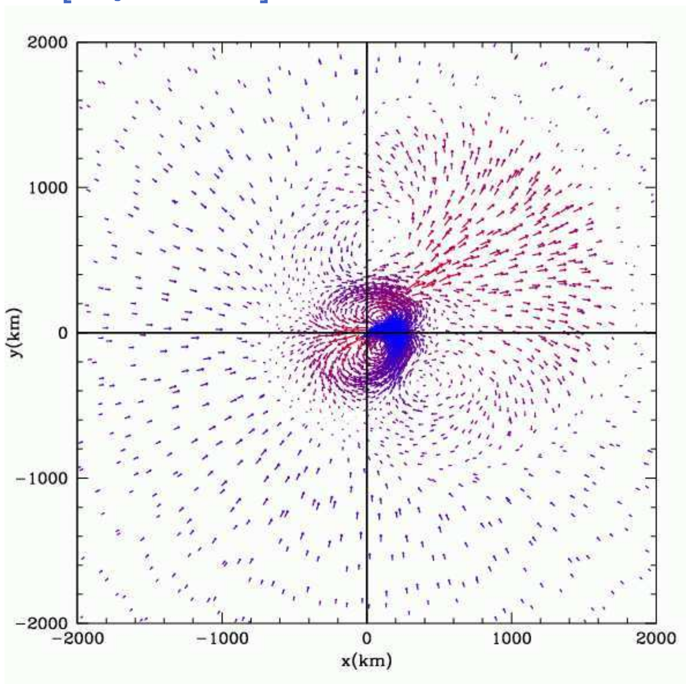
Other predictions of the pulsar kick mechanism

- Stronger supernova shock [Fryer, AK]



Other predictions of the pulsar kick mechanism

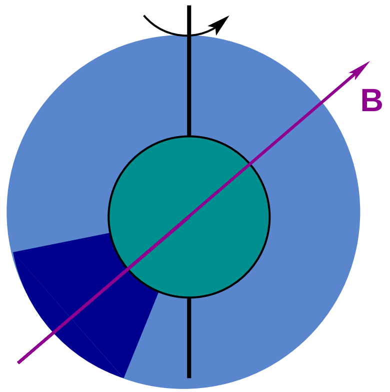
- Stronger supernova shock [Fryer, AK]



Other predictions of the pulsar kick mechanism

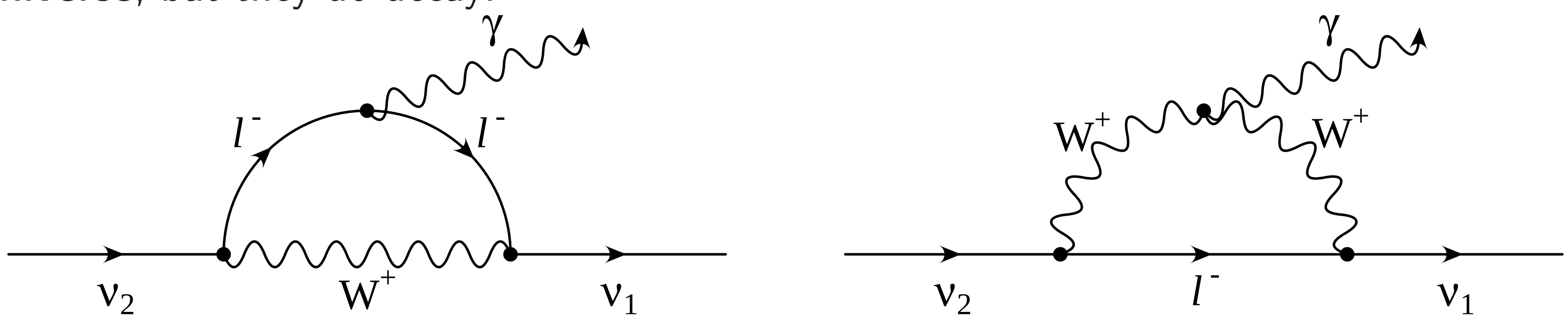
- Stronger supernova shock [Fryer, AK; Fuller, Hidaka]
- **No $B - v$ correlation** expected because
  - the magnetic field *inside* a hot neutron star during the *first ten seconds* is very different from the surface magnetic field of a cold pulsar
  - rotation washes out the x, y components
- **Directional $\vec{\Omega} - \vec{v}$ correlation** was predicted, because
  - the direction of rotation remains unchanged
  - only the z-component survives

This correlation has been confirmed by recent data.



Radiative decay

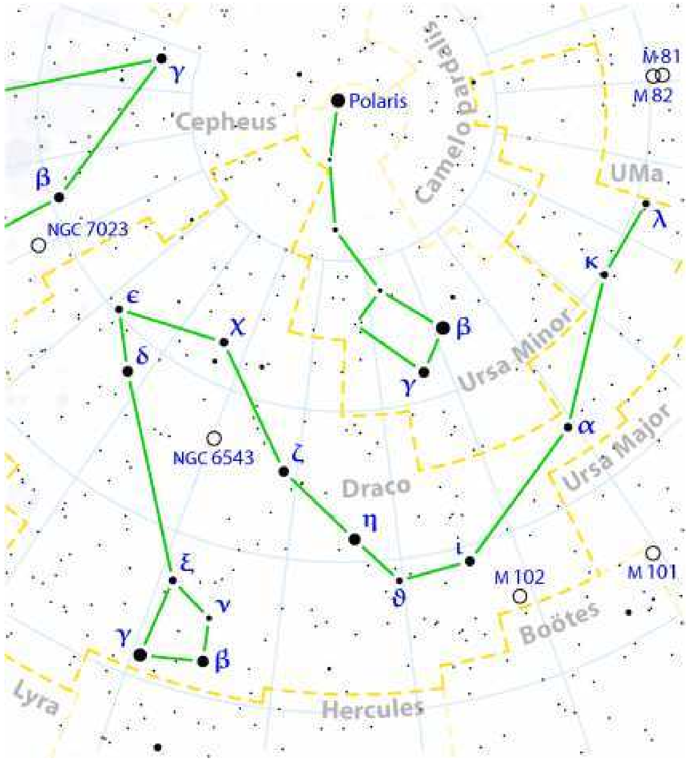
Sterile neutrino in the mass range of interest have lifetimes **longer than the age of the universe**, but they do decay:



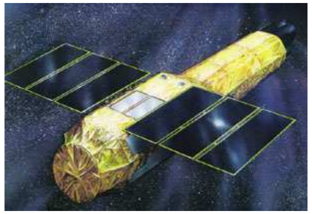
Photons have energies $m/2$: X-rays. Concentrations of dark matter emit X-rays. [Abazajian, Fuller, Tucker; Dolgov, Hansen; Shaposhnikov et al.]

[Loewenstein]

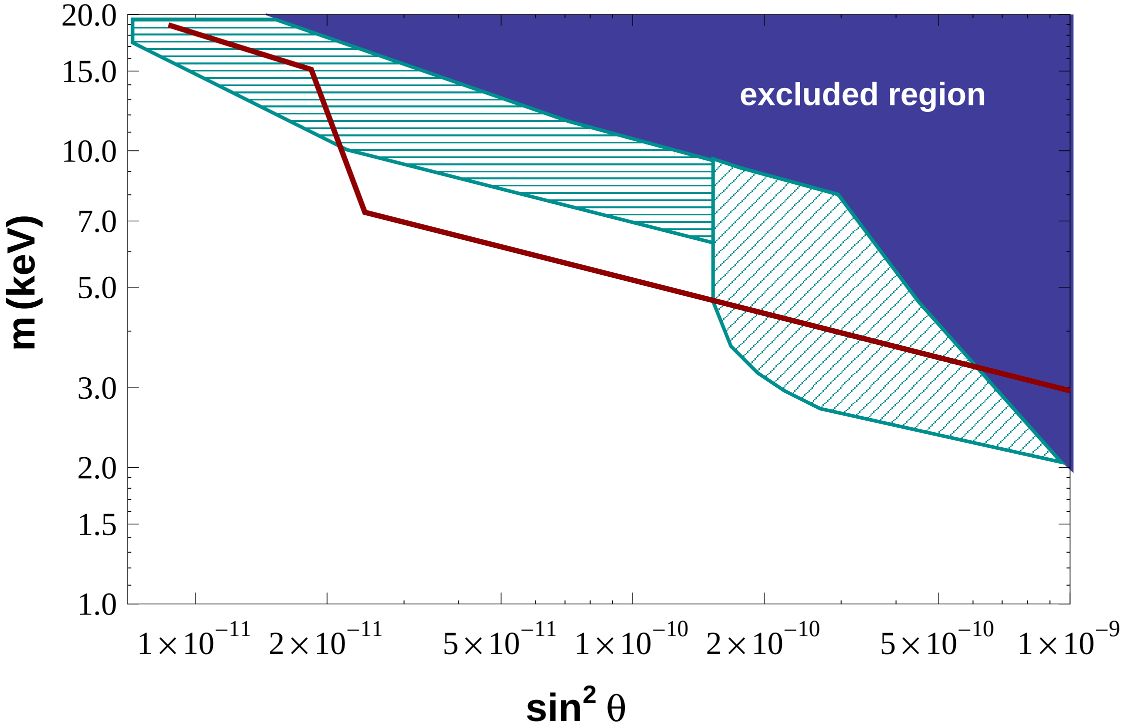
Suzaku observations of dSphs Draco and Ursa Minor



[Biermann, AK, Loewenstein, in preparation]



X-ray observations (very preliminary)



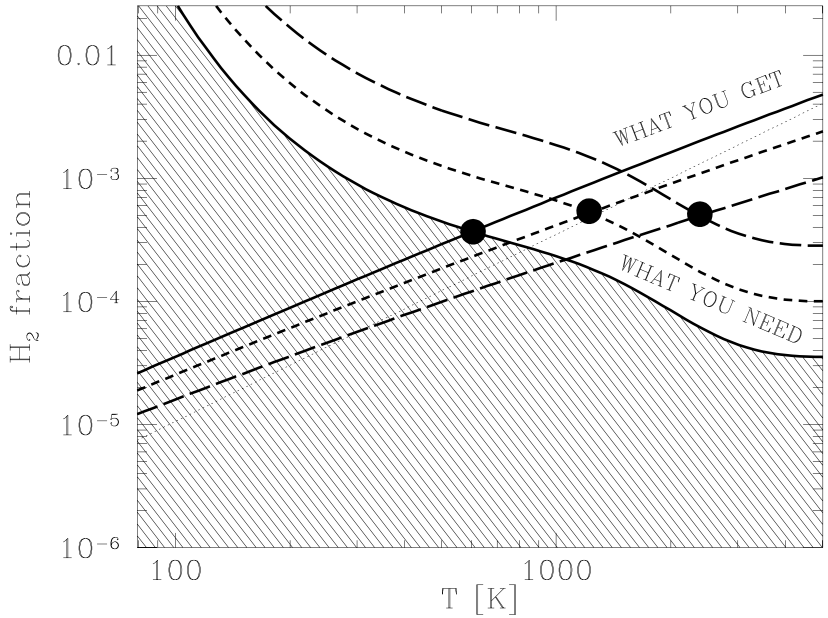
[Loewenstein, Biermann, AK]

Some challenges to CDM

- overproduction of the satellite halos for galaxies of the size of Milky Way [Bullock]
- WDM can reduce the number of halos in low-density voids. [Peebles]
- observed densities of the galactic cores (from the rotation curves) are lower than what is predicted based on the ΛCDM power spectrum. [Dalcanton et al.; van den Bosch et al.; Moore; Abazajian]
- The “angular-momentum problem”: in CDM halos, gas should cool at very early times into small halos and lead to massive low-angular-momentum gas cores in galaxies. [Dolgov]
- disk-dominated (pure-disk) galaxies are observed, but not produced in CDM because of high merger rate. [Governato et al.; Kormendy et al.]
- observations of dwarf spheroidal galaxies favor WDM [Gilmore et al.; Strigari et al.]
- Correlations in galaxy parameters (gallaxies ”too simple“) [Disney et al.]

Star formation and reionization

Molecular hydrogen is necessary for star formation



[Tegmark, et al., ApJ **474**, 1 (1997)]

Molecular hydrogen

$$H + H \longrightarrow H_2 + \gamma \quad - \quad \text{very slow!}$$

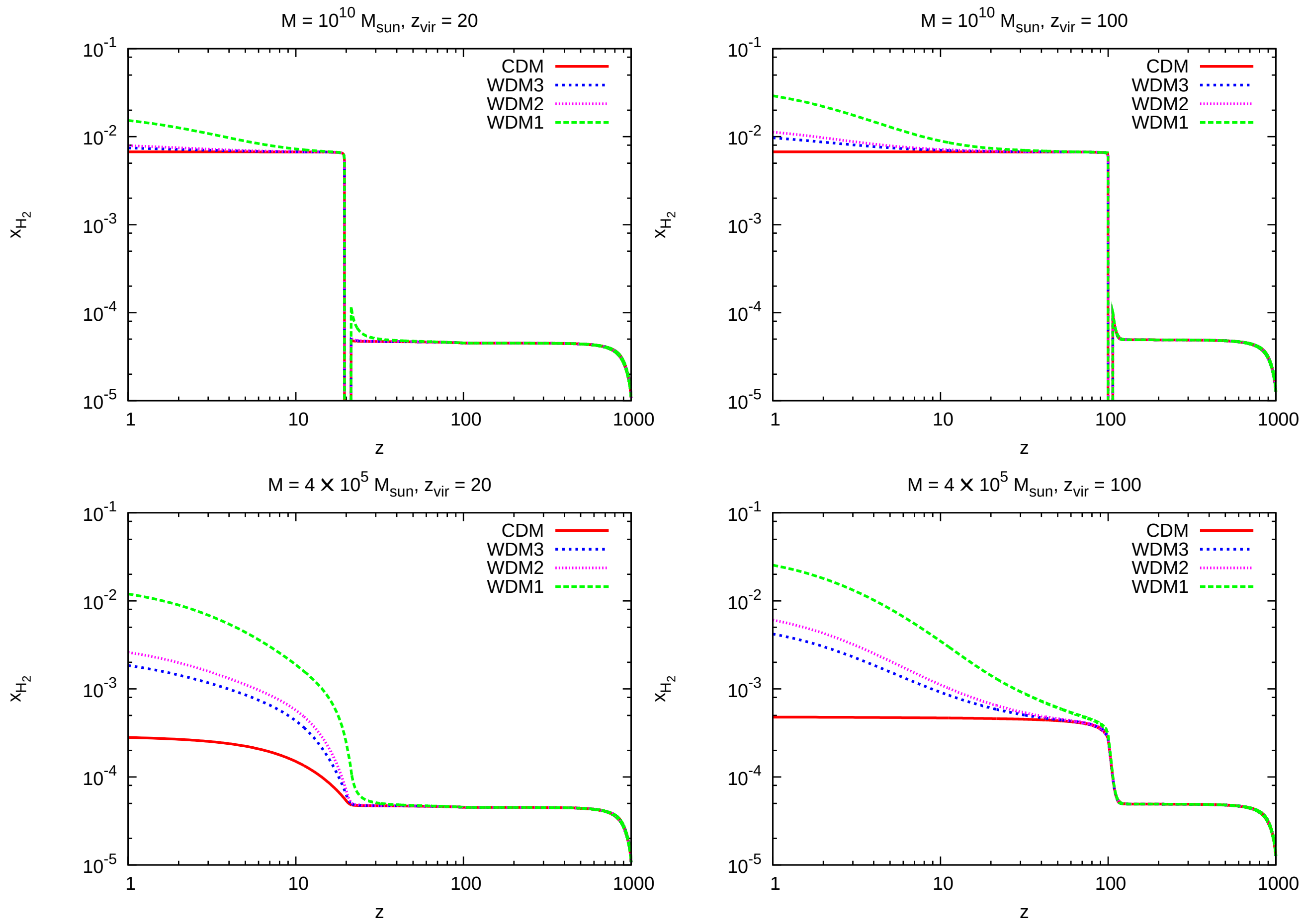
In the presence of ions the following reactions are faster:

$$\begin{aligned} H^+ + H &\longrightarrow H_2^+ + \gamma, \\ H_2^+ + H &\longrightarrow H_2 + H^+. \end{aligned}$$

H^+ produced by X-rays from $\nu_2 \longrightarrow \nu_1\gamma$ catalyze the formation of molecular hydrogen

[Biermann, AK, PRL **96**, 091301 (2006)]

[Stasielak, Biermann, AK, ApJ.654:290 (2007)]



[Biermann, AK; Stasielak, Biermann, AK]

0.2 GeV *eosphoric* sterile neutrinos

- allowed for $m \sim 0.2$ GeV, $\sin^2\theta \sim 10^{-8}$
- decay into pions, hence photons, on the time scale ~ 0.1 s
- can alter the energy transport dramatically and increase the explosion energy
- supernova is the place to look for them

[Fuller, AK, Petraki]

Summary

- SU(2) singlets with masses between eV and the Planck scale are needed to explain the observed neutrino masses ("seesaw" mechanism). The question is **not** whether they exist, but **what is the mass?** If one of these new degrees of freedom has a relatively small mass, it can be dark matter.
- A sterile neutrino with 2-50 keV mass can explain all the present data, including
  - dark matter (warm or cold, depending on the mass)
  - baryon asymmetry of the universe
  - pulsar velocities
- Supernova is a great laboratory
- X-ray telescopes, such as Suzaku, have the capabilities to discover this form of dark matter
- If discovered, the line from the relic sterile neutrinos can be used to map out the redshift distribution of dark matter. This can be used to study the structure and the expansion history of the universe.